

SPACE WEATHER INTRODUCTORY COURSE



Collaboration of



Solar-Terrestrial Centre of Excellence



Koninklijke luchtmacht



Koninklijk Nederlands
Meteorologisch Instituut
Ministerie van Infrastructuur en Milieu



SPACE WEATHER

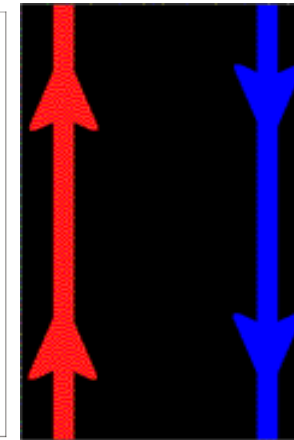
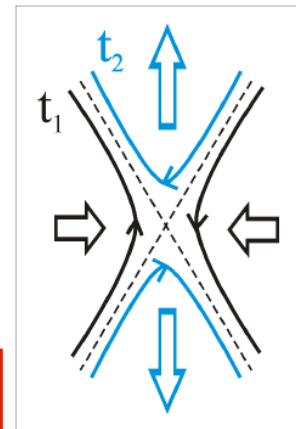
Drivers

Elke D'Huys, Jan Janssens, Elisabeth Dom & Petra Vanlommel



MAGNETIC RECONNECTION

- Conversion of magnetic energy into kinetic energy and radiation
- Where
 - Solar flares 
 - CME release from the Sun 
 - Geomagnetic storms at Earth 



Magnetic reconnection occurs when antiparallel magnetic field lines are pushed towards each other and interact. The field lines disconnect at the X-point and form new connections, thus restructuring the magnetic field and creating new topological configurations.

During the process of magnetic reconnection (stored) magnetic energy is converted into kinetic energy and heat (radiation).

The first observational evidence for the reconnection process was found in solar flares.

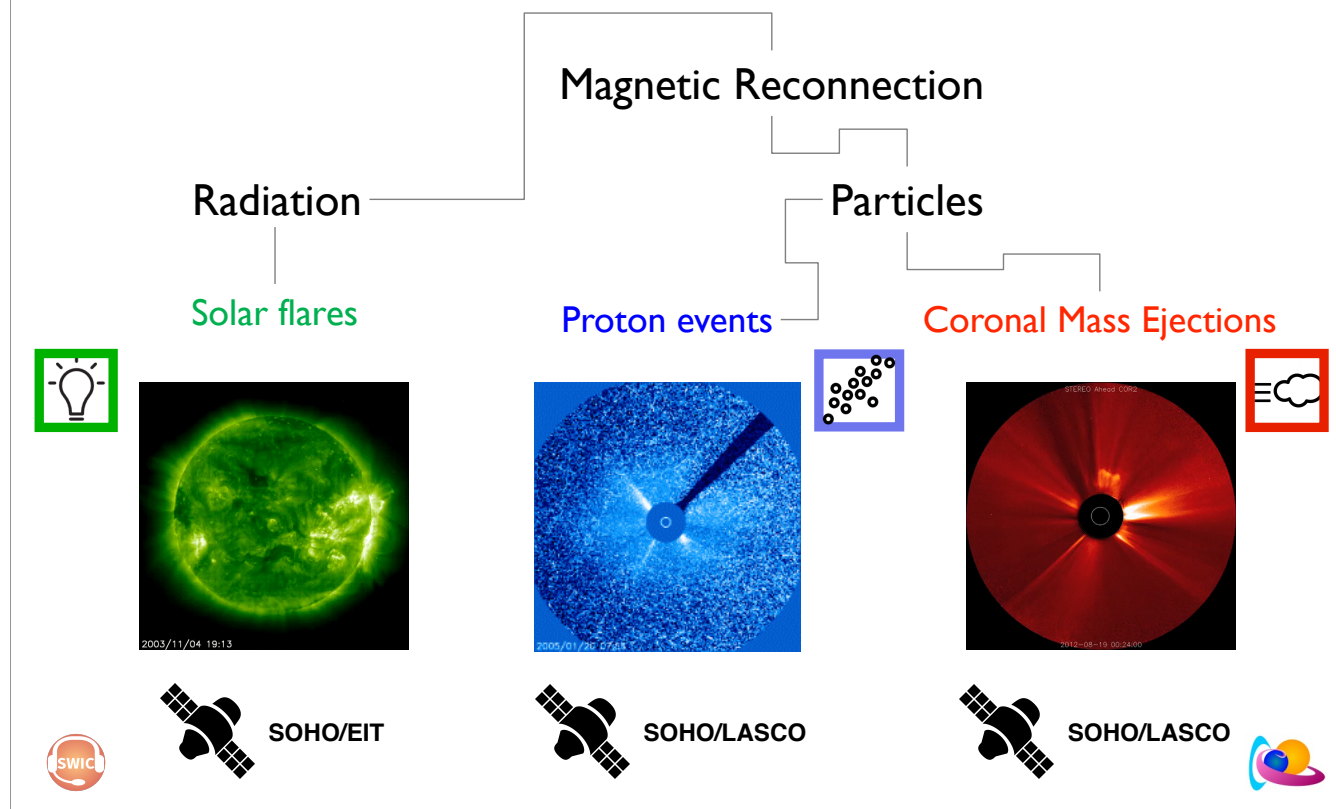
Sources:

Text from the CISM Summer School (Boulder, August 2013) – SW101_4_Flares

<https://www.bu.edu/cism/SummerSchool/summerlist.html>

Animation from ESA: <http://sci.esa.int/cluster/36447-direct-observation-of-3d-magnetic-reconnection/>

SOLAR ERUPTIONS



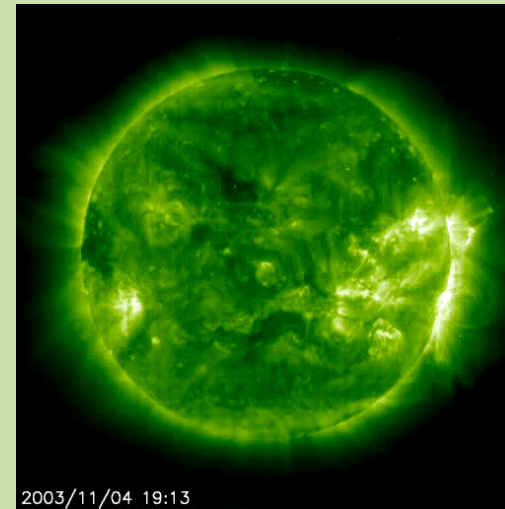
In this chapter we focus on eruptive solar events.

Magnetic reconnection is the physical principle at the base of the three types of space weather drivers: flares, proton events and coronal mass ejections.

FLARES - OVERVIEW



- Flare Characteristics
 - Definition
 - Standard model
 - Flare triggers
 - Flare features
- Flare Classification
- Flare Predictions



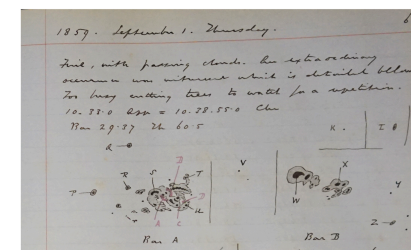
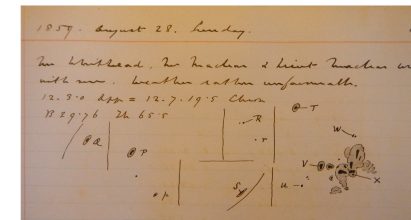
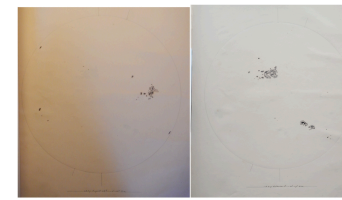
SOHO/EIT





SOLAR FLARES - WHAT?

- Sudden burst of radiation
- First observation in 1859 (white light) by Carrington and Hodgson
- Large quantity of energy is released
 $> 10^{20}$ Joule
- Required:
 A very rapid means of converting stored magnetic energy into particle energy and radiation – magnetic reconnection



Solar flares are sudden bursts of radiation lasting minutes to hours. Flares are observed in a wide range of electromagnetic waves such as radio, visible light, X-rays, and gamma rays.

Solar flares occur in a power-law spectrum of magnitudes; an energy release of typically 10^{20} joules of energy suffices to produce a clearly observable event, while a major event can emit up to 10^{25} joules. (For comparison: atomic bombs release energies of the order of 10^{12} J to 10^{15} J.)

(Other comparison: 10^{25} is 5 orders of magnitude more than annual world energy consumption)

A large quantity of energy is released from a small volume in a short period of time. This requires either a large amount of energy stored in that small volume that can be quickly transformed and released as energetic electrons and photons or very efficient transport of energy into that volume where it is then converted into the observed forms. The only viable energy source is intense solar magnetic fields.

Thus we need a very rapid means of converting stored magnetic energy into particle energy and heat: magnetic reconnection.

Magnetic energy is converted to thermal/radiative energy (flare, radio bursts) and kinetic energy (mass movement from CMEs and Solar Energetic Particles).

Flares in the visible continuum are particularly called white-light flares (WLFs), first observed by Carrington (1859). **Solar WLFs are usually rare** events compared to the H α and soft X-ray (SXR) flares because of the **short durations** (typically a few minutes; Hudson et al. 1992; Xu et al. 2006) and the **low contrast** (typically 5%–50%, at most 300%; Lin & Hudson 1976; Jess et al. 2008). (<https://iopscience.iop.org/article/10.3847/1538-4357/aa9b34>)

Sources:

From the CISM Summer School (Boulder, August 2013) – SW101_4_Flares

<https://www.bu.edu/cism/SummerSchool/summerlist.html>

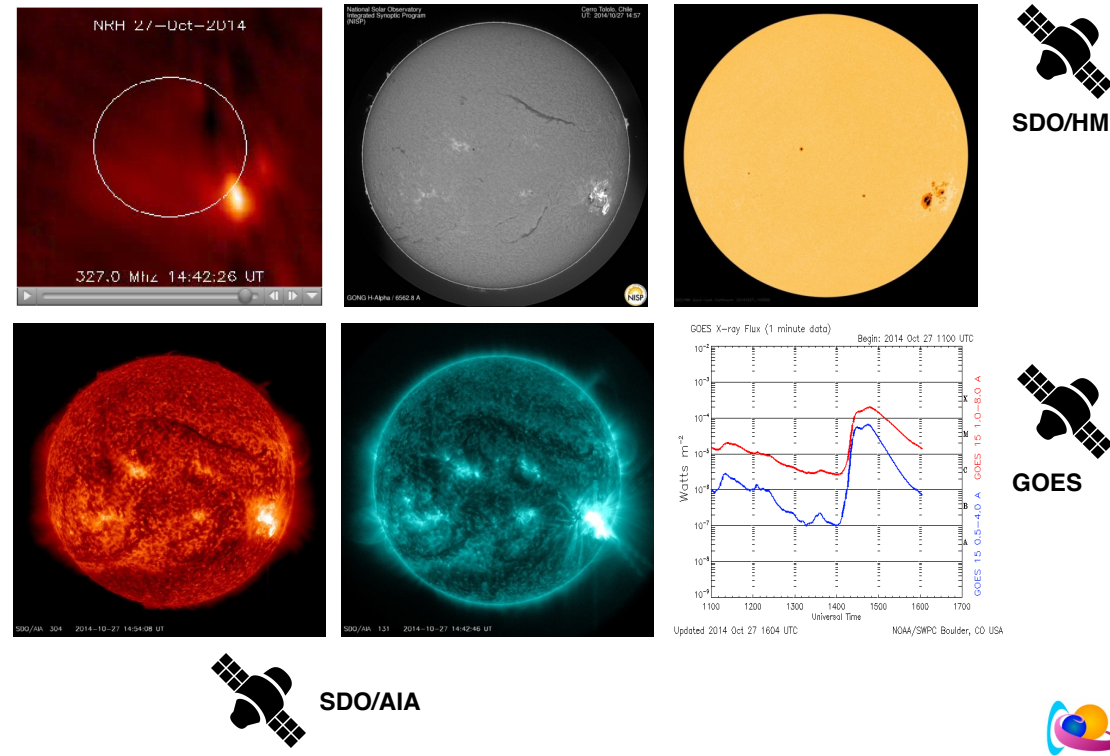
<http://solarphysics.livingreviews.org/Articles/lrsp-2011-6/>

Solar Flares: Magnetohydrodynamic Processes (Kazunari Shibata and Tetsuya Magara)

Images taken from <https://iopscience.iop.org/article/10.3847/1538-4357/aae47c>

Carrington's sunspot drawings on August 28 and September 1, shown in projected images. The whole disk drawings on August 28 and September 1 are shown above. The relevant parts of his logbook on August 28 and September 1 are shown below. These manuscripts are currently preserved in the archive of the Royal Astronomical Society.

SOLAR FLARES - OBSERVATIONS



Solar flares are sudden bursts of radiation lasting minutes – hours at wavelengths over the entire electromagnetic spectrum: Gamma-rays, HXR, SXR, EUV; H-alpha, radio

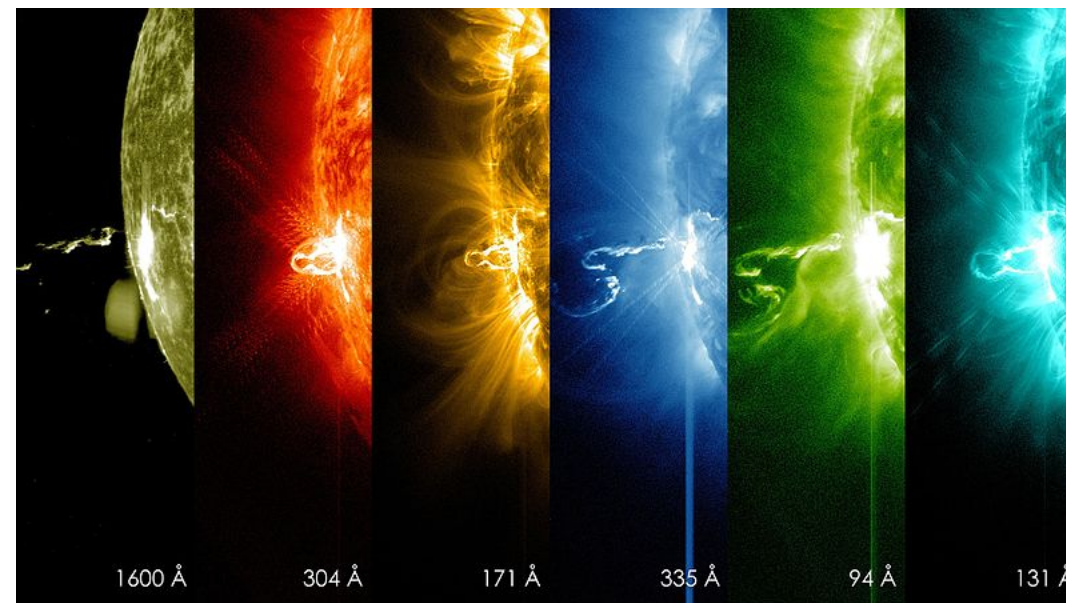
The total irradiance enhancement is dominated by white light and infra-red emission (77%). UV and soft X-ray emissions <200 nm amount to 23%.
(Living Reviews in Solar Physics – Flare Observations (Benz, 2017), <https://link.springer.com/article/10.1007/s41116-016-0004-3>)

Sources:

<http://bass2000.obspm.fr/home.php> (Nançay Radio Heliograph)

<http://www.stce.be/news/279/welcome.html>

SOLAR FLARES - OBSERVATIONS



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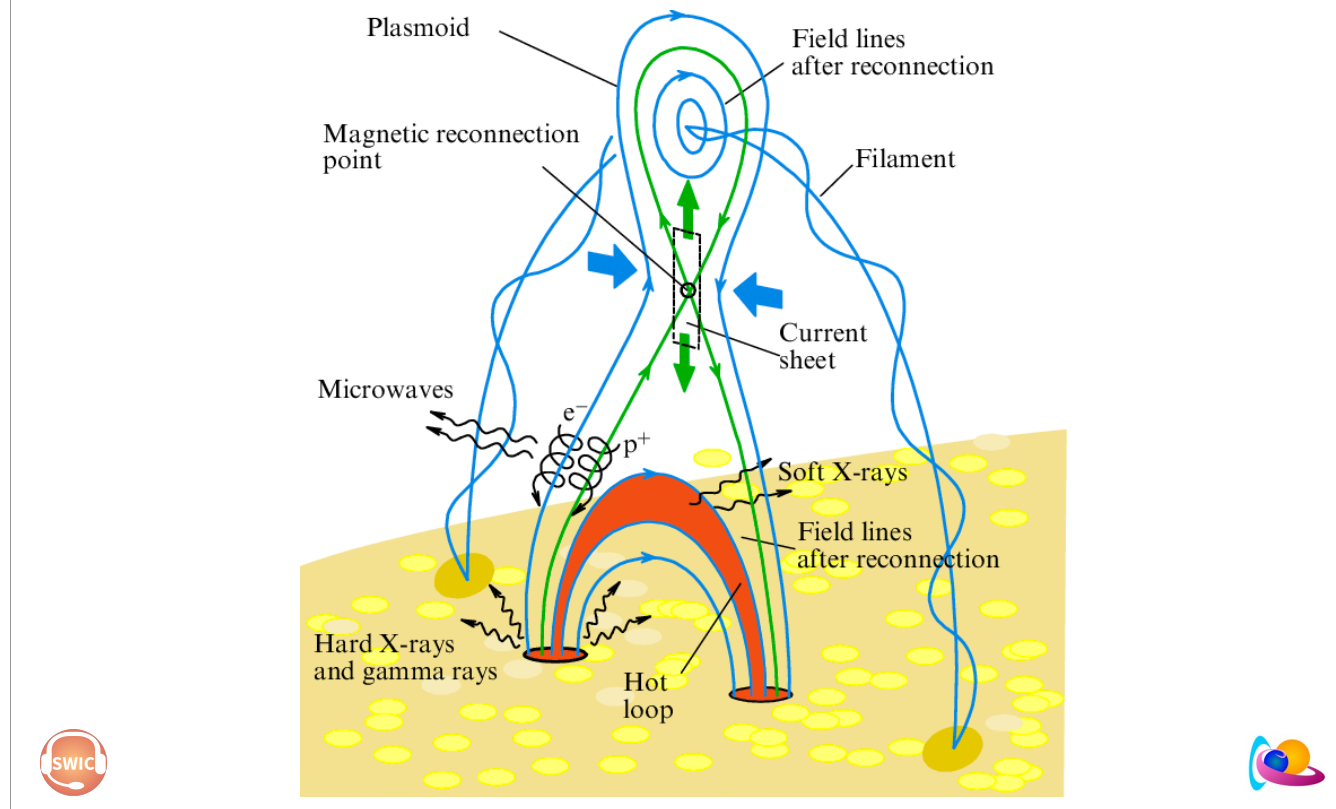
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Sources:

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<http://www.stce.be/news/279/welcome.html>

STANDARD FLARE MODEL



The CSHKP model is a model of solar flares that explains their observable features on the basis of magnetic reconnection. The basic idea of this model was proposed and developed by Carmichael, Sturrock, Hirayama, Kopp and Pneuman (Carmichael 1964; Sturrock 1966; Hirayama 1974; Kopp & Pneuman 1976), which is why this model is named CSHKP after these five scientists.

The model was originally proposed by Kopp and Pneuman and then refined. In broad lines it consists of the following phases:

- 0) Magnetic reconnection occurs
 - 1) It requires a „transient“ that opens up the magnetic field lines.
 - 2) As they close down and reconnect, energy is released that goes into accelerating electrons which travel down the magnetic field lines.
 - 3) These highly energetic particles will heat the dense chromosphere at the footpoints
 - 4) and this plasma is heated and conducted into the loops
 - 5) Post-eruption loop arcade appears successively high, because of the reconnection site rises with time
 - 6) The ribbon separates with time because of the increasing distance between footpoints due to higher loop arcades

Polarity inversion line (PIL, also called neutral line): The line that separates solar magnetic fields of opposite polarity, typically determined from solar magnetograms.

From: <https://www.swpc.noaa.gov/content/space-weather-glossary#n>

Sources:

From Maria Massi (What is a solar flare?)

<http://www3.mpifr-bonn.mpg.de/staff/mmassi/#coronae1>

Image:

https://www.researchgate.net/publication/341075910_X-ray_and_gamma-ray_emission_of_solar_flares

An animated model from a solar flare can be found at:

Cheung et al. (2018): A comprehensive three-dimensional radiative magnetohydrodynamic simulation of a solar flare

<http://adsabs.harvard.edu/abs/2018NatAs...3..160C>

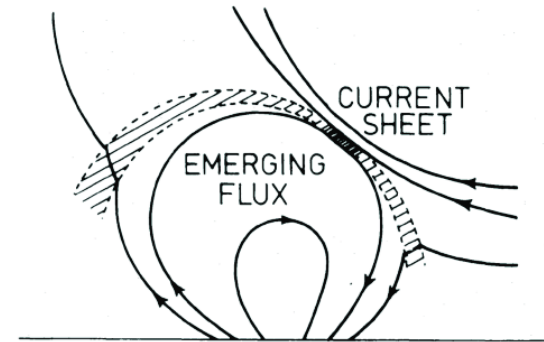
YouTube movie at https://www.youtube.com/watch?v=kyhsBqB2x_Y&feature=youtu.be

TRIGGERS OF SOLAR FLARES



Magnetic restructuring

- Magnetic flux emergence
- Helical energy storage
- Instability surrounding fields
- Interaction with nearby CH
- ...



There are various processes that can trigger a solar flare, but the underlying basic principle remains the same: magnetic reconnection transforms stored magnetic energy into kinetic and radiative energy.

Sources:

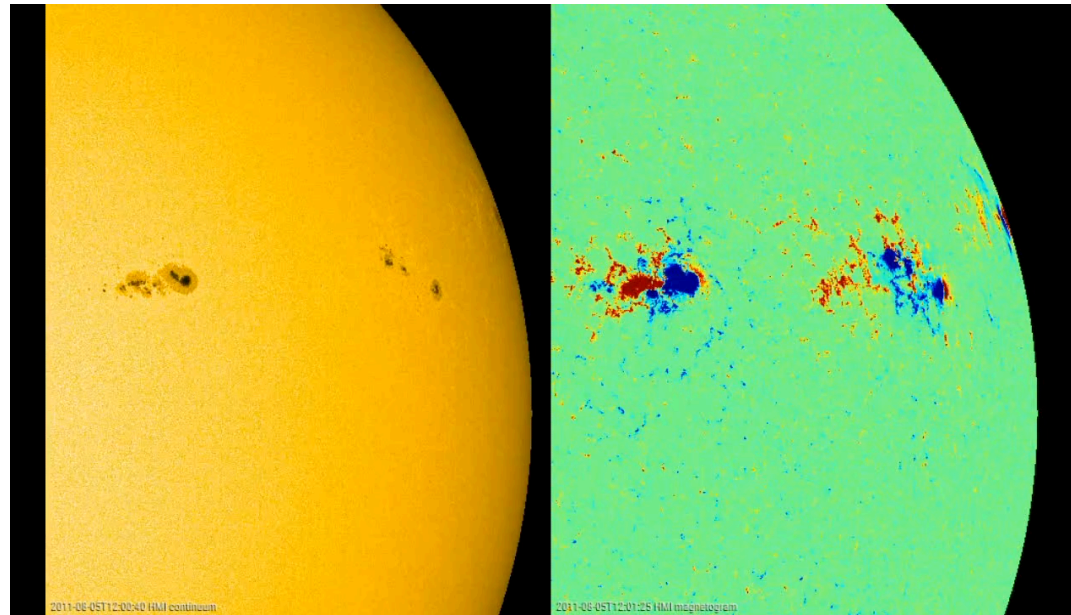
Image from: https://www.researchgate.net/publication/274542590_Flux_Emergence_Theory

Other example: Solar flare mechanism: <http://www.stce.be/news/265/welcome.html>

TRIGGERS OF SOLAR FLARES



Magnetic flux emergence



This movie shows how magnetic flux is emerging in the trail of the sunspot regions. This emerging field interacts with the existing magnetic configuration. When it creates an instability, an X-flare is observed.

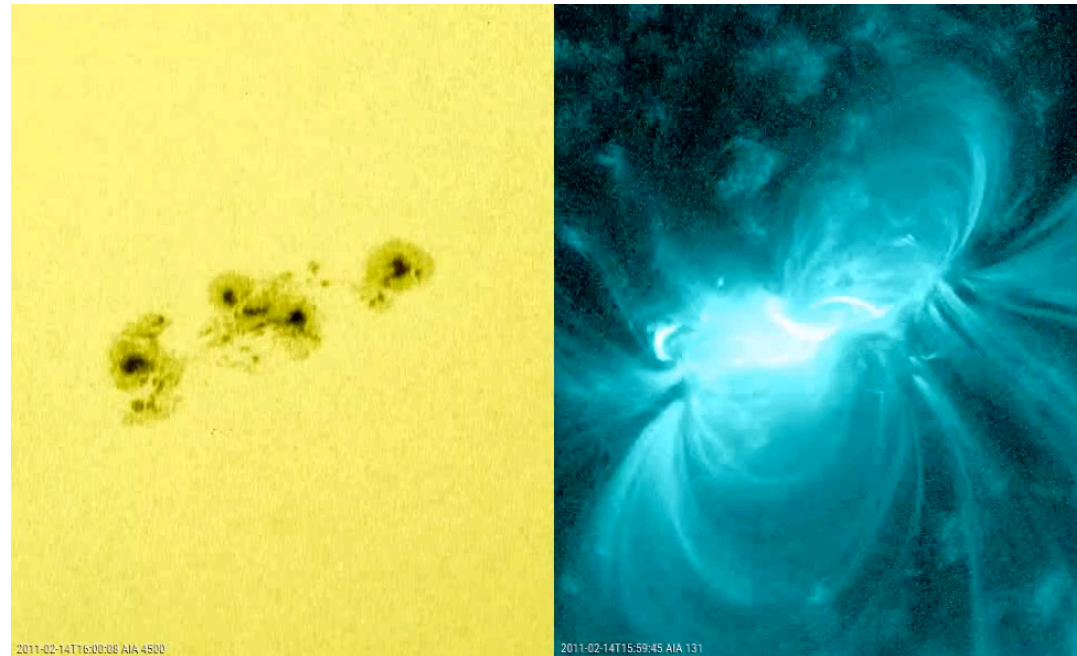
X6.9 flare on 9 August 2011: <http://www.stce.be/news/353/welcome.html>

Blue/black is negative (inward) magnetic polarity, red/white is positive (outward) polarity.

TRIGGERS OF SOLAR FLARES



Helical energy storage



This movie shows rotating sunspots that cause a solar flare.

Rotating sunspots are an extremely efficient way to inject energy into the magnetic field of the Sun's atmosphere. Twisting the Sun's magnetic field is like twisting an **elastic band**. At first you store energy in the elastic, but if you twist too much the elastic band snaps, releasing the stored energy. Similarly, rotating sunspots store energy in the Sun's atmospheric magnetic field. If they twist too much, the magnetic field breaks releasing energy in a flare.

X2.2 flare on 15 February 2011

Sources:

Velareddi et al. (2012): On the role of rotating sunspots in the activity of solar active region NOAA 11158

<http://iopscience.iop.org/article/10.1088/0004-637X/761/1/60/pdf>

Jiang et al. (2011): Rapid sunspot rotation associated with the X2.2 flare on 2011 February 15

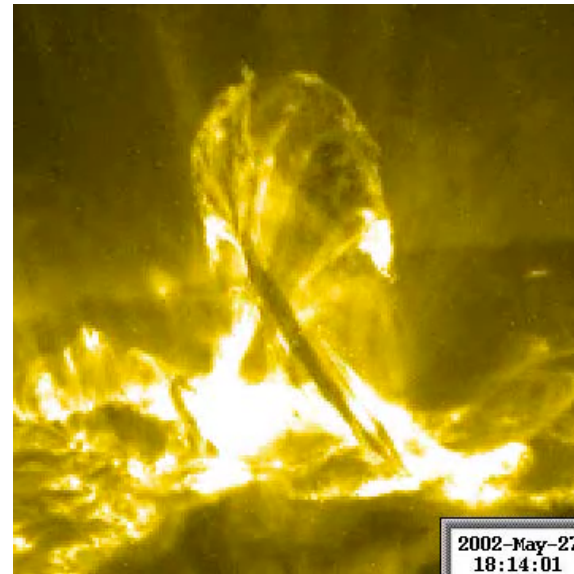
<http://iopscience.iop.org/article/10.1088/0004-637X/744/1/50>

Also at PhysOrg: <https://phys.org/news/2011-04-rotating-sunspots-super-solar-flare.html>

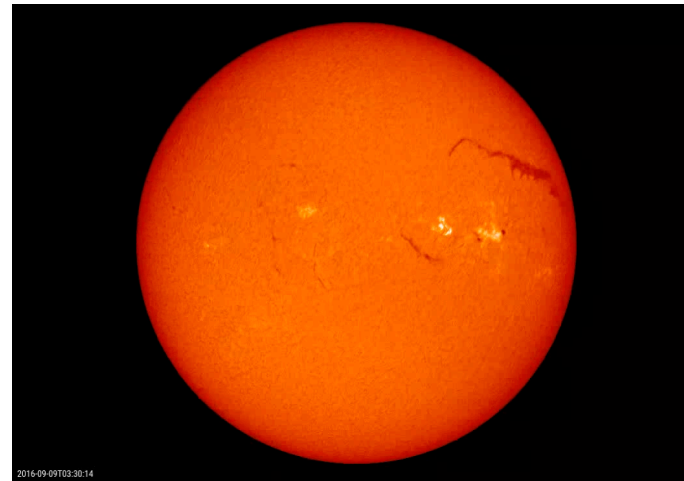
TRIGGERS OF SOLAR FLARES



Kink instability



Unstable magnetic fields



The left movies shows a kink magnetic flux tube, this is a bunch of magnetic strands that are intertwined. If it twist too much, it breaks.

The movie on the right is taken in H-Alpha and shows a B7 flare that did not only produce electromagnetic radiation. It apparently also created an imbalance in the magnetic fields. As a result, about an hour after the B7 flare, first the middle portion and then the western portion of the nearby filament disappeared entirely from view in H-alpha. The 40-degrees long filament to the upper right remained unchanged.

Sources:

Kink instability

Török et al. (2010): The writhe of helical structures in the solar corona

https://www.aanda.org/articles/aa/full_html/2010/08/aa13578-09/aa13578-09.html

<http://www.lmsal.com/TRACE/POD/TRACEpodarchive14.html#movie61> (27 May 2002; M2 ; NOAA 9957)

Unstable if twist $\sim 2.5\pi$ (Török et al., 2003: <https://www.aanda.org/articles/aa/pdf/2003/30/aah4206.pdf>).

Unstable magnetic fields

Collateral damage: <http://www.stce.be/news/361/welcome.html>

Shen et al. (2014): A Chain of Winking (Oscillating) Filaments Triggered by an Invisible Extreme-ultraviolet Wave

<http://adsabs.harvard.edu/abs/2014ApJ...786..151S>

:Issued: 2023 Mar 18 1231 UTC
:Product: documentation at <http://www.sidc.be/products/meu>
#-----#
DAILY BULLETIN ON SOLAR AND GEOMAGNETIC ACTIVITY from the SIDC #
(RWC Belgium) #
#-----#
SIDC URSIGRAM 30318
SIDC SOLAR BULLETIN 18 Mar 2023, 1230UT
SIDC FORECAST (valid from 1230UT, 18 Mar 2023 until 20 Mar 2023)
SOLAR FLARES : C-class flares expected, (probability >=50%)
GEOMAGNETISM : Quiet (A<20 and K<4)
SOLAR PROTONS : Quiet
PREDICTIONS FOR 18 Mar 2023 10CM FLUX: 138 / AP: 007
PREDICTIONS FOR 19 Mar 2023 10CM FLUX: 140 / AP: 019
PREDICTIONS FOR 20 Mar 2023 10CM FLUX: 138 / AP: 029

COMMENT: The solar flaring activity was at moderate levels during the last 24 hours with one M-class flare and several C-class flares detected, with the most frequent sources being NOAA active regions 3254 and 3256. The largest flare was a M1.1 flare, peaking at 15:07 UTC on March 17, associated with active region NOAA 3254 (beta class). NOAA active region 3256 produced an impulsive C9.4 flare at 07:10 UTC on March 18. This event was also associated with Type IV radio emission. Other regions on the disc did not show any significant flaring activity. Further M-class flare activity is possible but not probable, while frequently C-class activity is expected in the next 24 hours.

A filament eruption in the southwestern quadrant was observed on March 17 from around 09:20UTC. The associated CME appears in SoHO/LASCO C2 coronagraph data from 10:23UTC onwards. The CME is directed to the south-west and the bulk of the CME is not expected to be Earth directed. However, a glancing blow of the shock may impact Earth at around 19:00 UTC on March 19. Another small filament eruptions occurred in the northwestern quadrant from around 17:09UTC and 20:09UTC on March 17. We are awaiting corresponding coronagraph data for further analysis. During the last 24 hours there were no other potentially Earth-directed CMEs detected in the available coronagraph observations.

The greater than 10 MeV proton flux was at almost nominal levels over the past 24 hours and is expected to remain so for the next 24 hours. The greater than 2 MeV electron flux remained below the 1000 pfu alert threshold and is expected to remain below this threshold during the next 24 hours. The 24h electron fluence was at normal levels and is expected to remain so.

Over the past 24 hours the solar wind parameters (ACE and DSCOVR) have been indicative of slow solar wind conditions. The solar wind speed ranged between 400 km/s and 450 km/s. The interplanetary magnetic field magnitude was about 6 nT. The magnetic field orientation was predominantly in the positive sector (field directed away from the Sun). Similar slow solar wind regime is expected on March 18 with a slight wind speed enhancement possible for late on March 19, due to expected influence of the small equatorial coronal hole of positive polarity with a chance of being mixed with glancing blow from a CME which left the solar surface around 10 UTC on March 17th.

The geomagnetic conditions over the past 24 hours were globally and locally quiet to unsettled (NOAA Kp and K Bel 1-3). Quiet conditions are expected for March 18 with active to minor storm conditions possible for late on March 19 and March 20, due to expected arrival of the high speed stream and a possible glancing blow from a CME.

TODAY'S ESTIMATED ISN : 044, BASED ON 14 STATIONS.

SOLAR INDICES FOR 17 Mar 2023
WOLF NUMBER CATANIA : 110
10CM SOLAR FLUX : 134
AK CHAMBON LA FORET : 014
AK WINGST : 007
ESTIMATED AP : 007
ESTIMATED ISN : 073, BASED ON 18 STATIONS.

NOTICEABLE EVENTS SUMMARY
DAY BEGIN MAX END LOC XRAY OP 10CM Catania/NOAA RADIO_BURST_TYPES
17 1504 1507 1511 S22W65 M1.0 SN 12/3247
END



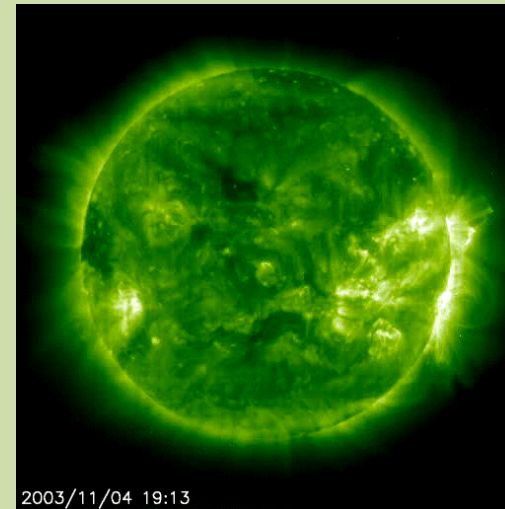
Flare features





FLARES - OVERVIEW

- Flare Characteristics
- **Flare Classification**
 - H-alpha
 - X-ray
- Flare predictions



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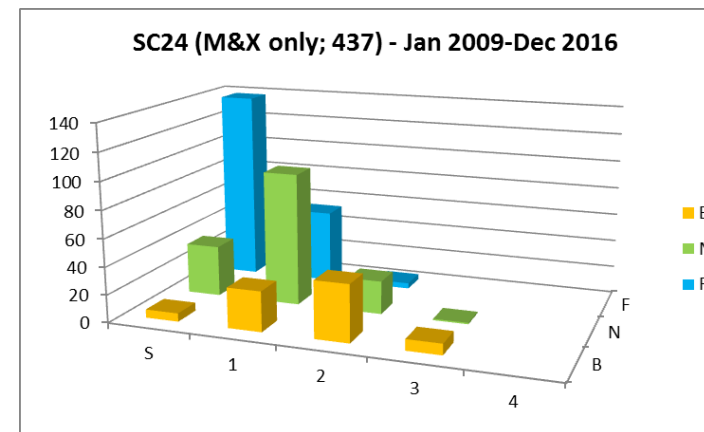


SOLAR FLARES CLASSIFICATION

H-alpha

- Visual classification
- Importance
 - Area at max. brightness
- Brightness
 - Faint, Normal, Brilliant
- Optical class
 - E.g. SF, 3B,...
- Limited correlation with geophysical effects
- Depends strongly on seeing conditions

Importance	A_c (MH)	A_c (°²)
S	$10 \leq A_c < 100$	$0,2 \leq A_c < 2,1$
1	$100 \leq A_c < 250$	$2,1 \leq A_c < 5,2$
2	$250 \leq A_c < 600$	$5,2 \leq A_c < 12,4$
3	$600 \leq A_c < 1200$	$12,4 \leq A_c < 24,7$
4	$1200 \leq A_c$	$24,7 \leq A_c$



Optical Information (Op): The optical classification and location of an associated flare, observed in H α . It contains an importance and a Brightness parameter:

* Importance is the corrected area of the flare in heliospheric square degrees at maximum brightness, observed in the H α line (656.3 nm).

S – Subflare (area ≤ 2.0 deg.²)

1 – Importance 1 ($2.1 \leq \text{area} \leq 5.1$ deg.²)

2 – Importance 2 ($5.2 \leq \text{area} \leq 12.4$ deg.²)

3 – Importance 3 ($12.5 \leq \text{area} \leq 24.7$ deg.²)

4 – Importance 4 (area ≥ 24.8 deg.²)

* Brightness is the relative maximum brightness of flare in H α .

F – faint ; N – normal ; B – brilliant

In photometric observations: **If a flaring region reaches a brightness of 1.6 times (160%) the surrounding quiet sun background brightness, then it is considered a faint (F) flare.** If the flaring region reaches a brightness **between 160% and 360%** the background brightness, it is catalogued as a flare of **normal** intensity (N). Finally, in order to be considered a "brilliant" flare (B), the flaring region should reach an intensity of 360% the background brightness and its area should cover at least 10 millionths of the solar hemisphere.

* Location (°Lat. °CMD) gives the spherical, heliographic coordinates of the solar flare in H α . The field is blank for x-ray events with no optical correlation (no optical flare observed or no optical patrol at the time) and for flares that occasionally occur in unassigned regions).

Note however, that the observed intensity is **strongly dependent on the seeing conditions**, and only a slight amount of atmospheric pollution can drastically alter the measured intensity.

This classification is still **widely used**, e.g. in the daily SWPC (event) reports, The Weekly, the SIDC's Ursigrams and weekly bulletins,...

Sources:

H-alpha flare classification: Australian SWS: <http://www.sws.bom.gov.au/Educational/2/4/2>

H-alpha observing: <http://users.telenet.be/j.janssens/Halpha/Halfaeng.html#Flares>

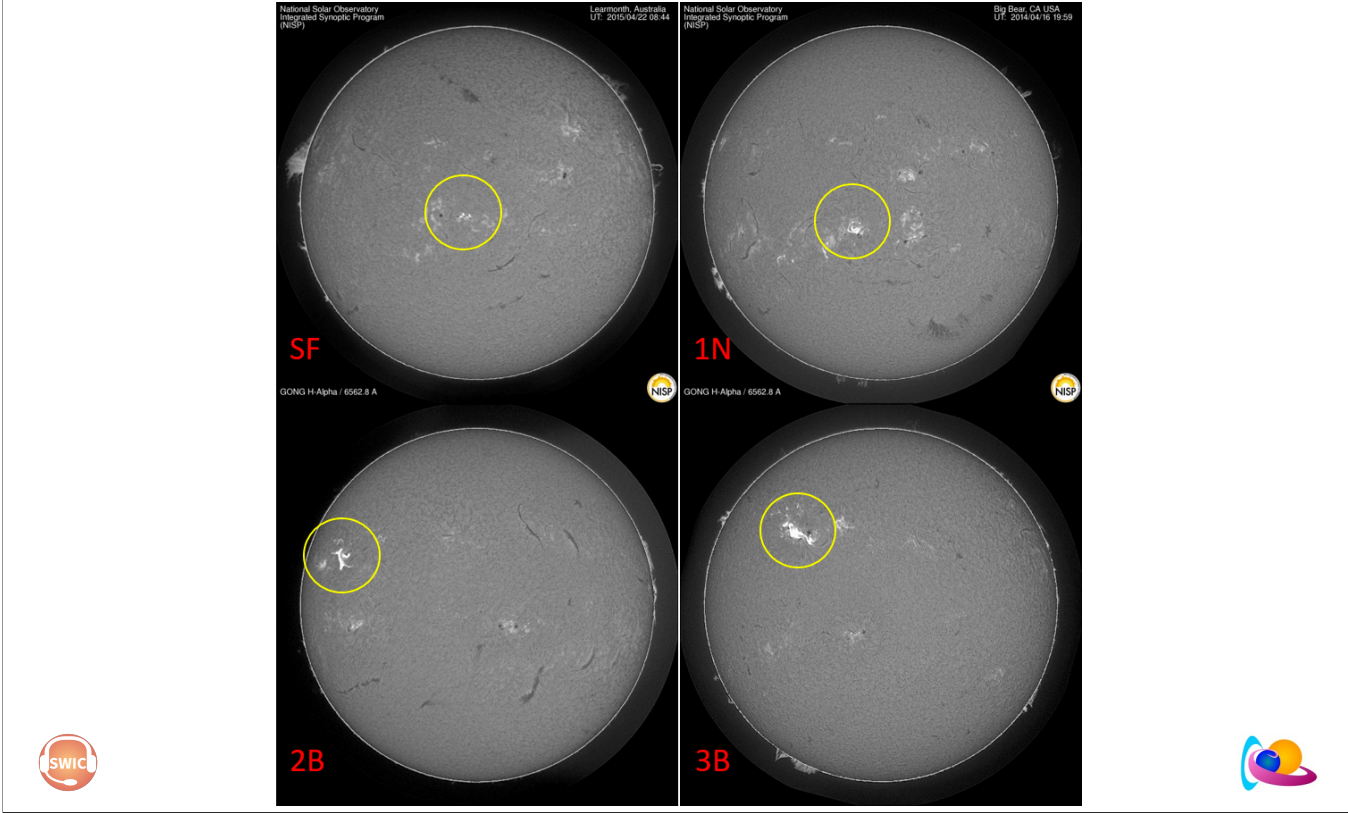
From SWPC's « The Weekly » User guide (https://www.swpc.noaa.gov/sites/default/files/images/u2/Usr_guide.pdf ; page 4)

A detailed analysis of H-alpha flare properties is by Temmer et al. (2001): Statistical analysis of solar H flares

<https://www.aanda.org/articles/aa/pdf/2001/33/aa1413.pdf>

See also: <https://www.stce.be/educational/classification>

SOLAR FLARE CLASSIFICATION: H-ALPHA

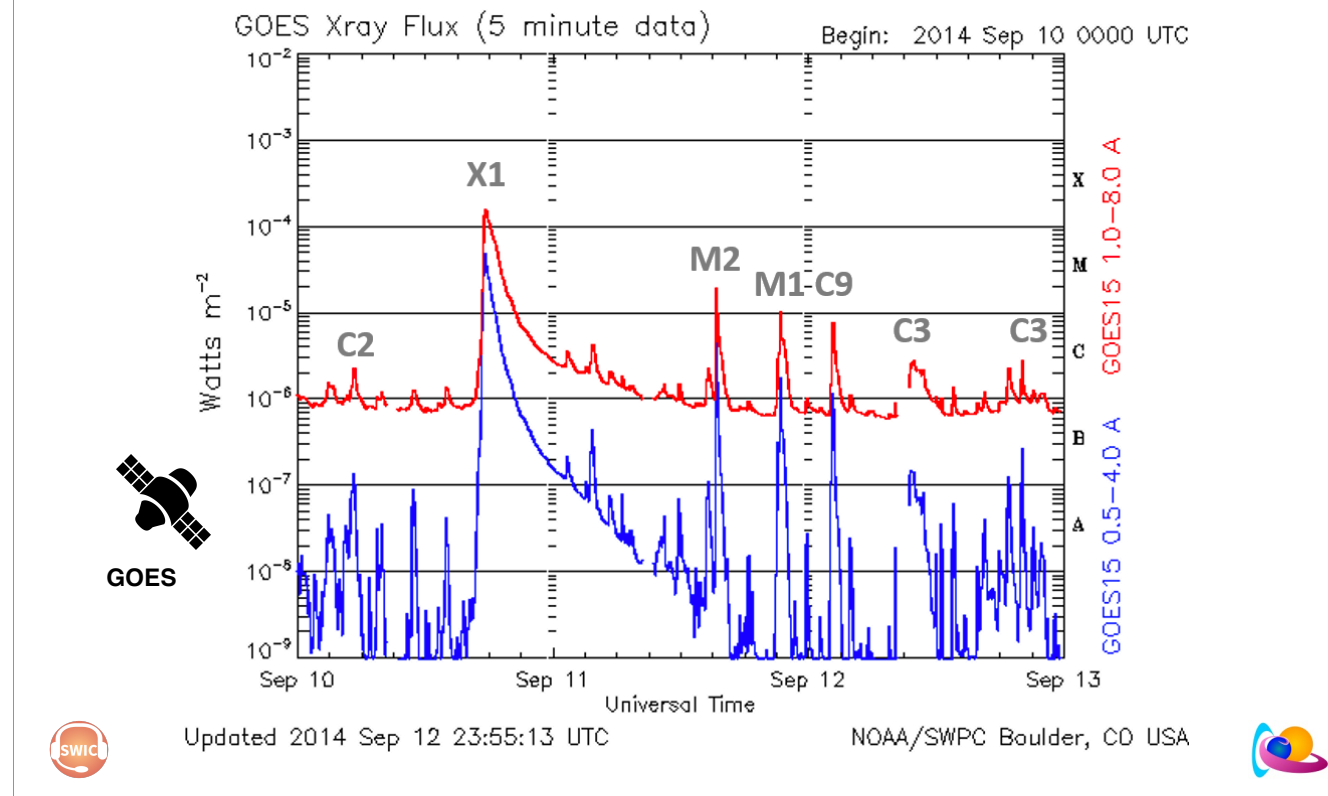


NOTICABLE EVENTS SUMMARY									
DAY	BEGIN	MAX	END	LOC	XRAY	OP	10CM	Catania/NOAA	RADIO_BURST_TYPES
SF: 22 Apr 2015	- 0830	0844	0858	S09E05	M1.1	SF			
1N: 16 Apr 2014	-1954	1959	2004	S14E09	M1.0	1N	24/2035		II/2
2B: 02 Apr 2014	- 1318	1405	1428	N14E53	M6.5	2B	3700	09/2027	II/1IV/2
3B: 07 Mar 2012	- 0002	0024	0040	N17E27	X5.4	3B	7200		IV/1,II/2,V/2

Data are from the SIDC / Daily Ursigrams (<http://www.sidc.be/archive>)
Images are from GONG/NSO H-alpha Network (<ftp://gong2.nso.edu/HA/hag/>)



SOLAR FLARE CLASSIFICATION: X-RAY



Source: <http://iopscience.iop.org/article/10.1086/304521/fulltext/36016.text.html>

From SWPC 's « The Weekly » User guide (https://www.swpc.noaa.gov/sites/default/files/images/u2/Usr_guide.pdf ; page 2)

The letter classification of solar flares was initiated on 01 January 1969. This classification ranks solar activity by its peak x-ray intensity in the 0.1–0.8 nm band as measured by the Geostationary Operational Environmental Satellites (GOES). This x-ray classification offers at least two distinct advantages compared with the standard optical classifications: it gives a **better measure of the geophysical significance** of a solar event, and it provides an objective means of classifying geophysically significant activity **regardless of its location** on the solar disk.

Table 1. The SWPC x-ray flare classification

Peak Flux Range (0.1–0.8 nm)

Classification	mks system (W m ⁻²)	cgs system (erg cm ⁻² s ⁻¹)
A	$\Phi < 10^{-7}$	$\Phi < 10^{-4}$
B	$10^{-7} \leq \Phi < 10^{-6}$	$10^{-4} \leq \Phi < 10^{-3}$
C	$10^{-6} \leq \Phi < 10^{-5}$	$10^{-3} \leq \Phi < 10^{-2}$
M	$10^{-5} \leq \Phi < 10^{-4}$	$10^{-2} \leq \Phi < 10^{-1}$
X	$10^{-4} \leq \Phi$	$10^{-1} \leq \Phi$

The letter designates the order of magnitude of the peak value and the number following the letter is the multiplicative factor. A C3.2 event for example, indicates an x-ray burst with $3.2 \times 10^{-6} \text{ W m}^{-2}$ peak flux. Solar flare forecasts are usually issued only in terms of the broad C, M, and X categories. Since x-ray bursts are observed as a full-Sun value, bursts below the x-ray background level are not discernible. The **background** drops to class A level during solar minimum; only bursts that exceed B1.0 are classified as x-ray events. During solar maximum the background is often at the class M level, therefore class A, B, or C x-ray bursts cannot be discerned. Data are measured by the NOAA GOES satellites, monitored in real time in Boulder (Grubb 1975).

The C is often referred to as « Common », M as « Medium (or moderate) », and X as « eXtreme »

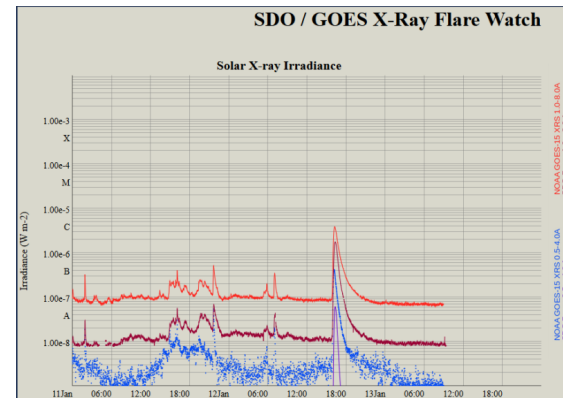
See also: <https://www.stce.be/educational/classification>

Note the peak intensity of a flare is based on the 1-minute averages by GOES, not on the **5-minute averages** which are often used for the graphs. Also, the peak intensity may not be rounded up, as the true flux never reached that level (e.g. a C5.8 is a C5 flare, not a C6 event).

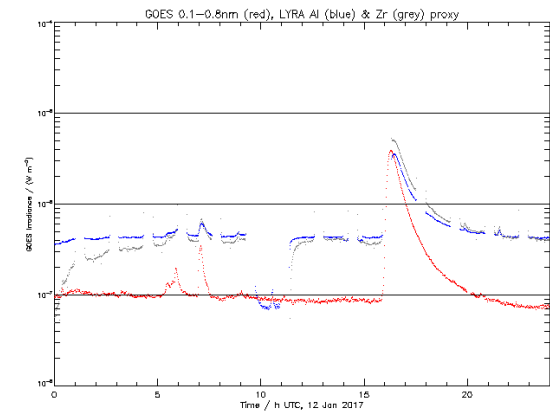


SOLAR FLARE CLASSIFICATION: X-RAY

Back-up for GOES x-ray



SDO/EVE



PROBA2/LYRA



SDO/EVE: http://lasp.colorado.edu/eve/data_access/sdo-goes-eve-flare-watch/index.html

PROBA2/LYRA: <http://proba2.oma.be/ssa>

These instruments measure the solar EUV output which is then scaled to GOES so that they can be reliably compared and substituted. In that sense, the scaled EUV measurements are proxies for the GOES x-ray measurements.



SOLAR FLARE CLASSIFICATION: X-RAY

Frequency terminology

- Solar (flaring) activity
- For a 24 hour period

Terms Used to Describe Solar Activity

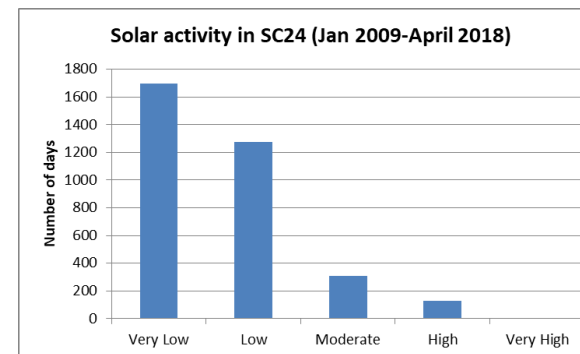
Very Low: x-ray events less than C-class.

Low: C-class x-ray events.

Moderate: isolated (one to four) M-class x-ray events.

High: several (5 or more) M-class x-ray events, or isolated (one to four) M5 or greater x-ray events.

Very High: several (5 or more) M5 or greater x-ray events.



PRF User Guide – August 2012



Source: https://www.swpc.noaa.gov/sites/default/files/images/u2/Usr_guide.pdf

Solar Activity in SC24 (Jan 2009 – Dec 2016)

Very Low 1691 days

Low 1214 days

Moderate 299 days

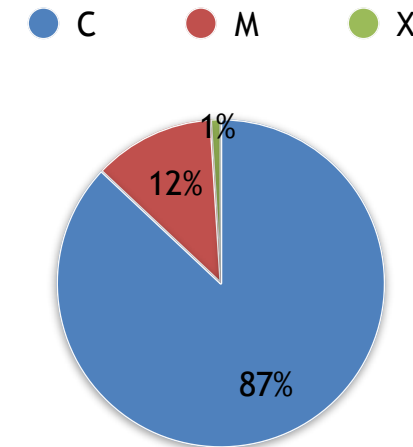
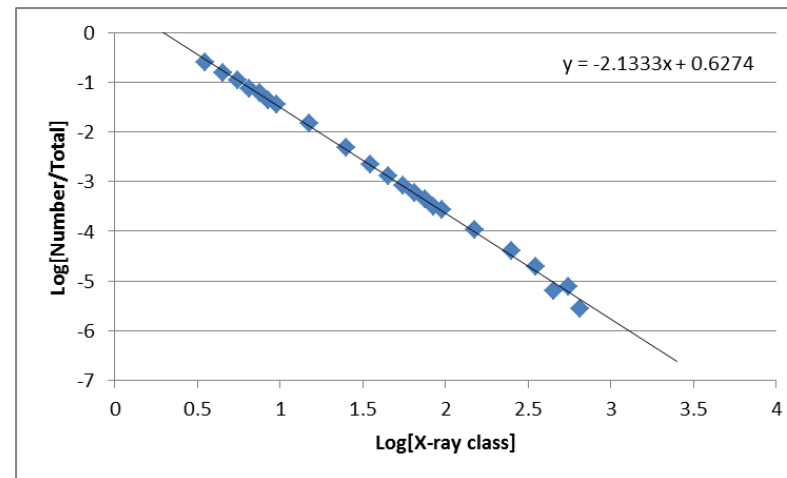
High 118 days

Very High 0 days

See also: <https://www.stce.be/educational/classification>



SOLAR FLARE STATISTICS: X-RAY



The above chart shows for each bin of solar flare intensity (C3–X6) the ratio of the number of flares for that bin vs. the total number of flares (25031 flares; January 1976 – May 2016). Both axes are logarithmic in nature. The C2 and lower classes were omitted as these numbers are affected during high solar activity (high x-ray background). The X7 and higher intensities were omitted for not sufficient data.

The linear expression between these two quantities is $y = 0.6274 - 2.1333x$

This means that **the number of flares N for a bin can be calculated from a power law equation**: $N = 25031 \cdot \delta^{4.24 - 2.13}$, with delta equalling 1, 10 or 100 for the resp. class C, M or X.

Another rule of thumb: Since 1976, there have been a total of 55000 x-ray flares. About 48000 were C-class flares, 6500 were M-class flares, and 500 were X-class flares. Or in percentages: **For every 100 solar flares, there are 87 C-class flares, 12 M-class flares, and 1 X-class flare.**

More on this (for the period 1976–1993) is at the Australian SWS: <http://www.sws.bom.gov.au/Educational/2/4/5>



SOLAR FLARE CLASSIFICATION: X-RAY

NOAA-scales: R-scale

Scale	Description	Effect	Physical measure	Average Frequency (1 cycle = 11 years)
R 5	Extreme		X20 (2×10^{-3})	Less than 1 per cycle
R 4	Severe		X10 (10^{-3})	8 per cycle (8 days per cycle)
R 3	Strong		X1 (10^{-4})	175 per cycle (140 days per cycle)
R 2	Moderate		M5 (5×10^{-5})	350 per cycle (300 days per cycle)
R 1	Minor		M1 (10^{-5})	2000 per cycle (950 days per cycle)

R = Radio Blackout



From the SWPC webpage:

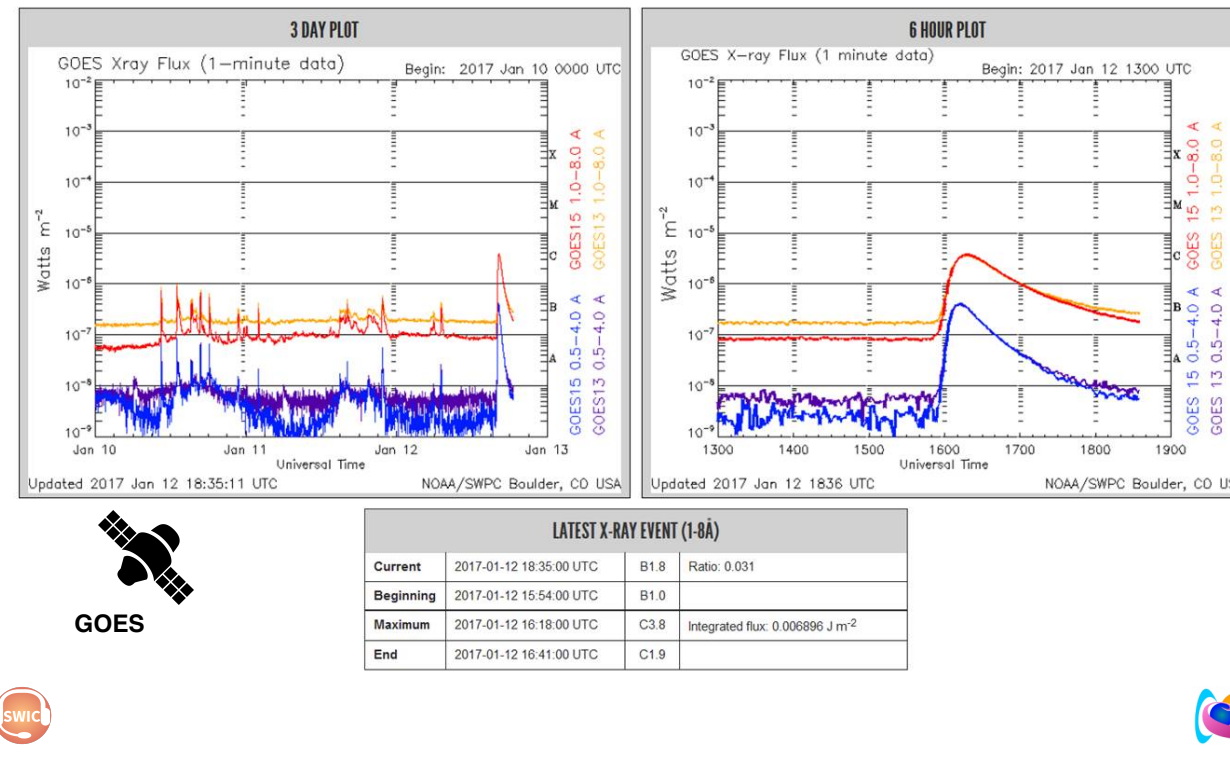
NOAA Space Weather Scales

The NOAA Space Weather Scales were introduced as a way to communicate to the general public the current and future space weather conditions and their possible effects on people and systems. Many of the SWPC products describe the space environment, but few have described the effects that can be experienced as the result of environmental disturbances. These scales are useful to users of our products and those who are interested in space weather effects. The scales describe the environmental disturbances for three event types: geomagnetic storms, solar radiation storms, and radio blackouts. The scales have numbered levels, analogous to hurricanes, tornadoes, and earthquakes that convey severity. They list possible effects at each level. They also show how often such events happen, and give a measure of the intensity of the physical causes.

The « R » stands for Radio Blackout. Note it starts only from M1 class flares and higher.

More at <http://www.stce.be/news/366/welcome.html>

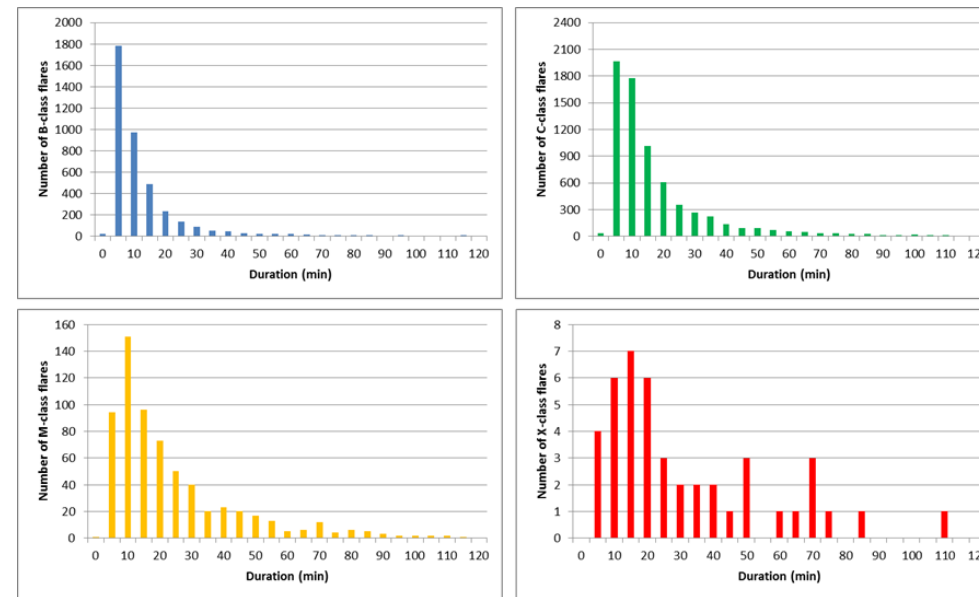
SOLAR FLARE DURATION



From SWPC 's « The Weekly » User guide (https://www.swpc.noaa.gov/sites/default/files/images/u2/Usr_guide.pdf ; page 15)

The **start** of an x-ray event is defined as the **first minute in a sequence of 4 minutes** of steep monotonic increase in 0.1-0.8 nm flux. The time of x-ray maximum is defined as the time tag of the peak 1-minute averaged value x-ray flux. The **end time** is the time when the flux level **decays to a point halfway (1/2 peak) between the maximum flux and the pre-flare background level**.

SOLAR FLARE DURATION



Jan 1976 - Dec 2000			Jan 2009 - Nov 2015		
Class	Number	Median	Class	Number	Median
B	8844	10	B	4041	10
C	16507	12	C	7015	14
M	1331	24	M	659	19
X	63	30	X	45	24
T	26745	12	T	11760	13



These graphs show the distribution of the duration (in minutes) of flares in the different categories. It is clear that stronger flares tend to last longer.

Sources:

From SWPC 's « The Weekly » User guide (https://www.swpc.noaa.gov/sites/default/files/images/u2/Usr_guide.pdf ; page 15)

From **Temporal aspects and frequency distributions of solar soft X-ray flares**

Veronig et al. (2002): <https://www.aanda.org/articles/aa/pdf/2002/06/aa1910.pdf>

And from **The duration of solar flares**

<http://www.stce.be/news/332/welcome.html>

SOLAR FLARE DURATION



Impulsive flare

- M- and X-class only
- Duration
 - Total duration < 10 minutes
- Usually NOT associated with CMEs
- Compact

Long Duration Event

- All flare classes
- Duration
 - Total duration > 1 hour
 - Decay time > 30 minutes (SWPC)
- Association with CMEs increases with increased duration



From the SWPC glossary at <https://www.swpc.noaa.gov/content/space-weather-glossary#longduration> (operational definition)

Long duration X-ray events (LDE) are not impulsive in appearance. The exact time threshold separating impulsive from long-duration events is not well defined, but **operationally**, any event requiring 30 minutes or more to decay to one-half peak flux is regarded as an LDE. It has been shown that the likelihood of a coronal mass ejection increases with the duration of an x-ray event, and becomes **virtually certain for durations of 6 hours or more**.

SOLAR FLARE DURATION

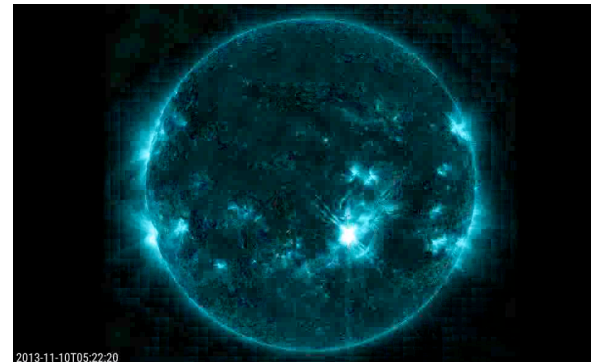


Impulsive flare

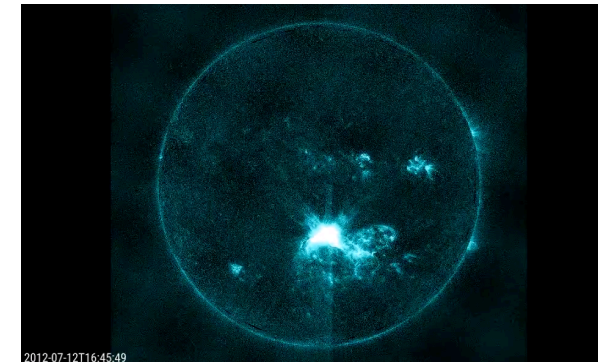
Long Duration Event



SDO/AIA



X1 - NOAA 1890 - 10 Nov 2013 (duration: 10 minutes)



X1 - NOAA 1520 - 12 July 2012 (duration: 113 minutes)



Imagery from STCE: <http://www.stce.be/news/332/welcome.html>

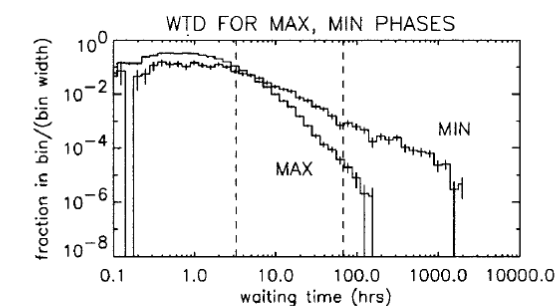
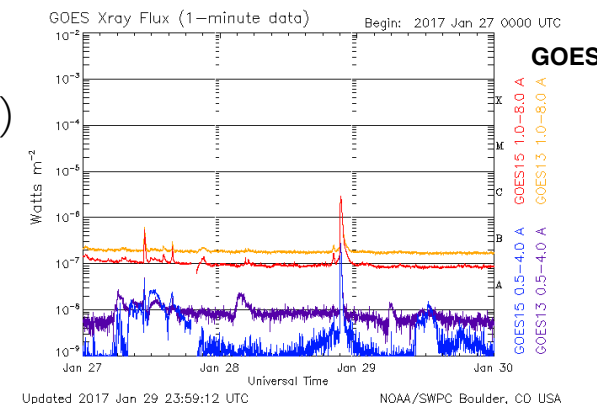
A short and a long duration X1 flaring event. These took place resp. in NOAA 11890 on 10 November 2013 (duration: 10 minutes) and in NOAA 11520 on 12 July 2012 (duration: 113 minutes or nearly 2 hours). The latter was accompanied by a full halo CME (no surprise), but also the 2013 X1 flare was associated with a partial halo CME.

SOLAR FLARE FREQUENCY



Isolated flare

- Usually specified per class (B, C, M, X)
- From entire Sun
 - 1-4 per day
- In practice
 - 1 event in 24 hours
- Average waiting time
 - 6.5 hours ($\geq$ C1)
 - SCmax: 3 hours
 - SCmin: 3 days



From SWPC 's « The Weekly » User guide (https://www.swpc.noaa.gov/sites/default/files/images/u2/Usr_guide.pdf ; page 1)

Terms Used to Describe Solar Activity

Very Low: x-ray events less than C-class.

Low: C-class x-ray events.

Moderate: isolated (one to four) M-class x-ray events.

High: several (5 or more) M-class x-ray events, or isolated (one to four) M5 or greater x-ray events.

Very High: several (5 or more) M5 or greater x-ray events.

Wheatland et al. (2002): Understanding solar flare waiting-time distributions
http://www.physics.usyd.edu.au/wheat/papers/pdfs/understanding_WTD.pdf

The bottom figure shows the Waiting Time Distribution for all years 1975–2001, and reproduces the power-law tail reported by Boffeta et al. (1999). The distribution for the maximum phase of the solar cycle has a steeper distribution, because the rate of flaring is higher around solar maximum, and so the average waiting time is less. The average waiting times for the two phases are indicated by the dashed vertical lines (average waiting time for all flares regardless the phase of the cycle is 6,5 hours). The maximum and minimum distributions both exhibit approximate power-law tails.

From <http://users.telenet.be/j.janssens/Archives/Archives.html#021109>

The longest stretch without C-class flares was from 3 April till 3 November 2008, that's 214 consecutive days of very low activity.

Since the start of systematic GOES observations, there have been only 9 periods with more than 60 consecutive days with no C-class flares, 6 of those happened during the most recent SC23–SC24 minimum... The longest stretch without M-class flares was from 25 March 2008 till 19 January 2010 (665 days). ... The longest stretch without X-class flares was from 14 Dec 2006 till 15 February 2011 (1524 days).

:Issued: 2023 Mar 18 1231 UTC
:Product: documentation at <http://www.sidc.be/products/meu>
#-----#
DAILY BULLETIN ON SOLAR AND GEOMAGNETIC ACTIVITY from the SIDC #
(RWC Belgium) #
#-----#
SIDC URSIGRAM 30318
SIDC SOLAR BULLETIN 18 Mar 2023, 1230UT
SIDC FORECAST (valid from 1230UT, 18 Mar 2023 until 20 Mar 2023)
SOLAR FLARES : C-class flares expected, (probability >=50%)
GEOMAGNETISM : Quiet (A<20 and K<4)
SOLAR PROTONS : Quiet
PREDICTIONS FOR 18 Mar 2023 10CM FLUX: 138 / AP: 007
PREDICTIONS FOR 19 Mar 2023 10CM FLUX: 140 / AP: 019
PREDICTIONS FOR 20 Mar 2023 10CM FLUX: 138 / AP: 029

COMMENT: The solar flaring activity was at moderate levels during the last 24 hours with one M-class flare and several C-class flares detected, with the most frequent sources being NOAA active regions 3254 and 3256. The largest flare was a M1.1 flare, peaking at 15:07 UTC on March 17, associated with active region NOAA 3254 (beta class). NOAA active region 3256 produced an impulsive C9.4 flare at 07:10 UTC on March 18. This event was also associated with Type IV radio emission. Other regions on the disc did not show any significant flaring activity. Further M-class flare activity is possible but not probable, while frequently C-class activity is expected in the next 24 hours.

A filament eruption in the southwestern quadrant was observed on March 17 from around 09:20UTC. The associated CME appears in SoHO/LASCO C2 coronagraph data from 10:23UTC onwards. The CME is directed to the south-west and the bulk of the CME is not expected to be Earth directed. However, a glancing blow of the shock may impact Earth at around 19:00 UTC on March 19. Another small filament eruptions occurred in the northwestern quadrant from around 17:09UTC and 20:09UTC on March 17. We are awaiting corresponding coronagraph data for further analysis. During the last 24 hours there were no other potentially Earth-directed CMEs detected in the available coronagraph observations.

The greater than 10 MeV proton flux was at almost nominal levels over the past 24 hours and is expected to remain so for the next 24 hours. The greater than 2 MeV electron flux remained below the 1000 pfu alert threshold and is expected to remain below this threshold during the next 24 hours. The 24h electron fluence was at normal levels and is expected to remain so.

Over the past 24 hours the solar wind parameters (ACE and DSCOVR) have been indicative of slow solar wind conditions. The solar wind speed ranged between 400 km/s and 450 km/s. The interplanetary magnetic field magnitude was about 6 nT. The magnetic field orientation was predominantly in the positive sector (field directed away from the Sun). Similar slow solar wind regime is expected on March 18 with a slight wind speed enhancement possible for late on March 19, due to expected influence of the small equatorial coronal hole of positive polarity with a chance of being mixed with glancing blow from a CME which left the solar surface around 10 UTC on March 17th.

The geomagnetic conditions over the past 24 hours were globally and locally quiet to unsettled (NOAA Kp and K Bel 1-3). Quiet conditions are expected for March 18 with active to minor storm conditions possible for late on March 19 and March 20, due to expected arrival of the high speed stream and a possible glancing blow from a CME.

TODAY'S ESTIMATED ISN : 044, BASED ON 14 STATIONS.

SOLAR INDICES FOR 17 Mar 2023
WOLF NUMBER CATANIA : 110
10CM SOLAR FLUX : 134
AK CHAMBON LA FORET : 014
AK WINGST : 007
ESTIMATED AP : 007
ESTIMATED ISN : 073, BASED ON 18 STATIONS.

NOTICEABLE EVENTS SUMMARY

DAY	BEGIN	MAX	END	LOC	XRAY	OP	10CM	Catania/NOAA	RADIO_BURST_TYPES
17	1504	1507	1511	S22W65	M1.0	SN			12/3247

END



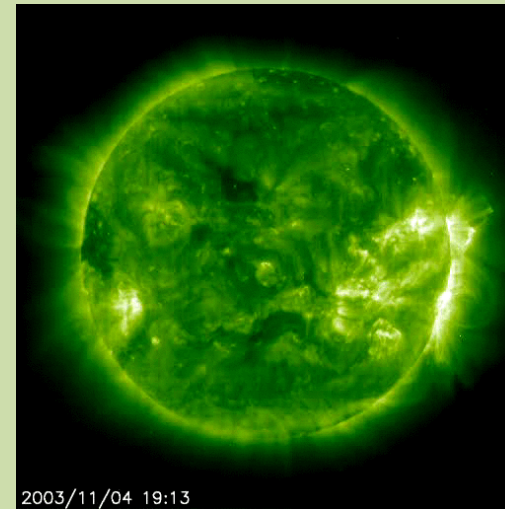
Flare classification



FLARES - OVERVIEW



- Flare Characteristics
- Flare Classification
- **Flare predictions**
 - McIntosh
 - Hale



SOHO/EIT

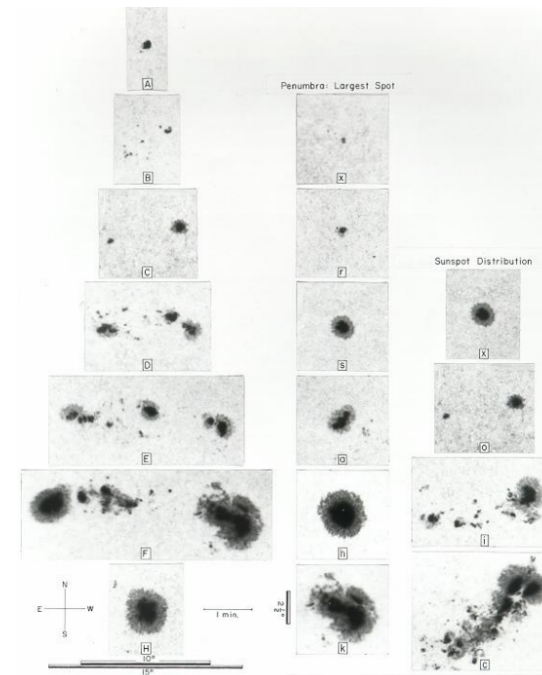


FLARE PREDICTIONS



McIntosh classification

- Zpc (*3-letter code*)
 - Z - Modified Zürich classification
 - p - Penumbra largest spot
 - c - Interior sunspot distribution
- 60 possible combinations
 - Linked to flare intensity
 - Rather large uncertainties
- Used worldwide



See: <https://www.stce.be/educational/classification>

The McIntosh classification for sunspot groups was developed by Patrick McIntosh (McIntosh, 1990) based on an earlier scheme devised by Max Waldmeier in 1947, called the Zürich classification. The general form of the McIntosh classification is Zpc (overview image above), where Z is the modified Zurich class, p is the type of principal spot, primarily describing the penumbra, and c is the degree of **compactness** in the interior of the group.

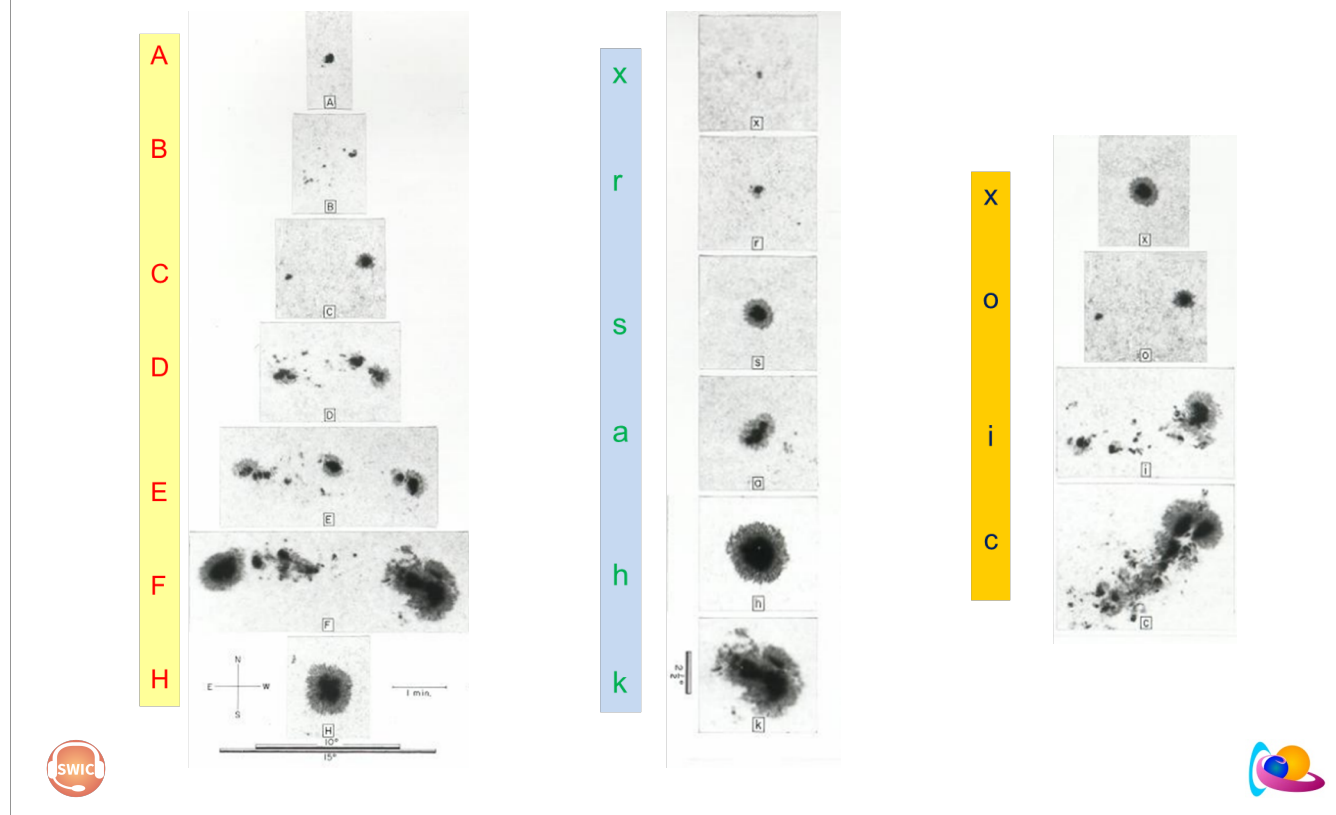
The modified **Zurich classes** are defined on the basis of whether **penumbra** is present, how penumbra is distributed, and by the **length** of the group.

McIntosh, P.S. (1990): The classification of sunspot groups

<http://adsabs.harvard.edu/abs/1990SoPh..125..251M>

<https://wwwbis.sidc.be/educational/classification.php#:~:text=The%203%20component%20McIntosh%20classification,the%20interior%20of%20the%20group.>

FLARE PREDICTIONS



McIntosh, P.S. (1990): The classification of sunspot groups
<http://adsabs.harvard.edu/abs/1990SoPh..125..251M>

Questions to ask (Table 1 from McIntosh paper)

Z – General outlook of the sunspot group:

- => Unipolar or bipolar group?
- => Penumbra or no penumbra?
- => Penumbra on one or both sides of the group?
- => Length of the group ($>10^\circ$? $>15^\circ$?)
- * 7 options: A, B, C, D, E, F, H

p – Penumbra largest spot

- => Rudimentary or mature penumbra?
- => Symmetric or asymmetric penumbra main spot?
- => N–S–diameter of the largest spot ($>2,5^\circ$?)
- * 6 options: x, r, s, a, h, k

c – Sunspot distribution interior (“compactness”)

- => Several spots between leading and trailing main spot?
- => Internally, is there at least one spot with a mature penumbra?
- * 4 options: x, o, i, c (open, intermediate, compact)



FLARE PREDICTIONS

THE ASTROPHYSICAL JOURNAL LETTERS, 747:L41 (7pp), 2012 March 10

BLOOMFIELD ET AL.

Table 2

McIntosh Classification Flare Statistics

McIntosh	SWPC (1988–1996)				Kildahl (1969–1976) ^b				Combined Flare Rate (24 hr ⁻¹)				Poisson Flare Probability (%)				
Region	Region	Total Flares			Region	Total Flares			In GOES Class				In GOES Class			Above GOES ^d	
Classes ^a	Count	C	M	X	Count	C ^c	M	X	C	M	X	±σ	C	M	X	M1.0	C1.0
ESO	95	37	6	0	82	31.9	14	0	0.39	0.11	0.00	0.08	32	11	0	11	39
ESI	18	33	1	0	78	143.0	22	2	1.83	0.24	0.02	0.10	84	21	2	23	88
EAO	459	267	61	0	47	27.3	10	4	0.58	0.14	0.01	0.04	44	13	1	14	52
EAI	295	370	83	2	82	102.8	48	1	1.25	0.35	0.01	0.05	71	29	1	30	80
EAC	3	5	1	0	17	28.3	6	3	1.67	0.35	0.15	0.22	81	30	14	39	89
EHO	42	31	6	0	39	28.8	6	0	0.74	0.15	0.00	0.11	52	14	0	14	59
EHI	15	24	6	0	45	72.0	28	4	1.60	0.57	0.07	0.13	80	43	6	47	89
EHC	2	9	0	0	4	18.0	8	0	4.50	1.33	0.00	0.41	99	74	0	74	100
EKO	185	173	35	3	52	48.6	20	1	0.94	0.23	0.02	0.06	61	21	2	22	69
EKI	423	703	173	23	81	134.6	103	11	1.66	0.55	0.07	0.04	81	42	7	46	90
EKC	103	278	132	17	63	170.0	149	21	2.70	1.69	0.23	0.08	93	82	20	85	99
FRI	0	0	0	0	2	0.0	1	0	0.00	0.50	0.00	0.71	0	39	0	39	39
FSO	14	9	3	0	13	8.4	6	1	0.64	0.33	0.04	0.19	47	28	4	31	64
FSI	6	12	0	0	8	16.0	15	0	2.00	1.07	0.00	0.27	86	66	0	66	95
FAO	73	63	16	0	3	2.6	0	0	0.86	0.21	0.00	0.11	58	19	0	19	66
FAI	91	106	35	3	12	14.0	8	0	1.16	0.42	0.03	0.10	69	34	3	36	80
FHO	9	5	1	0	10	5.6	0	0	0.56	0.05	0.00	0.23	43	5	0	5	46
FHI	10	17	9	0	18	30.6	15	0	1.70	0.86	0.00	0.19	82	58	0	58	92
FHC	0	0	0	0	5	0.0	4	0	0.00	0.80	0.00	0.45	0	55	0	55	55
FKO	97	165	29	1	19	32.3	6	0	1.70	0.30	0.01	0.09	82	26	1	27	87
FKI	235	517	161	17	47	103.4	106	17	2.20	0.95	0.12	0.06	89	61	11	66	96
FKC	93	233	146	24	27	67.6	39	13	2.51	1.54	0.31	0.09	92	79	27	84	99

Notes.
^a Only includes classifications producing ≥1 C-, M-, or X-class flare in either time range.
^b From Kildahl (1980).
^c Non-integer flare numbers result from use of observed C-class rates from SWPC (1988–1996).
^d “Above GOES X1.0” is equivalent to “In GOES Class X.”



For space weather forecasting it is important to know how this classification relates to the probability of producing a flare.

Bloomfield S. et al. (2012): Toward Reliable Benchmarking of Solar Flare Forecasting Methods
<http://adsabs.harvard.edu/abs/2012ApJ...747L..41B>



FLARE PREDICTIONS

Likelihood (terminology)

RWC Belgium: Decision on use of scales and wording

Flare forecast

ISES	
0 = Quiet (<50% probability of C-class flares) 1 = Eruptive (C-class flares expected, probability >=50%) 2 = Active (M-class flares expected, probability >=50%) 3 = Major flares expected (X-class flares expected, probability >=50%)	
Activity level	wording for bulletin
<50% probability of C-class flares	Quiet solar conditions
C-class flares expected, probability >=50%	C-class flares expected C-class flaring activity/conditions expected we expect solar active conditions ([C,M,X]-class flares) with a high/small probability for a [C, M, X]-flare
M-class flares expected, probability >=50%	M-class flares expected / idem as above
X-class flares expected, probability >=50%	X-class flares expected / idem as above

Unofficial terminology
deduced from various
SWx reports/services



Probability (%)	Terminology
0-10	Unlikely
10-25	Small chance
25-50	A chance; Possible
50-75	Likely
75-100	Very likely; Expected

This is a topic under continued improvement / discussion....



Nearly equivalent terms are used by the different space weather service providers (SWPC, SIDC,...)

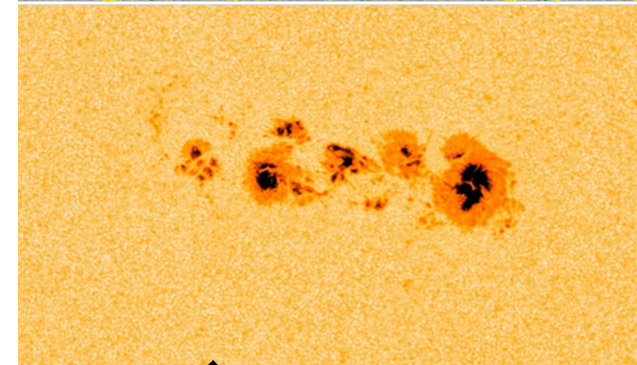
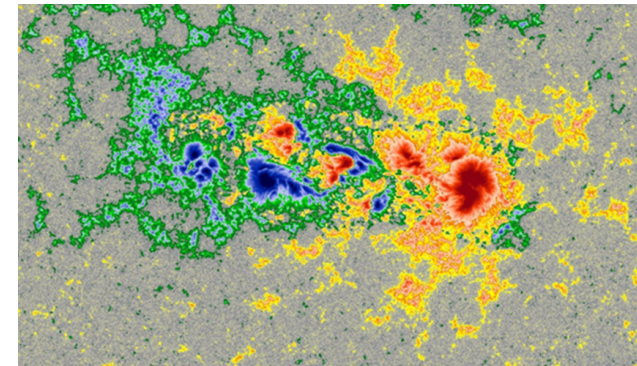
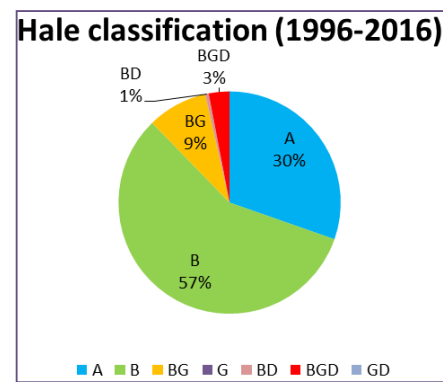
As far as we know, there's no clear terminology consistently applied internally or between the space weather prediction services.



FLARE PREDICTIONS

Mount Wilson (Hale) classification

- Based on the magnetic properties of a sunspot group
- 7 options:
 - A (α)
 - B (β), BG ($\beta\gamma$), G (γ)
 - BD ($\beta\delta$), BGD ($\beta\gamma\delta$), GD ($\gamma\delta$)



SDO/HMI



Jaeggli and Norton (2016): The Magnetic Classification of Solar Active Regions 1992–2015

<http://adsabs.harvard.edu/abs/2016ApJ...820L..11J>

<http://iopscience.iop.org/article/10.3847/2041-8205/820/1/L11/pdf>

Magnetic classifications provide a simple way to describe the configuration of the magnetic flux and sunspots in a solar active region (AR). The Mount Wilson (or Hale) classification system for sunspot groups put forward by Hale et al. (1919) has been used for nearly a century. In the original Hale classification scheme, the designation (**alpha**) is given to regions that contain a single sunspot or sunspot group all having the same polarity. Generally, these also have a weaker opposite polarity counterpart that is not strong or concentrated enough to produce sunspots. (**beta**) is assigned to regions that have two sunspots or sunspot groups of opposite polarity. The classification (**gamma**) is appended to the above classes to indicate the AR has a complex region of sunspots with **intermixed polarity**. This classification can also be used individually to describe an AR that has no organized magnetic behavior. As an addendum to the original scheme, Kunzel (1965) proposed an additional classification to modify the existing three. (**delta**) indicates that at least one sunspot in the region **contains opposite magnetic polarities inside of a common penumbra** separated by no more than 2° in heliographic distance (24 Mm or 33" at disk center).

Also at STCE: <http://www.stce.be/news/222/welcome.html>

Make sure to avoid classifying too quickly a sunspot group as a delta or a gamma type when this sunspot group is still very close to the **limb**. Indeed, line-of-sight may come into play that show an unipolar spot as if it would have a delta structure.

See STCE: <http://www.stce.be/news/188/welcome.html>

The pictures to the right are from SDO/HMI and show a magnetogram and a white light image of NOAA 1875 on 23 October 2013.

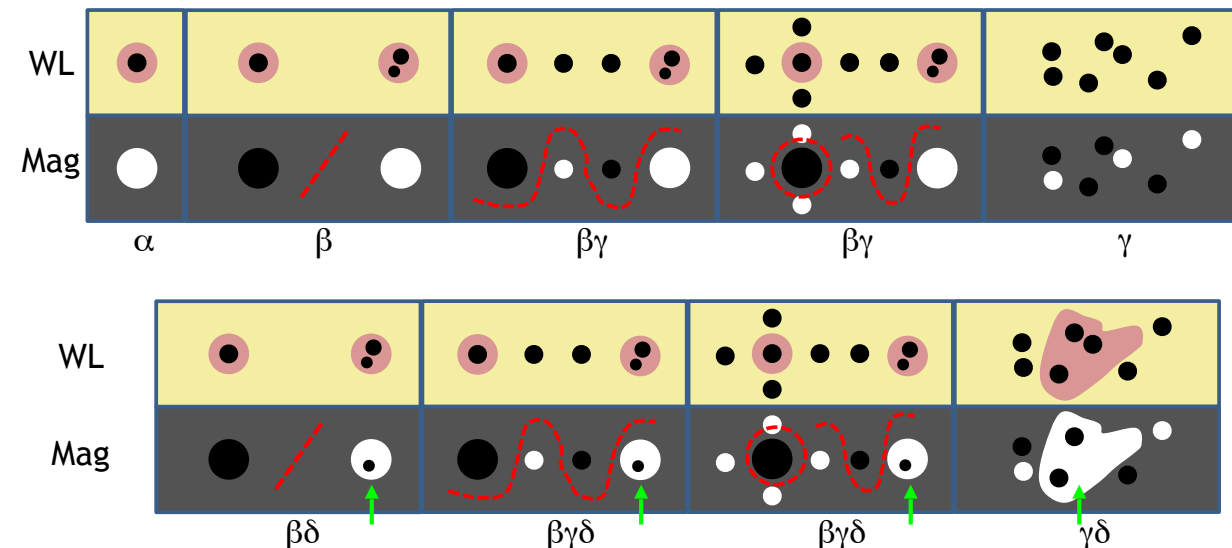
The chart with the % of Hale classification was based on the NOAA reports at <https://solarscience.msfc.nasa.gov/greenwch.shtml> for the period Jan 1996–September 2016. A total of 32965 classifications were made. The percentage of reported delta's in the sunspot groups is 3.5%.

See also: <https://www.stce.be/educational/classification>

FLARE PREDICTIONS



Mount Wilson (Hale) classification - Sketches



Text underneath based on SWPC User Guide, SWPC Glossary (<https://www.swpc.noaa.gov/content/space-weather-glossary#m>), Mount Wilson (<http://obs.astro.ucla.edu/spotlnd.html>) and SIDC old webpages (<http://sidc.oma.be/educational/classification.php#magnetic>).

Alpha – Unipolar group; that is, all plus or all minus magnetic field

Beta – A bipolar group; that is a mix of plus and minus magnetic polarities exist, with the plus well divided from the minus with one polarity in each end (E-W) of the group, i.e. “easily divided by a simple line”.

Beta-Gamma – A group which is generally bipolar but which is lacking a well marked dividing line between the opposite polarity regions (“you need to lift your pencil to divide the polarities” or “no single, continuous line can be drawn between spots of opposite polarities”).

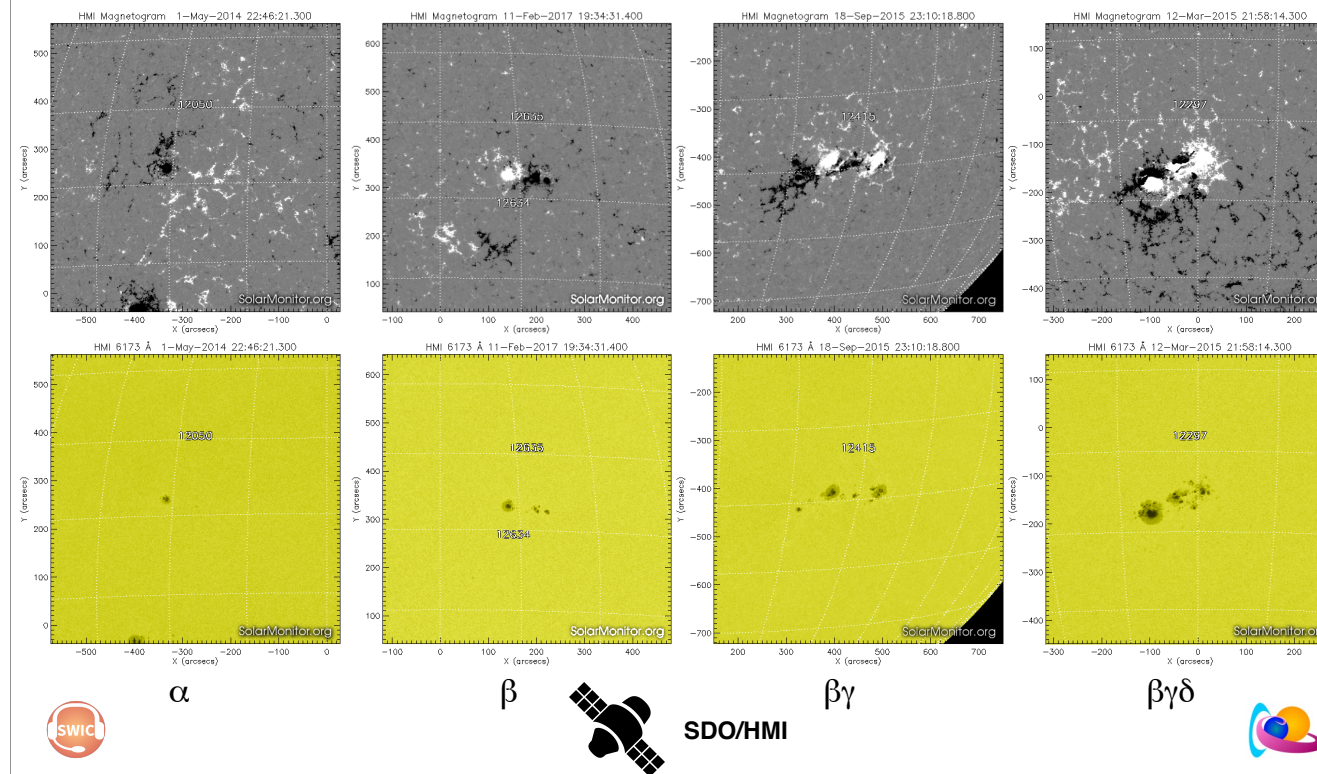
Gamma – a group in which the polarities are so completely mixed that **no bipolar structure** can obviously be recognized.

Delta – This is a sub-classification for non-unipolar regions. It means at least two opposing polarity umbrae are **within two heliographic degrees** of each other and share the same penumbra.

[The determination of the Hale class is done on the (magnetic polarity of the) sunspots, NOT the magnetograms!]



EXAMPLES OF HALE CLASSIFICATION



Examples from <https://www.solarmonitor.org/index.php>

Delta means we have opposite polarities very close together, increasing the chances of a flare significantly.

FLARE PREDICTIONS

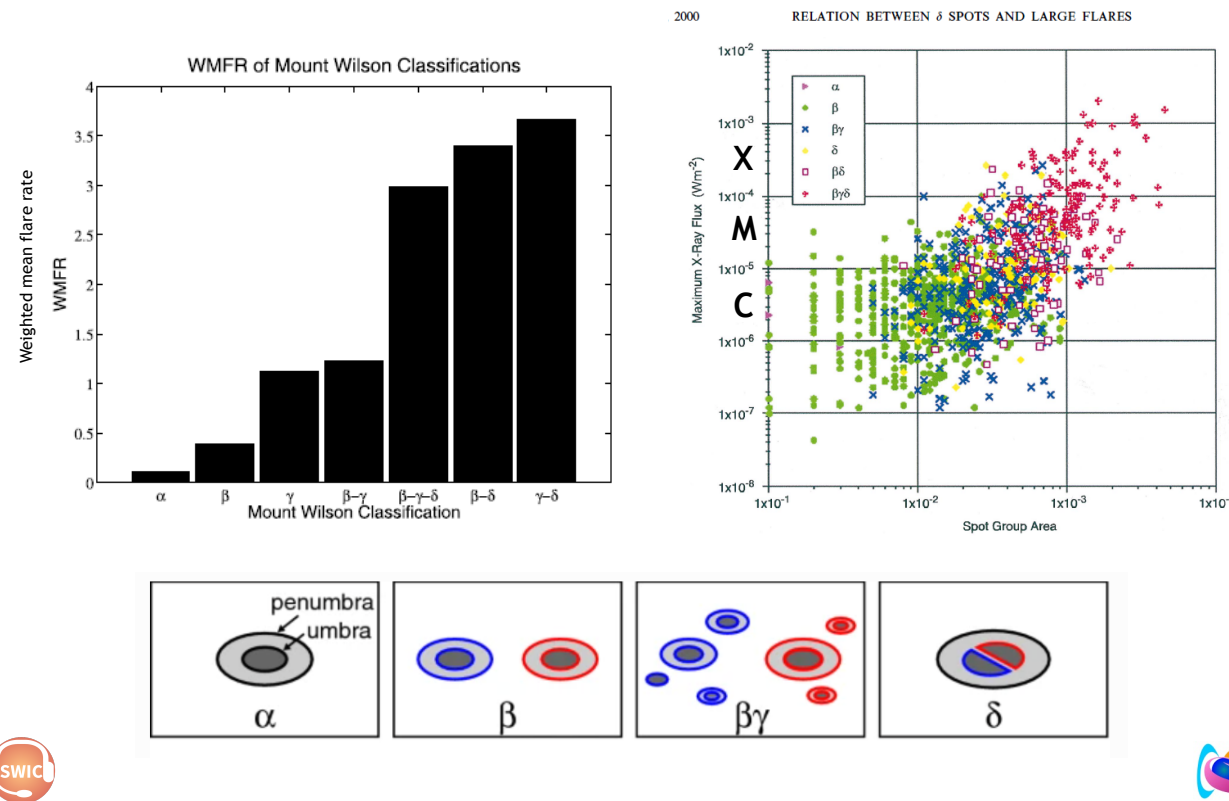


Figure left: Shin et al. (2016): Development of Daily Maximum Flare-Flux Forecast Models for Strong Solar Flares – <http://adsabs.harvard.edu/abs/2016SoPh..291..897S>

Most of the complex sunspots of the Mount Wilson magnetic classification that are characterized by gamma and/or delta show higher WMFR values (WMFR: weighted mean flare rate).

Figure right: Sammis et al. (2000): The Dependence of Large Flare Occurrence on the Magnetic Structure of Sunspots
<http://adsabs.harvard.edu/abs/2000ApJ...540..583S>
<http://iopscience.iop.org/article/10.1086/309303/pdf>

In Figure 2, we plot the largest flare from each active region against the largest reported area from that region, for each magnetic class. This shows a roughly linear connection between the logs of SXR flux and active region maximum areas. The general slope of Figure 2, upward and to the right, confirms the well-known fact that large active regions have more large flares than small ones and also tend to be more complex.

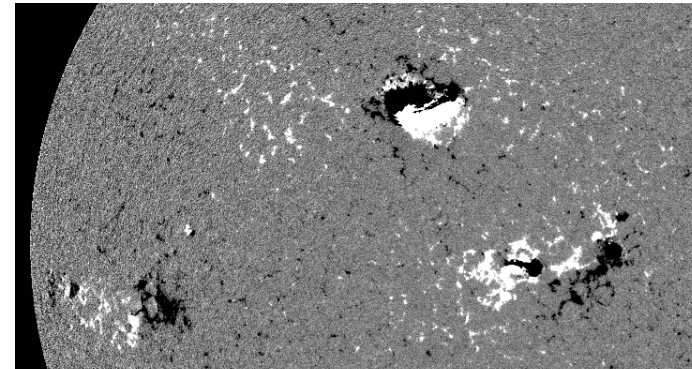
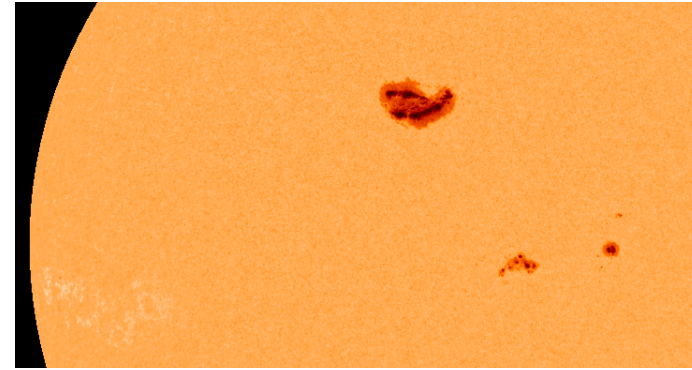
The increase in flare size with spot size shows that although the sharp gradient and currents of the delta configuration provide the appropriate situation for flare occurrence, the scale offered by a large spot is important in producing great flares. All large flares (X4 or higher) occur in spot groups of area greater than 1000 MH (micro hemispheres, **disk fraction**) classified bgd. Predictions that X1 flares will occur for such a class will enjoy a 41% probability of success with no other considerations. Adding some of the considerations mentioned by Zirin & Liggett (1987) and Zirin & Marquette (1991), particularly H-alpha brightness and flux emergence, should improve these predictions considerably.

Bottom figure from <https://link.springer.com/article/10.1007/s41116-019-0019-7>

FLARE PREDICTIONS



- Magnetic shear
- Magnetic helicity
- Flaring history
- Evolution of group
- Size of the spots
- ...



Massi et al.: <http://www3.mpifr-bonn.mpg.de/staff/mmassi/c4-Model.pdf>

⇒Magnetic shear: the vector magnetic field is oriented more parallel to the neutral line than perpendicular to it.
(Shear: laterally shifted in relation to each other.)

From <https://solarscience.msfc.nasa.gov/flares.shtml>

Stable sunspots tend to be fairly symmetrical unless there is extensive magnetic shear nearby from emerging magnetic flux or the passing of an area of opposite magnetic polarity. Magnetic shearing can cause large portions of sunspot penumbras to distort or vanish.

Lee et al. (2012): Solar Flare Occurrence Rate and Probability in Terms of the Sunspot Classification Supplemented with Sunspot Area and Its Changes

<http://adsabs.harvard.edu/abs/2012SoPh..281..639L>

We used sunspot data from 1996 to 2010. We noted that sunspot area and its changes can be a proxy of magnetic flux and its emergence/cancellation, respectively.

When the sunspot area increases, the flare occurrence rates and probabilities noticeably increase, especially for major flares. This is statistical evidence that magnetic flux emergence is an important mechanism for triggering solar flares, because sunspot area can be a good proxy of magnetic flux.

Filament: e..g 7 June 2011:

STCE: <http://www.stce.be/news/353/welcome.html>

STCE: <http://www.stce.be/news/x137x/welcome.html>

Science at NASA: https://science.nasa.gov/science-news/science-at-nasa/2011/11jul_darkfireworks

:Issued: 2023 Mar 18 1231 UTC
:Product: documentation at <http://www.sidc.be/products/meu>
#-----#
DAILY BULLETIN ON SOLAR AND GEOMAGNETIC ACTIVITY from the SIDC #
(RWC Belgium) #
#-----#

SIDC URSIGRAM 30318
SIDC SOLAR BULLETIN 18 Mar 2023, 1230UT
SIDC FORECAST (valid from 1230UT, 18 Mar 2023 until 20 Mar 2023)

SOLAR FLARES : C-class flares expected, (probability >=50%)

GEOMAGNETISM : Quiet (A<20 and K<4)

SOLAR PROTONS : Quiet

PREDICTIONS FOR 18 Mar 2023 10CM FLUX: 138 / AP: 007

PREDICTIONS FOR 19 Mar 2023 10CM FLUX: 140 / AP: 019

PREDICTIONS FOR 20 Mar 2023 10CM FLUX: 138 / AP: 029

COMMENT: The solar flaring activity was at moderate levels during the last 24 hours with one M-class flare and several C-class flares detected, with the most frequent sources being NOAA active regions 3254 and 3256. The largest flare was a M1.1 flare, peaking at 15:07 UTC on March 17, associated with active region NOAA 3254 (beta class). NOAA active region 3256 produced an impulsive C9.4 flare at 07:10 UTC on March 18. This event was also associated with Type IV radio emission. Other regions on the disc did not show any significant flaring activity. Further M-class flare activity is possible but not probable, while frequently C-class activity is expected in the next 24 hours.

A filament eruption in the southwestern quadrant was observed on March 17 from around 09:20UTC. The associated CME appears in SoHO/LASCO C2 coronagraph data from 10:23UTC onwards. The CME is directed to the south-west and the bulk of the CME is not expected to be Earth directed. However, a glancing blow of the shock may impact Earth at around 19:00 UTC on March 19. Another small filament eruptions occurred in the northwestern quadrant from around 17:09UTC and 20:09UTC on March 17. We are awaiting corresponding coronagraph data for further analysis. During the last 24 hours there were no other potentially Earth-directed CMEs detected in the available coronagraph observations.

The greater than 10 MeV proton flux was at almost nominal levels over the past 24 hours and is expected to remain so for the next 24 hours. The greater than 2 MeV electron flux remained below the 1000 pfu alert threshold and is expected to remain below this threshold during the next 24 hours. The 24h electron fluence was at normal levels and is expected to remain so.

Over the past 24 hours the solar wind parameters (ACE and DSCOVR) have been indicative of slow solar wind conditions. The solar wind speed ranged between 400 km/s and 450 km/s. The interplanetary magnetic field magnitude was about 6 nT. The magnetic field orientation was predominantly in the positive sector (field directed away from the Sun). Similar slow solar wind regime is expected on March 18 with a slight wind speed enhancement possible for late on March 19, due to expected influence of the small equatorial coronal hole of positive polarity with a chance of being mixed with glancing blow from a CME which left the solar surface around 10 UTC on March 17th.

The geomagnetic conditions over the past 24 hours were globally and locally quiet to unsettled (NOAA Kp and K Bel 1-3). Quiet conditions are expected for March 18 with active to minor storm conditions possible for late on March 19 and March 20, due to expected arrival of the high speed stream and a possible glancing blow from a CME.

TODAY'S ESTIMATED ISN : 044, BASED ON 14 STATIONS.

SOLAR INDICES FOR 17 Mar 2023

WOLF NUMBER CATANIA : 110

10CM SOLAR FLUX : 134

AK CHAMBON LA FORET : 014

AK WINGST : 007

ESTIMATED AP : 007

ESTIMATED ISN : 073, BASED ON 18 STATIONS.

NOTICEABLE EVENTS SUMMARY

DAY BEGIN MAX END LOC XRAY OP 10CM Catania/NOAA RADIO_BURST_TYPES

17 1504 1507 1511 S22W65 M1.0 SN 12/3247

END

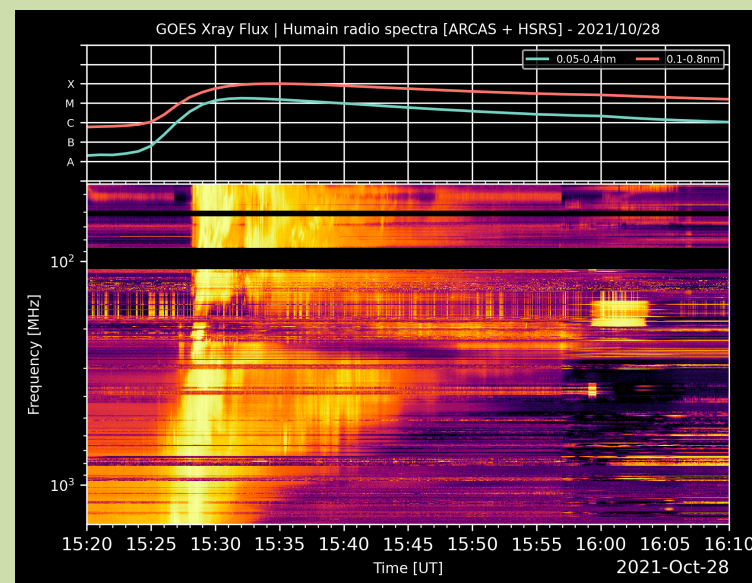


Flare prediction





RADIO BURSTS



Humain



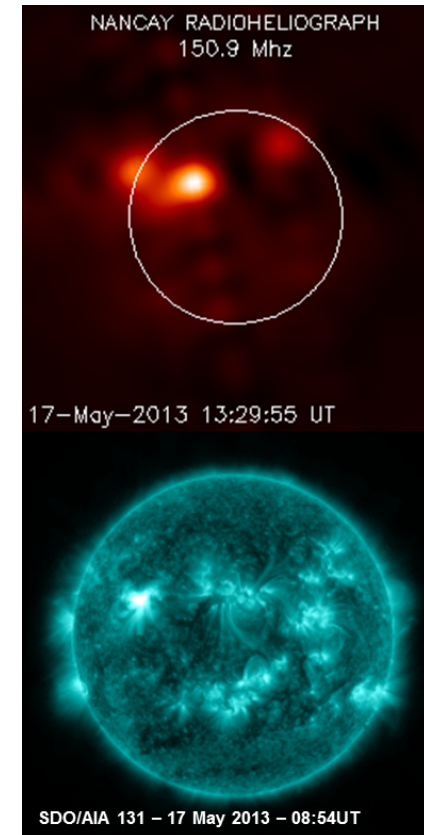
Observations in the e.m. spectrum at radio wavelengths:
• 10 cm flux → index for solar activity, similar like the sunspot index
• Solar Radio Bursts

RADIO BURSTS



Tenflare

- Compared to pre-flare background levels, the 10.7cm (2800 MHz) radio flux suddenly increases by at least 100%
- Example: 17 May 2013
 - M3 flare in NOAA 1748
 - From +/- 140 sfu to > 400 sfu
- May affect daily 10.7cm radio flux values
- Radio flares are also observed at other wavelengths



SDO/AIA



One value of the 10 cm flux value per day —> index for solar activity like the sunspot number

$$1 \text{ sfu} = 10^{-22} \text{ Wm}^{-2} \text{ Hz}^{-1}$$

Tapping (2013): The 10.7 cm solar radio flux (F10.7)

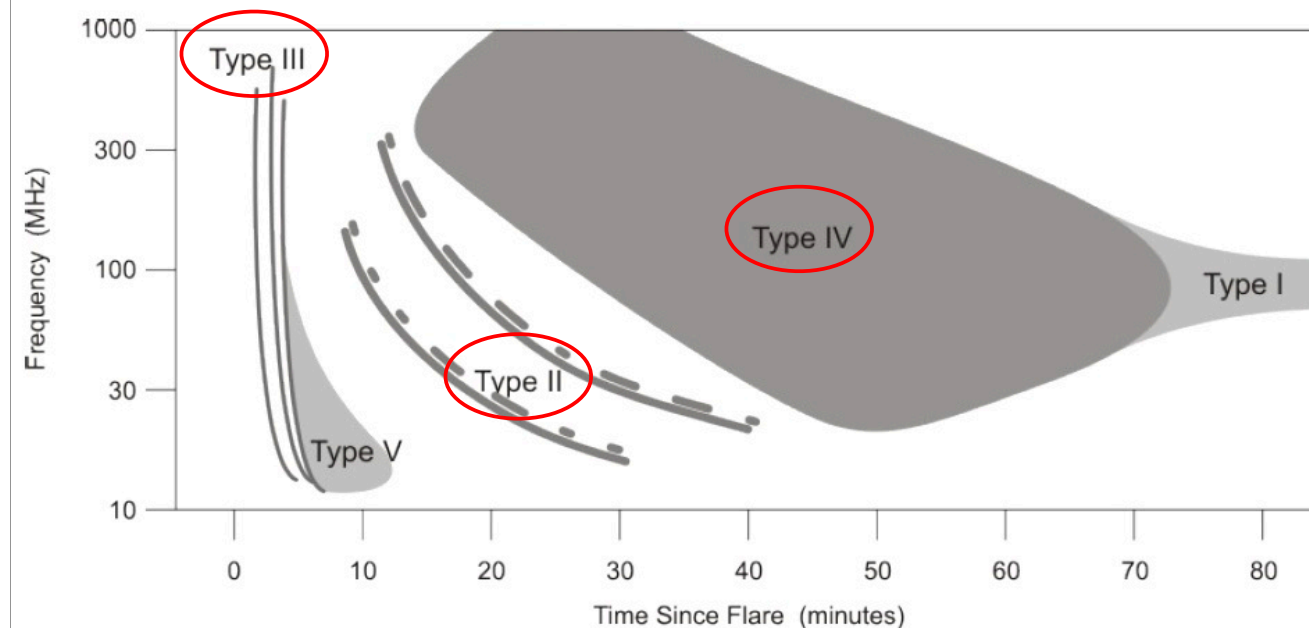
<http://onlinelibrary.wiley.com/doi/10.1002/swe.20064/epdf>

A 10.7 cm solar flux measurement is a determination of the strength of solar radio emission in a 100 MHz-wide band centered on 2800 MHz (a wavelength of 10.7 cm), averaged over an hour. It is expressed in solar flux units (sfu), where $1 \text{ sfu} = 10^{-22} \text{ Wm}^{-2} \text{ Hz}^{-1}$. It is daily measured at Penticton, British Columbia, Canada (DRAO: Dominion Radio Astrophysical Observatory). Measurements are taken at 17UT, 20UT and 23UT (winter period: 18–20–22UT), with the local noon value (20UT) as the value for that day. It is uncorrected for any flare influence. The daily values are at <http://www.spaceweather.ca/solarflux/sx-4a-en.php>

From SWPC Glossary at <https://www.swpc.noaa.gov/content/space-weather-glossary#t>

Tenflare: A solar flare accompanied by a 10cm radio burst of intensity greater than 100% of the pre-burst value.

RADIO BURSTS



Electrons start radiating in radio wavelengths because they get an energy boost from a solar event. A solar radio burst (SRB) is a signature of the presence of a CME or flare, not a direct consequence of magnetic reconnection.

(Note: 10.7cm $\rightarrow$ 2800 MHz: not in this spectrum)

SRBs are produced by electrons energised by solar eruptive events, like flares, coronal mass ejections. Their signature in a spectrogram gives information about the fate of these electrons.

Source of Figure: https://www.sws.bom.gov.au/World_Data_Centre/1/9/4

This diagram illustrates all of the major burst types in a typical configuration following a large flare. It should be noted that it is not common for all of these features to be observed after a flare.

Mind the orientation of the vertical axis! Other figures may have a reversed direction. As the frequency is proportional to the square root of the density, and the density decreases with increasing distance from the Sun, a **decreasing frequency means locations higher up in the solar atmosphere**.

The **ionospheric cut-off frequency** is around 15MHz (due to too low frequency and so **reflected by ionosphere**). In order to observe radio disturbances below this frequency, one has to use **satellites** (above the earth atmosphere) such as STEREO/SWAVES or WIND. Radio bursts at low frequencies (< 15 MHz) are of particular interest because they are associated with energetic CMEs that travel far into the interplanetary (IP) medium and affect Earth's space environment if Earth-directed. Low frequency radio emission needs to be observed from space because of the ionospheric cutoff.

Example: <https://stereo-ssc.nascom.nasa.gov/browse/2017/01/16/insitu.shtml>

Solar Radio Bursts and Space Weather, S.M. White

https://www.nrao.edu/astrores/gbsrbs/Pubs/AJP_07.pdf

Solar radio bursts at frequencies below a few hundred MHz were classified into 5 types in the 1960s (Wild et al., 1963).

Coronal Mass Ejections and solar radio emissions, N. Gopalswamy

<http://citeseerx.ist.psu.edu/viewdoc/download?doi=10.1.1.708.626&rep=rep1&type=pdf>

The three most relevant to space weather radio burst types are type II, III, and IV. Three types of low-frequency non-thermal radio bursts are associated with coronal mass ejections (CMEs): Type III bursts due to accelerated electrons propagating along open magnetic field lines, type II bursts due to electrons accelerated in shocks, and type IV bursts due to electrons trapped in post-eruption arcades behind CMEs.

[Radio burst type II, III, and IV are also the only ones that ever get mentioned in the Ursigrams.]

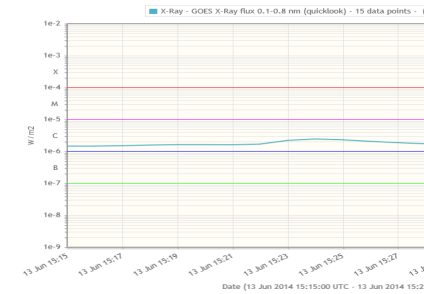
See also: <https://www.stce.be/educational/classification>

RADIO BURSTS

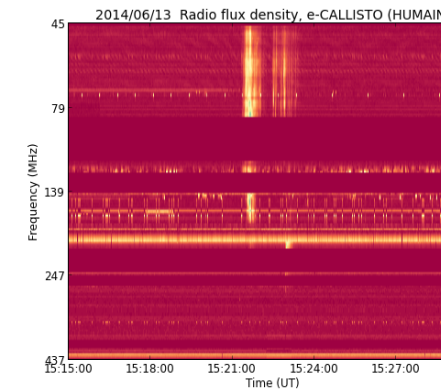


Type III

- Source:
accelerated electrons propagating
along open magnetic field lines
- During impulsive phase of
flares
- Duration
Seconds (isolated) to minutes
(groups)
- Frequency
10 kHz-1 GHz



GOES



HUMAN



Rather than examine each event Type in numerical order, this discussion treats each event Type in the order in which they are most likely to be encountered (i.e. from most to least common).

Image courtesy:

GOES-curve: STAFF viewer, <http://www.staff.oma.be>

Radio plot: ROB/Humain Radio Observatory, <http://www.sidc.be/humain/>

13 June 2014

3940.	1521	1524	1527	G15	5	XRA	1-8A	C2.4	5.2E-04	2087
3940 +	1521	1522	1523	SAG	G	RBR	245	290		2087
3940 +	1521	////	1523	SAG	C	RSP	025-180	III/2		2087
3940 +	1522	1522	1525	HOL	3	FLA	S19E38	SF		2087

Solar Radio Bursts and Space Weather, S.M. White

https://www.nrao.edu/astrores/gbsrbs/Pubs/AJP_07.pdf

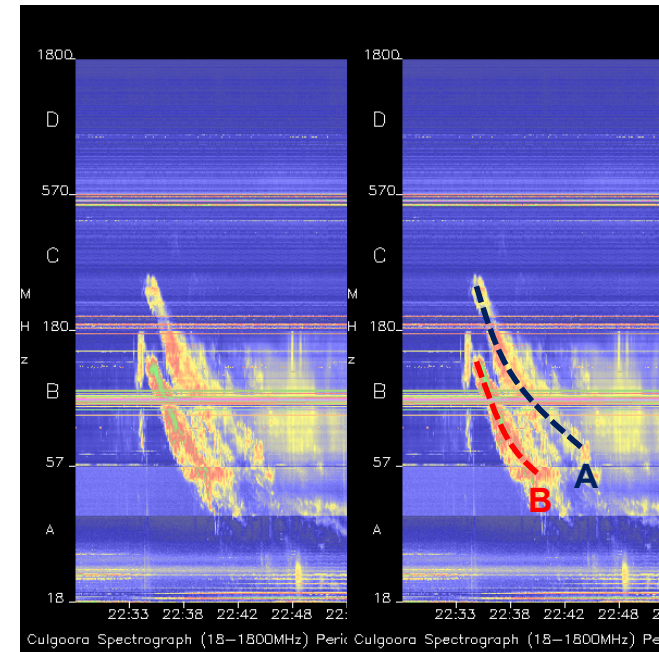
Type III bursts are brief radio bursts that drift very rapidly in frequency versus time (Fig. 1). For example, it can drift from 50 to 20 MHz in about 3 seconds, or 10 MHz s⁻¹. Type IIIs are commonly seen in the impulsive phase of solar flares, and the connection they imply between the acceleration region in solar flares and open field lines that reach the solar wind makes them important for understanding field line connectivity in flares and the access of flare-accelerated particles to the Earth.



RADIO BURSTS

Type II

- Source:
 - electrons accelerated in shocks
 - Indicates CME
 - Shock speed can be derived from fundamental band (B)
- Start at peak in soft X-ray flux of flare
- Duration
 - 3-30 minutes
- Frequency
 - 20-150 MHz



Culgoora



Culgoora spectrograph at 01 Nov 2003 – <http://www.sws.bom.gov.au/Solar/2/2/1> (Type II/2, 1079 km/s)

Solar Radio Bursts and Space Weather, S.M. White

https://www.nrao.edu/astrores/gbsrbs/Pubs/AJP_07.pdf

Type II bursts typically occur at around the time of the soft X-ray peak in a solar flare and are identified by a slow drift to lower frequencies with time in dynamic spectra, the frequent presence of both fundamental and **second-harmonic** bands (with a frequency ratio of 2), and splitting of each of these bands into two traces. The frequency drift rate is typically two orders of magnitude slower than that of the (“fast-drift”) Type III bursts, so the two burst types are readily distinguished.

The harmonic is almost always stronger than the fundamental.

The fundamental band is the one provoked by the shock of the CME and is the one that reaches the lowest frequencies first (track « B » in the image). It is the fundamental track that is used to calculate the (true) speed of the shock as it moves through the corona and away from the Sun (density decrease => frequency decrease).

Roberts (1959): Solar Radio Bursts of Spectral Type II : <http://adsabs.harvard.edu/abs/1959AuJPh..12..327R>

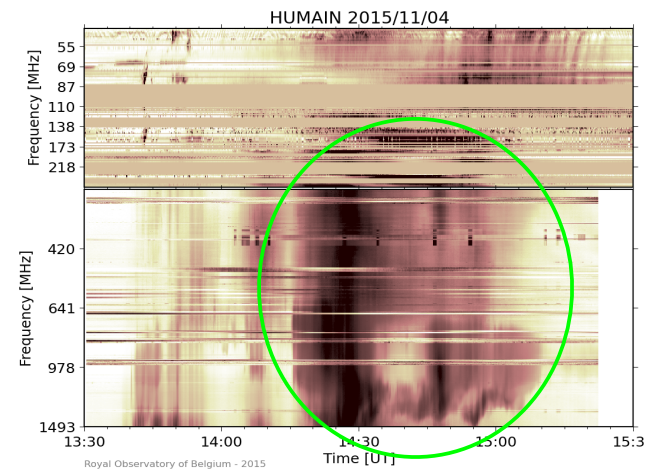
Gopalswamy: Coronal Mass Ejections and solar radio emissions : <http://citeseerx.ist.psu.edu/viewdoc/download?doi=10.1.1.708.626&rep=rep1&type=pdf>

RADIO BURSTS



Type IV

- Source
 - Electrons trapped in post-eruption arcades behind CMEs
 - Related to very energetic CMEs
(average speed: 1200 km/s)
- During decay phase of solar flares
Connection with SEPs
- Duration
Hours (to days)
- Frequency
 - 20 to >1000 MHz
 - Lowest: 8 +/-5 MHz



Gopalswamy: Coronal Mass Ejections and solar radio emissions

<http://citeseerx.ist.psu.edu/viewdoc/download?doi=10.1.1.708.626&rep=rep1&type=pdf>

The type IV bursts are associated with very energetic CMEs (average speed 1200 km/s), confirming the earlier finding by Robinson [1986] for the **continuum** events at metric wavelengths. The radio emission should originate from a heliocentric distance 3.5 to 4.5 Rs, depending on whether the radio emission occurs at the fundamental or harmonic of the plasma frequency. When the type IV burst attains the lowest frequency, the IP type II burst occurs at frequencies well below 1 MHz, which means the shock is much farther away. This suggests that the energetic electrons responsible for the type IV burst might come from the continued reconnection occurring beneath the CME.

[Comment by Dr Christophe Marqué (ROB): The height of type IV reported by Gopalswamy concerns the low frequency ones. The one for example observed in Humain (04 Nov 2015) is really taking place in the post flare loops close to the flare site.]

Solar Radio Bursts and Space Weather, S.M. White

https://www.nrao.edu/astrores/gbsrbs/Pubs/AJP_07.pdf

Type IV bursts are broadband quasi-continuum features associated with the decay phase of solar flares. They are attributed to electrons trapped in closed field lines in the post-flare arcades produced by flares; their presence implies ongoing acceleration somewhere in these arcades, possibly at the tops of the loops in a “helmet-streamer” configuration. Type IV bursts have long been of interest in Space Weather studies because they have a high degree of association with solar energetic particle events.

Example: 04 Nov 2015: <http://www.stce.be/news/326/welcome.html>

2340B1327 U1339 A1348 SVI 2 FLA N09W04 2B ERU 2443

2340 + 1331 1352 1413 G15 5 XRA 1-8A M3.7 5.9E-02 2443

2340 + 1336 1341 1438 SVI G RBR 4995 740 2443

2340 + 1337 1341 1442 SVI G RBR 2695 340 2443

2340 + 1337 1341 1429 SVI G RBR 8800 560 2443

2340 + 1338 1341 1414 SVI G RBR 15400 210 2443

2340 + 1343 //// 1358 SAG C RSP 048-180 II/2 955 2443

2340 + 1351 //// 1531 SVI C RSP 025-171 IV/1 2443

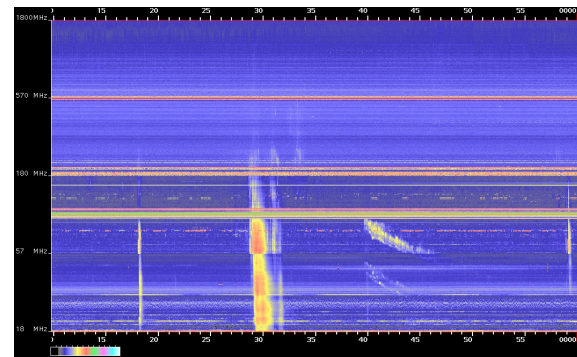
2340 + 1404 1426 1502 SAG G RBR 410 1400 2443

2340 + 1405 1433 1507 SAG G RBR 245 1400 2443

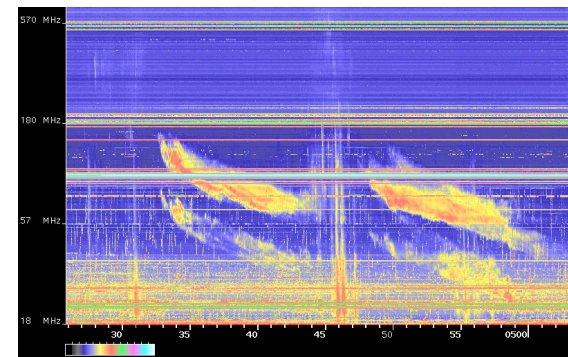
2340 + 1406 1427 1456 SAG G RBR 1415 5800 2443

2340 + 1406 1427 1458 SAG G RBR 610 1000 2443

RADIO BURSTS



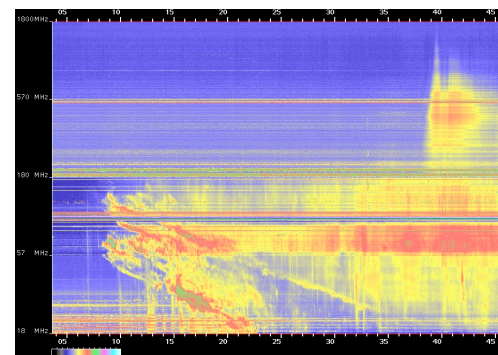
Type II preceded by Type III bursts



Two Type II bursts with background noise



Culgoora



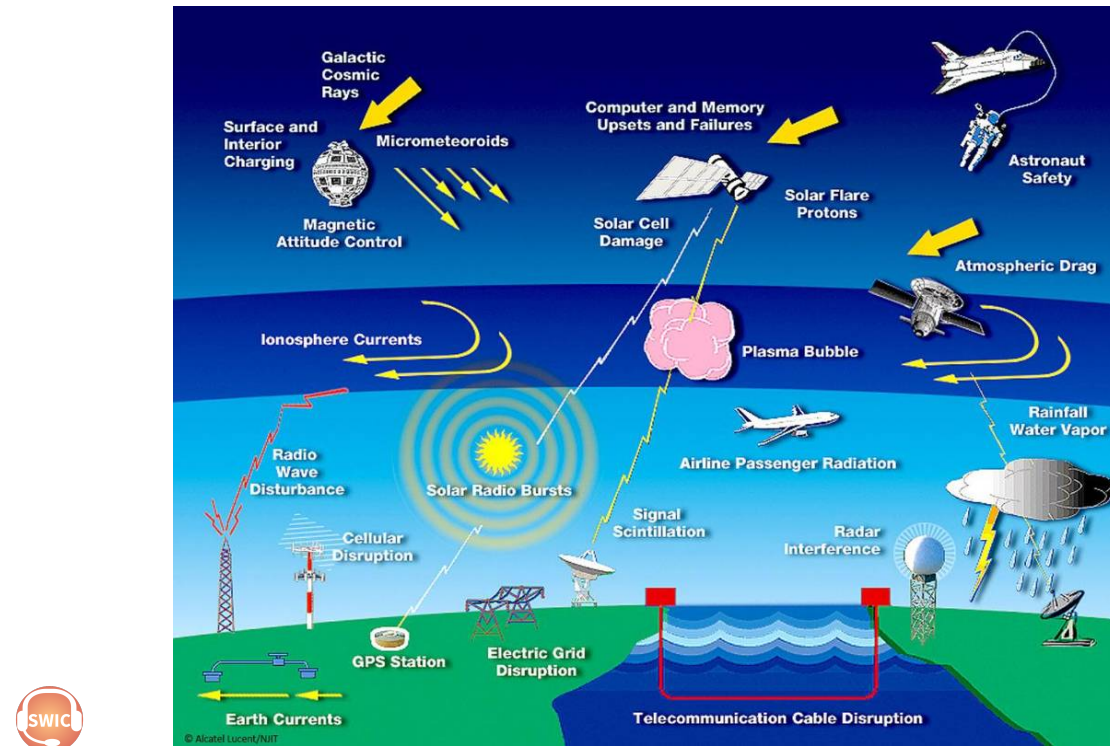
Long duration Type II
with Type IV continuum



Images taken from: https://www.sws.bom.gov.au/World_Data_Centre/1/9/5

Often, for strong bursts, multiple signatures can be seen at the same time.

IMPACT?



Can a Solar Radio Burst impact the ionosphere?

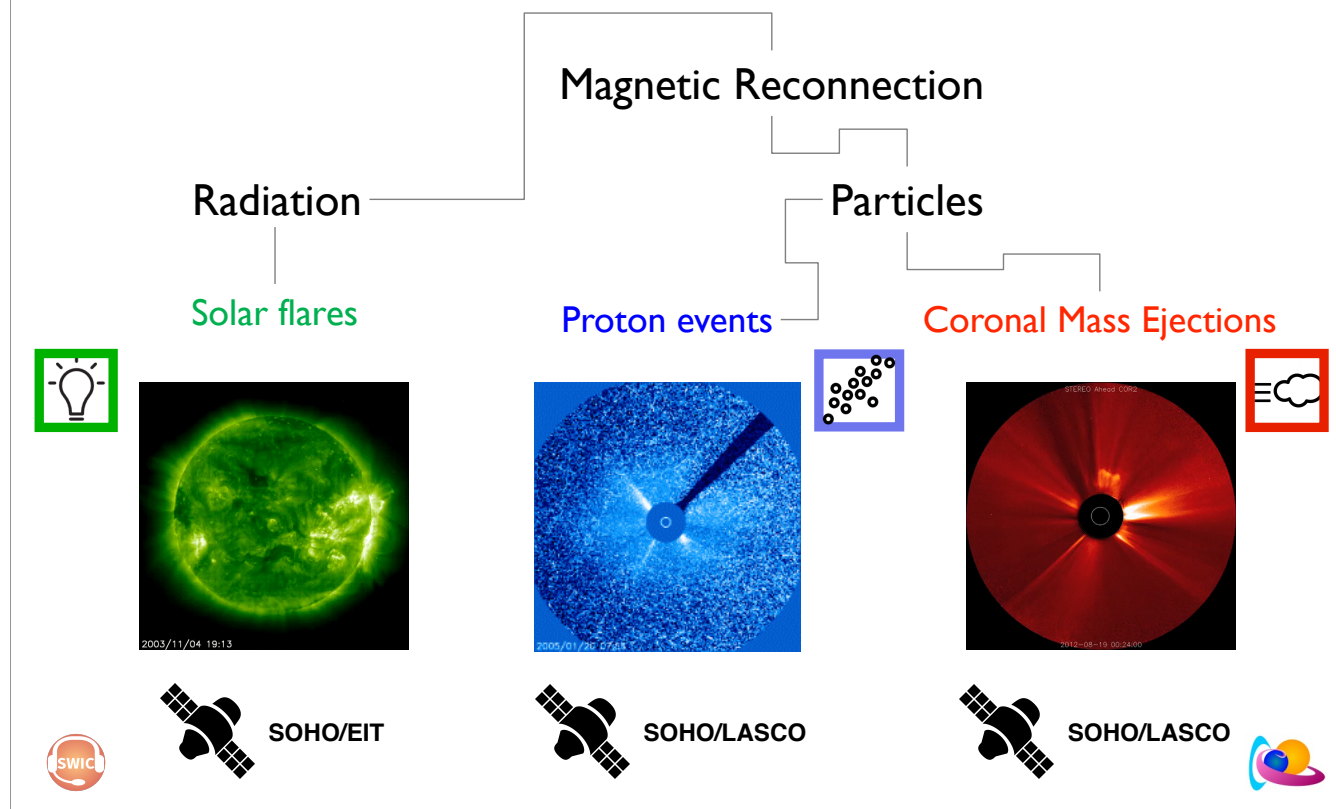
SRB can impact radar systems and GNSS but in a completely other way compared to flares. Flares, i.e. ionising radiation impacting the ionosphere. The radio waves from a SRB behave as a wave used by the GNSS and radar technology.

Noise increase - the ionosphere is not impacted but the signal itself. The noise of the Sun is too loud, the GNSS receiver can't hear the satellite signal clear enough. Or the radar interprets the radio waves coming from the Sun as being a plane.

Image Source: https://www.nasa.gov/mission_pages/sunearth/news/gallery/agu11-spaceweather.html

GPS station Signal/noise - signal is from the satellite. GPS receivers are designed to be sensitive to the signal above them, not at the horizon. When there is a strong radio burst - in the typical GPS frequencies - the noise increases. GPS receiver ontvangt signalen die niet van een satelliet komen maar van de Zon. De GPS ontvanger maakt geen onderscheid tussen solar noise en satelliet signaal. Radar interference Radars are monitoring the planes near the horizon - descending and ascending planes. Radar 'ziet' vliegtuigen door de reflectie van radio-signaal. Radio-signalen van de zon kunnen geïnterpreteerd worden als 'spook'-vliegtuigen: vliegtuigen die je ziet op het radar-scherm maar er in werkelijkheid niet zijn. HF Com f you have a strong radio burst in HF, your MUF might be full of solar noise and in practice not usable

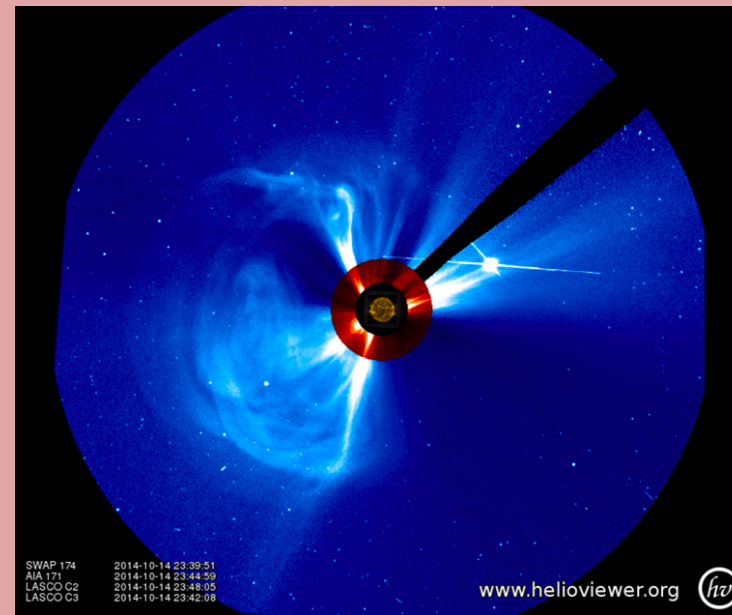
SOLAR ERUPTIONS



CME - OVERVIEW



- Model
- On-disk signatures
 - Filaments
 - Waves
 - Dimming
 - Post-eruption arcade
- Characteristics



PROBA2/SWAP
SDO/AIA
SOHO/LASCO



CORONAL MASS EJECTION

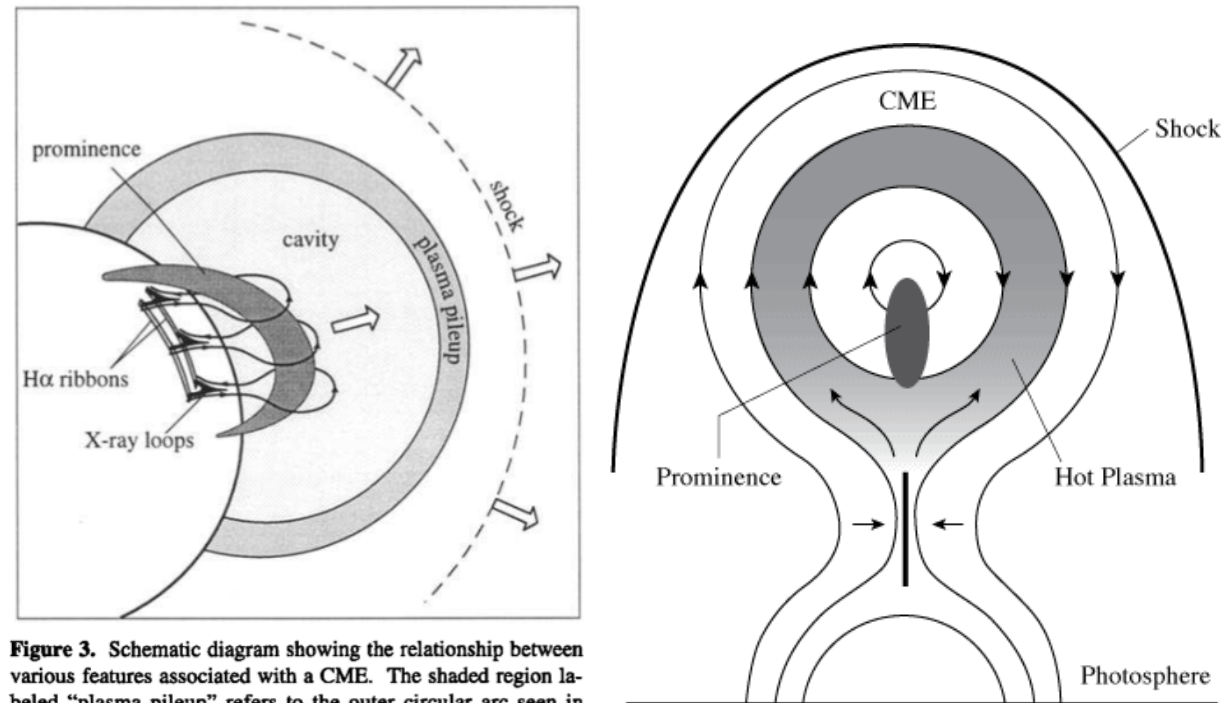


Figure 3. Schematic diagram showing the relationship between various features associated with a CME. The shaded region labeled “plasma pileup” refers to the outer circular arc seen in coronagraphs.



Figure to the right taken from https://ase.tufts.edu/cosmos/print_images.asp?id=27

A magnetic reconnection takes place at a current sheet (dark vertical line) beneath a prominence and above closed magnetic field lines. The coronal mass ejection (CME) traps hot plasma below it (hatched region). The solid curve at the top is the bow shock driven by the CME. The closed field region above the prominence (center) is supposed to become a flux rope in the interplanetary medium.

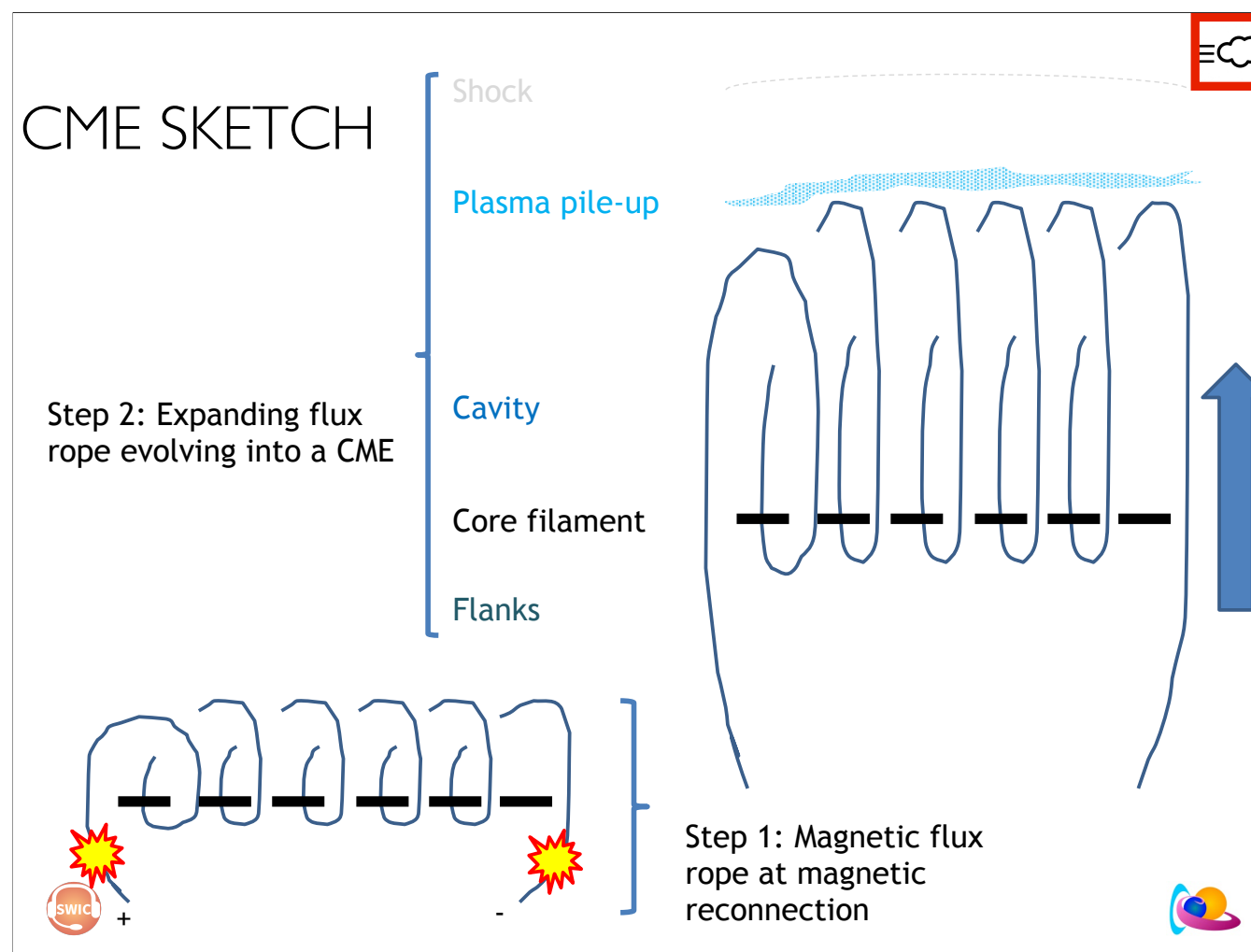
Left: In this model of a three-part coronal mass ejection, portrayed by Terry Forbes (2000), swept-up, compressed mass and a bow shock have been added to the eruptive-flare portrayal of Tadashi Hirayama (1974). The combined representation includes compressed material at the leading edge of a low-density, magnetic bubble or cavity, and dense prominence gas. The prominence and its surrounding cavity rise through the lower corona, followed by sequential magnetic reconnection and the formation of flare ribbons at the footpoints of a loop arcade. [Adapted from Hugh S. Hudson, Jean-Louis Bougeret and Joan Burkepille (2006).]

Figure to the left taken from Forbes (2000): A review on the genesis of coronal mass ejection

<http://adsabs.harvard.edu/abs/2000JGR...10523153F>

<http://onlinelibrary.wiley.com/doi/10.1029/2000JA000005/epdf>

When CMEs were first clearly identified by Skylab in 1973, many researchers assumed that they were caused by the outward expansion of hot plasma produced by a large flare. We now know that this is not the case, for several reasons. First, less than 20% of all CMEs are associated with large flares [Gosling, 1993]. Second, CMEs that are associated with flares often appear to start before the onset of the flare [Wagner et al., 1981; Simnett and Harrison, 1985]. Finally, the thermal pressure produced by a flare is too small to blow open the strong magnetic field of the corona.



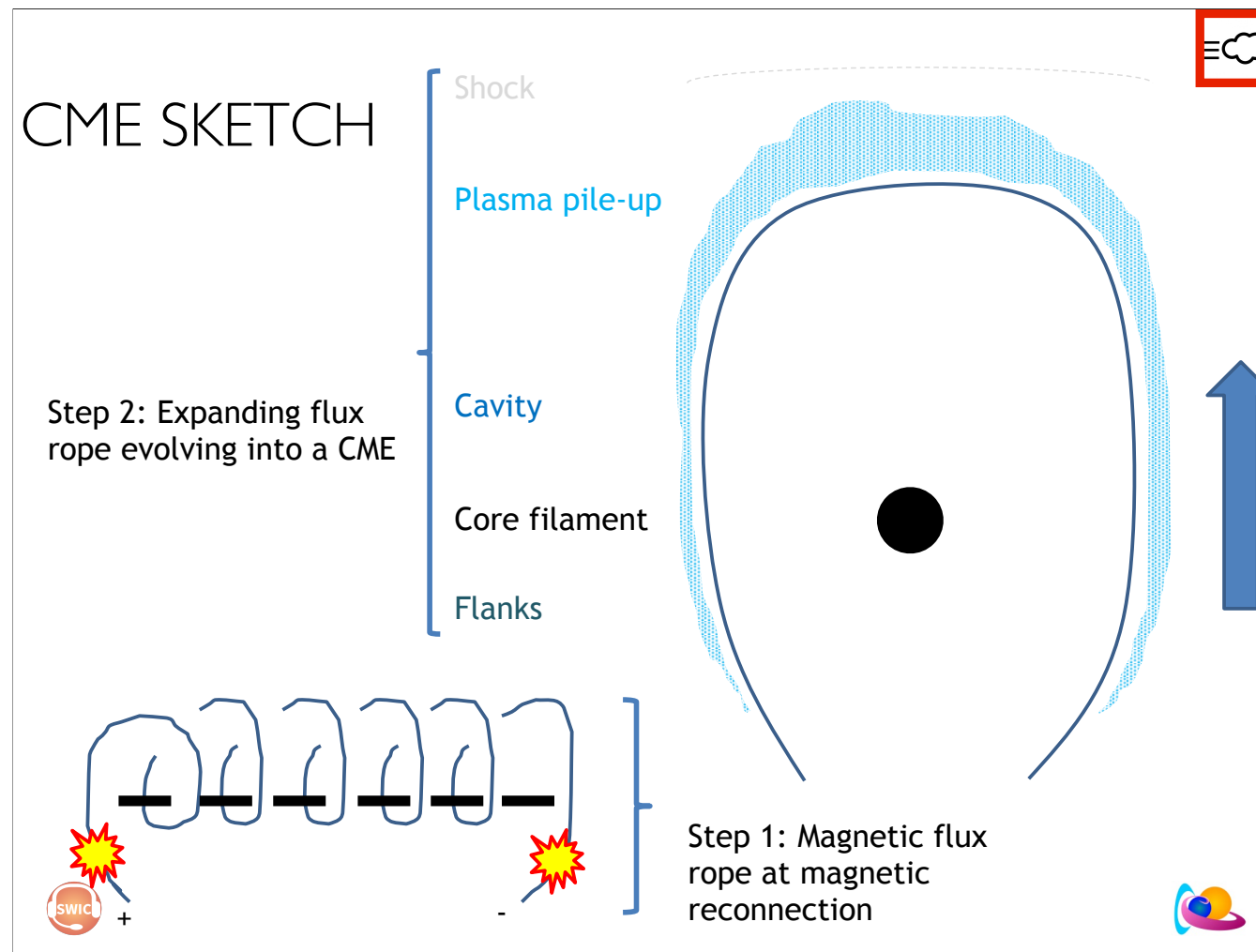
Expanding flux rope consisting of magnetic field lines and filament at bottom (left sketch) evolves into a CME with leading edge (pushed up by top of magnetic field lines from flux rope), cavity and core filament (right sketch). The shock is usually not visible, but measurable by e.g. Type II radio bursts or by observing the deflection of nearby streamers.

Source file: Webb et al. (2012): Coronal Mass Ejections: Observations

<http://link.springer.com/article/10.12942/lrsp-2012-3>

However, with the larger dynamic range of LASCO rims of material detected ahead of fast LASCO CMEs are now considered evidence of shock waves, and emission can be detected ahead of slower speed CMEs as low-level brightness enhancements due to the expanding streamer.

It should be kept in mind that these features are not necessarily present in all CMEs. Not all CMEs contain a prominence, nor do all CMEs have detectable chromospheric ribbons and shock waves

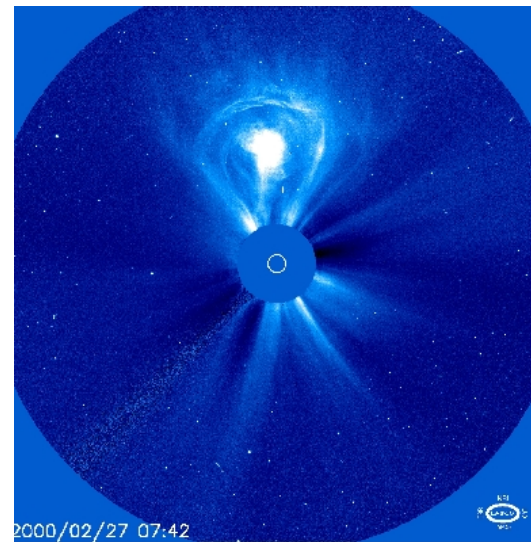


Expanding flux rope consisting of magnetic field lines and filament at bottom evolves into a CME with leading edge (pushed up by top of magnetic field lines from flux rope), cavity and core filament.

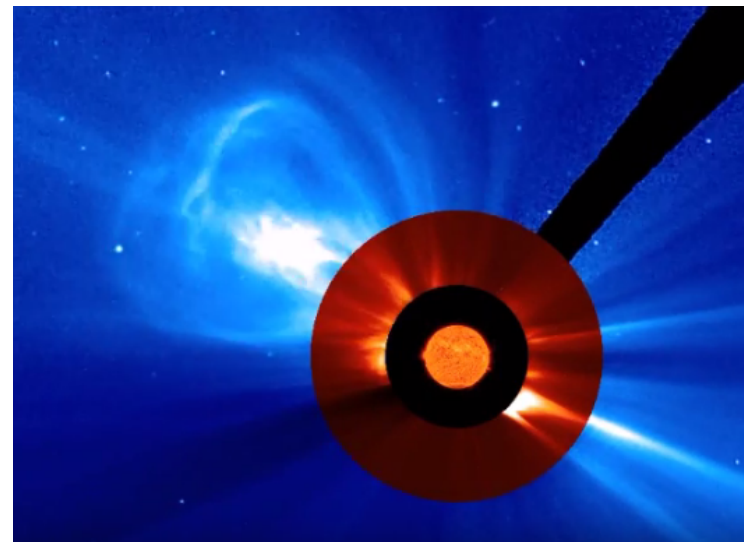
Vourlidas et al. (2013): How Many CMEs Have Flux Ropes? Deciphering the Signatures of Shocks, Flux Ropes, and Prominences in Coronagraph Observations of CMEs – <http://adsabs.harvard.edu/abs/2013SoPh..284..179V>

“... the brightness of the front originates from the pile-up of the overlying streamer material. ... the cavity, while not completely devoid of plasma, does contain less electrons (it is less bright) than its surroundings. ... ‘three-part’-CMEs are indeed systems of ejected FRs, where the cavity is the actual FR. ...”

CORONAL MASS EJECTION



SOHO/LASCO



SDO/AIA
SOHO/LASCO



Left picture: SOHO Gallery: <https://sohowww.nascom.nasa.gov/gallery/images/las02.html>

Right picture: STCE: <http://www.stce.be/news/342/welcome.html>

Coronagraph LASCO: https://lasco-www.nrl.navy.mil/index.php?p=content/handbook/hndbk_5

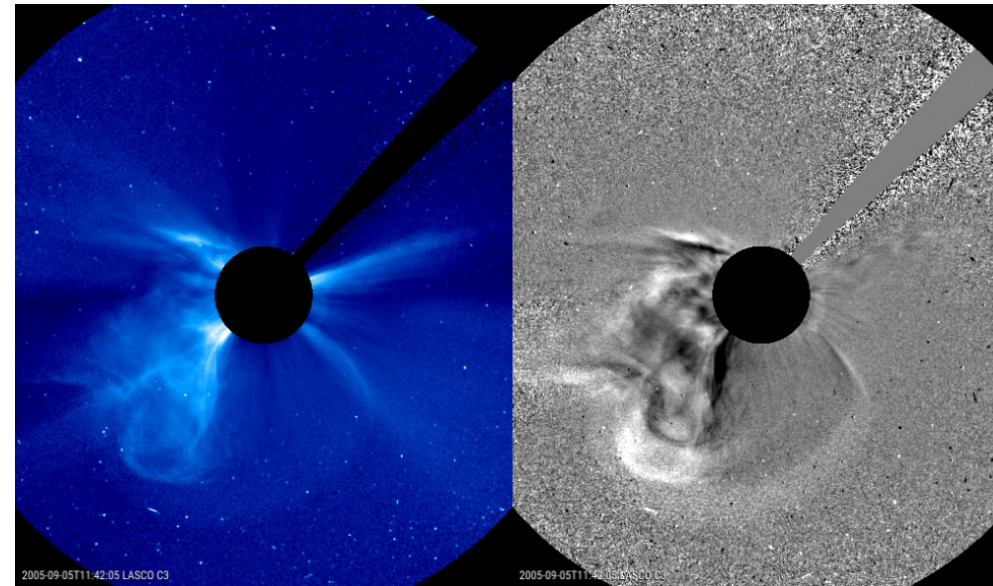
CMEs are mostly observed in white light by coronagraphs from space (SOHO, STEREO).

In order to make the faint CMEs better visible, **difference images** are used (one image subtracted from the other).

Ground-based observatories can observe CMEs very close to the Sun: MLSO (K-Cor):

<http://download.hao.ucar.edu/d5/www/fullres/latest/latest.kcor.gif>

DIFFERENCE IMAGES



SOHO/LASCO



Left picture: SOHO Gallery: <https://sohowww.nascom.nasa.gov/gallery/images/las02.html>

Right picture: STCE: <http://www.stce.be/news/342/welcome.html>

Coronagraph LASCO: https://lasco-www.nrl.navy.mil/index.php?p=content/handbook/hndbk_5

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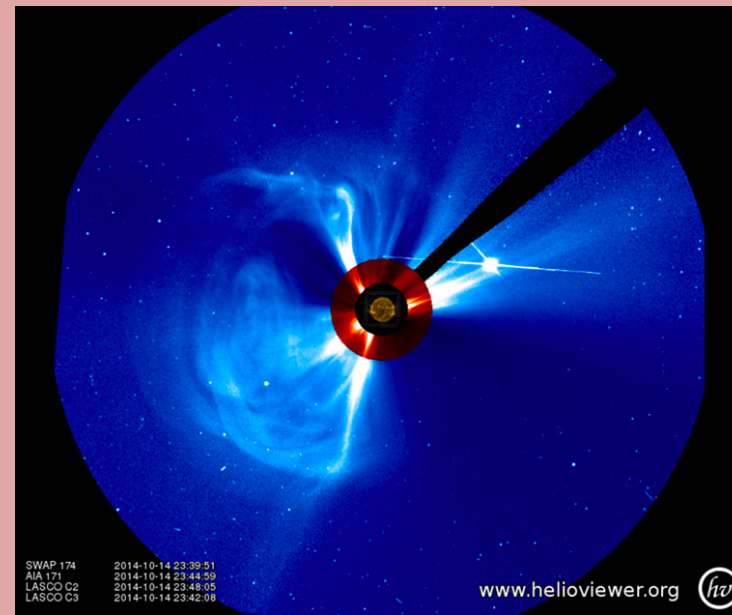
Ground-based observatories can observe CMEs very close to the Sun: MLSO (K-Cor):

<http://download.hao.ucar.edu/d5/www/fullres/latest/latest.kcor.gif>

CME - OVERVIEW



- Model
- On-disk signatures
 - Filaments
 - Waves
 - Dimming
 - Post-eruption arcade
- Characteristics



PROBA2/SWAP
SDO/AIA
SOHO/LASCO



FILAMENTS & PROMINENCES

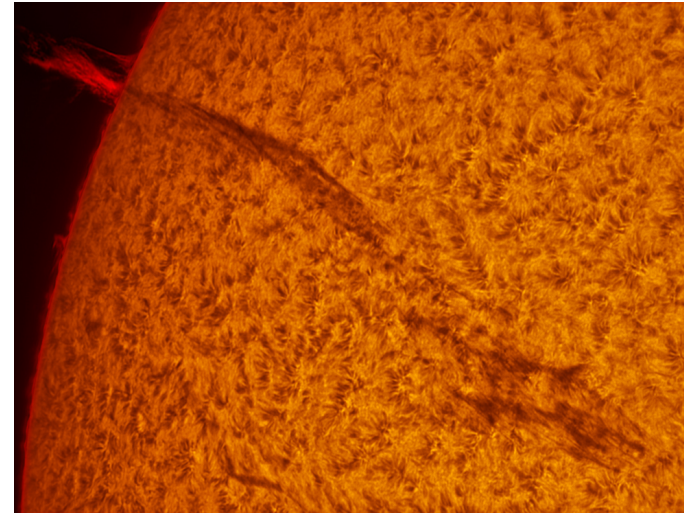
Chromospheric features

- Protruding into corona
- H-alpha and EUV
- Relatively cool (10.000K)

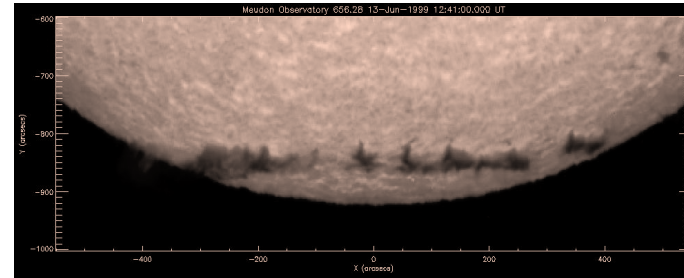
Mark the transition between positive and negative magnetic areas
Appear all over solar disk

Location

- In active region
Group filament
- Outside active region
Quiescent filament
- Polar crown filament



Large NE Solar Filament (Filament and Prominence) by Jett Aguilar, 01:45 UTC, April 28, 2015, Quezon City, Philippines
Lunt 100 mm/DF1800 H-Alpha telescope, 2.5x Televue Powermate, DMK31AU03 camera



Prominences are observed at the limb, filaments on the solar disk, but these are in fact the same features.

More info on filaments at <http://www.stce.be/news/219/welcome.html>

Polar crown filaments form above the polarity inversion line between the old magnetic flux of the previous cycle and the new magnetic flux of the current cycle. Studying their appearance and their properties can lead to a better understanding of the solar cycle.

More info on polar crown filaments at:

- <http://solar.physics.montana.edu/wood/99Prom.html>
- https://science.nasa.gov/science-news/science-at-nasa/2008/17sep_polarcrown

FILAMENTS & PROMINENCES

Quiescent

- Not associated to AR
- Weeks to months

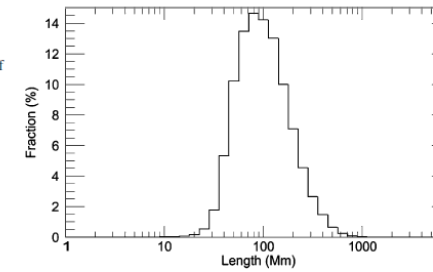
Length

- Long filaments ($> 15^\circ$)
- More prone to instabilities
 - Near (emerging) active regions
 - From chromospheric or coronal waves
 - Near coronal holes
- Flaring
 - 38% result in at least 1 flare
 - Intensity : C1-M1

1120

A.G. Tlatov et al.

Figure 4 Distribution of the filament length in a logarithmic scale. The relative number of filaments is given as a function of their length in Mm.



R. Mawad et al. / *Advances in Space Research* 55 (2015) 696–704

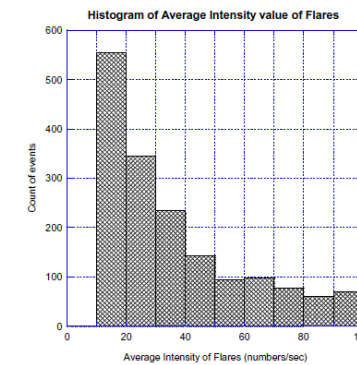


Fig. 6. Histogram of intensity of solar flares which are occurred during duration time of filament disappearance event during 1996–2010.



Tlatov et al. (2016): Tilt Angles of Solar Filaments over the Period of 1919 – 2014

<http://adsabs.harvard.edu/abs/2016SoPh..291.1115T>

Mawad et al. (2014): Filaments disappearances in relation to solar flares during the solar cycle 23

<http://adsabs.harvard.edu/abs/2015AdSpR..55..696M>

Hao et al. (2015): Statistical Analysis of Filament Features Based on the H α Solar Images from 1988 to 2013 by Computer Automated Detection Method

<http://adsabs.harvard.edu/abs/2015ApJS..221...33H>

FILAMENT/PROMINENCE ERUPTIONS

Height

- Zirin (1988)
 - If $> 50,000$ km, eruption likely within next 48 hours
- Filippov (2008)
 - Height cannot exceed critical height
 - Related to strength of and change in magnetic field

Other signs

- Darkening filament
- Change in tilt or length

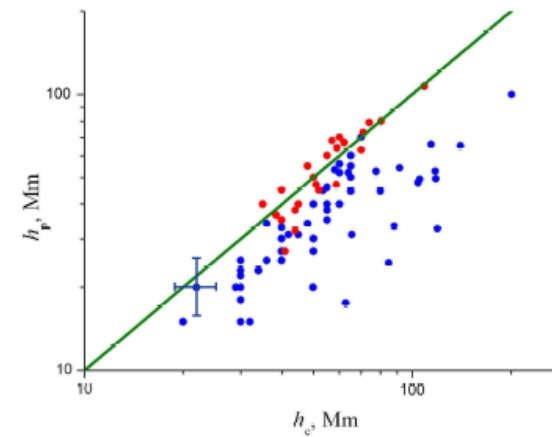


Fig. 7. The observed filament height above the chromosphere h_p versus the critical height of stable filament equilibrium h_c . The blue circles correspond to the filaments which safely passed the west limb. The red circles correspond to the filaments which disappeared from the disk. The straight green line corresponding to an equality of these quantities is the stability boundary.



Data taken from:

Filippov et al. (2008): Causal relationships between eruptive prominences and coronal mass ejections

<http://adsabs.harvard.edu/abs/2008AnGeo..26.3025F>

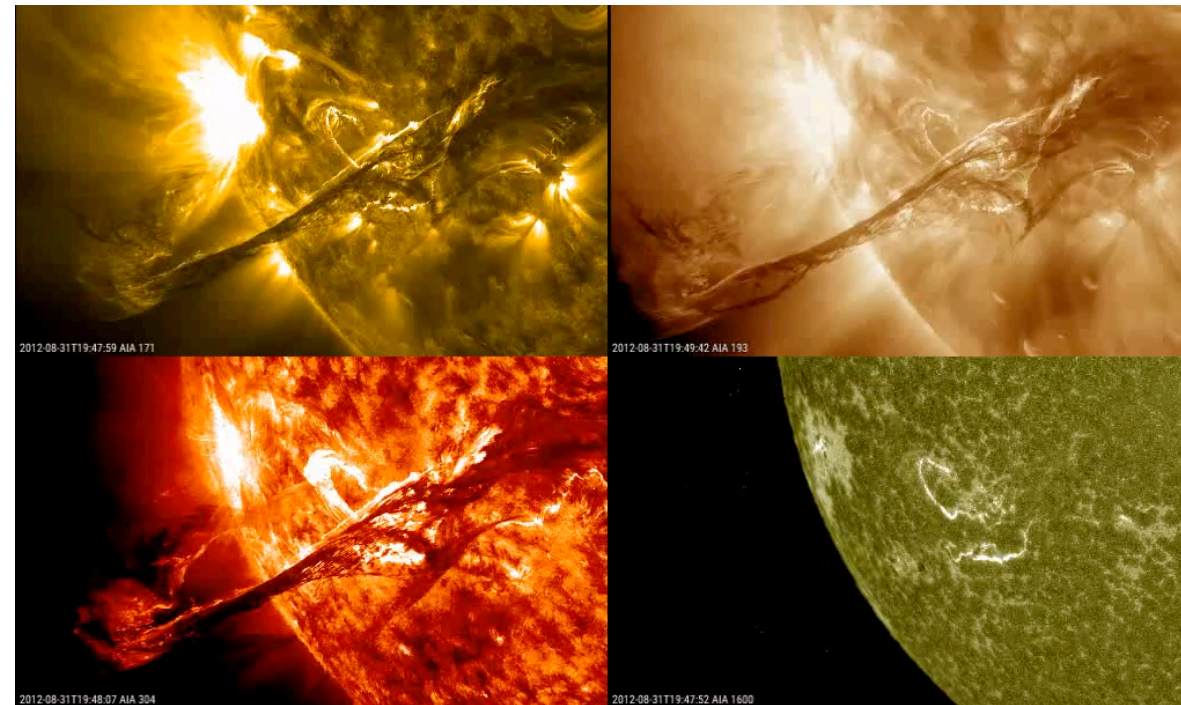
<http://www.ann-geophys.net/26/3025/2008/angeo-26-3025-2008.pdf>

Zirin (1988): Astrophysics of the Sun

The critical height is a model based parameter depending on the strength of the magnetic field and describing how the equilibrium of a prominence will be unstable when it reaches a certain height (See Filippov & Den (2000)).

Most disrupted prominences become unstable at a height of $0.06\text{--}0.14 R_\odot$ from the solar surface, and there are two most probable critical heights at which a prominence is very likely to become unstable, the first one is $0.13 R_\odot$ and the second one is $0.19 R_\odot$. (<https://iopscience.iop.org/article/10.1088/0004-637X/744/2/168>)

FILAMENT ERUPTION



SDO/AIA



This movie shows a spectacular filament eruption, where a very long-living filament (that was already seen at the previous solar rotation) came back around the solar limb. Only, now the filament was not stable anymore and erupted causing a coronal mass ejection. Due to the location of the filament on the east limb, only a glancing blow (flank) of this relatively slow moving particle cloud (± 520 km/s) struck Earth, sparking a minor geomagnetic storm. Aurorae were visible in Scandinavia, Scotland, Canada and Alaska.

In the lower right panel, the **flare ribbons** are clearly seen to move apart.

Sources:

<http://www.stce.be/news/157/welcome.html>

<http://www.stce.be/news/218/welcome.html>

:Issued: 2023 Mar 18 1231 UTC
:Product: documentation at <http://www.sidc.be/products/meu>
#-----#
DAILY BULLETIN ON SOLAR AND GEOMAGNETIC ACTIVITY from the SIDC #
(RWC Belgium) #
#-----#

SIDC URSIGRAM 30318
SIDC SOLAR BULLETIN 18 Mar 2023, 1230UT
SIDC FORECAST (valid from 1230UT, 18 Mar 2023 until 20 Mar 2023)
SOLAR FLARES : C-class flares expected, (probability >=50%)
GEOMAGNETISM : Quiet (A<20 and K<4)
SOLAR PROTONS : Quiet
PREDICTIONS FOR 18 Mar 2023 10CM FLUX: 138 / AP: 007
PREDICTIONS FOR 19 Mar 2023 10CM FLUX: 140 / AP: 019
PREDICTIONS FOR 20 Mar 2023 10CM FLUX: 138 / AP: 029

COMMENT: The solar flaring activity was at moderate levels during the last 24 hours with one M-class flare and several C-class flares detected, with the most frequent sources being NOAA active regions 3254 and 3256. The largest flare was a M1.1 flare, peaking at 15:07 UTC on March 17, associated with active region NOAA 3254 (beta class). NOAA active region 3256 produced an impulsive C9.4 flare at 07:10 UTC on March 18. This event was also associated with Type IV radio emission. Other regions on the disc did not show any significant flaring activity. Further M-class flare activity is possible but not probable, while frequently C-class activity is expected in the next 24 hours.

A filament eruption in the southwestern quadrant was observed on March 17 from around 09:20UTC. The associated CME appears in SoHO/LASCO C2 coronagraph data from 10:23UTC onwards. The CME is directed to the south-west and the bulk of the CME is not expected to be Earth directed. However, a glancing blow of the shock may impact Earth at around 19:00 UTC on March 19. Another small filament eruptions occurred in the northwestern quadrant from around 17:09UTC and 20:09UTC on March 17. We are awaiting corresponding coronagraph data for further analysis. During the last 24 hours there were no other potentially Earth-directed CMEs detected in the available coronagraph observations.

The greater than 10 MeV proton flux was at almost nominal levels over the past 24 hours and is expected to remain so for the next 24 hours. The greater than 2 MeV electron flux remained below the 1000 pfu alert threshold and is expected to remain below this threshold during the next 24 hours. The 24h electron fluence was at normal levels and is expected to remain so.

Over the past 24 hours the solar wind parameters (ACE and DSCOVR) have been indicative of slow solar wind conditions. The solar wind speed ranged between 400 km/s and 450 km/s. The interplanetary magnetic field magnitude was about 6 nT. The magnetic field orientation was predominantly in the positive sector (field directed away from the Sun). Similar slow solar wind regime is expected on March 18 with a slight wind speed enhancement possible for late on March 19, due to expected influence of the small equatorial coronal hole of positive polarity with a chance of being mixed with glancing blow from a CME which left the solar surface around 10 UTC on March 17th.

The geomagnetic conditions over the past 24 hours were globally and locally quiet to unsettled (NOAA Kp and K Bel 1-3). Quiet conditions are expected for March 18 with active to minor storm conditions possible for late on March 19 and March 20, due to expected arrival of the high speed stream and a possible glancing blow from a CME.

TODAY'S ESTIMATED ISN : 044, BASED ON 14 STATIONS.

SOLAR INDICES FOR 17 Mar 2023
WOLF NUMBER CATANIA : 110
10CM SOLAR FLUX : 134
AK CHAMBON LA FORET : 014
AK WINGST : 007
ESTIMATED AP : 007
ESTIMATED ISN : 073, BASED ON 18 STATIONS.

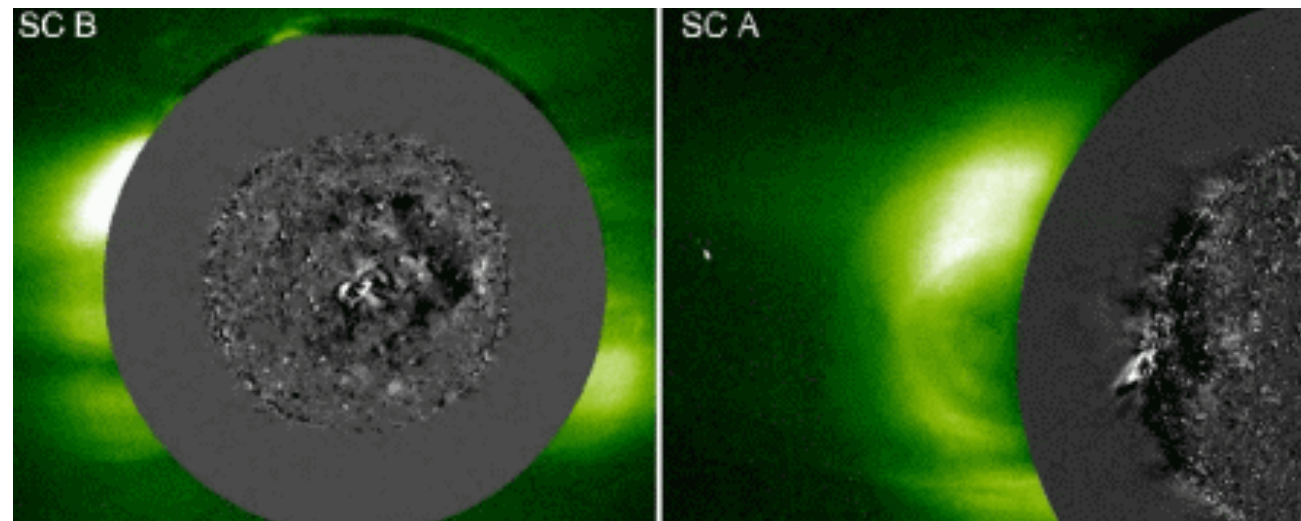
NOTICEABLE EVENTS SUMMARY
DAY BEGIN MAX END LOC XRAY OP 10CM Catania/NOAA RADIO_BURST_TYPES
17 1504 1507 1511 S22W65 M1.0 SN 12/3247
END



filaments / prominences



EIT WAVE



STEREO/EUVI & COR1



There are expanding large-scale waves in the solar atmosphere usually associated to strong solar eruptions (flares and CMEs). We distinguish two kinds of waves, Moreton waves are observed in H α and mostly related to flares, while EIT waves are seen in EUV and mostly related to CMEs.

Another name for this kind of feature is "solar tsunami".

EIT waves are named after the SOHO/EIT instrument with which they were first discovered. They are coronal waves with rather low speeds (200–600 km/s) and are typically related to CMEs. Since all EUV imagers have observed these events, we now call them EUV waves as well.

Example above on EIT wave from: <https://cor1.gsfc.nasa.gov/>

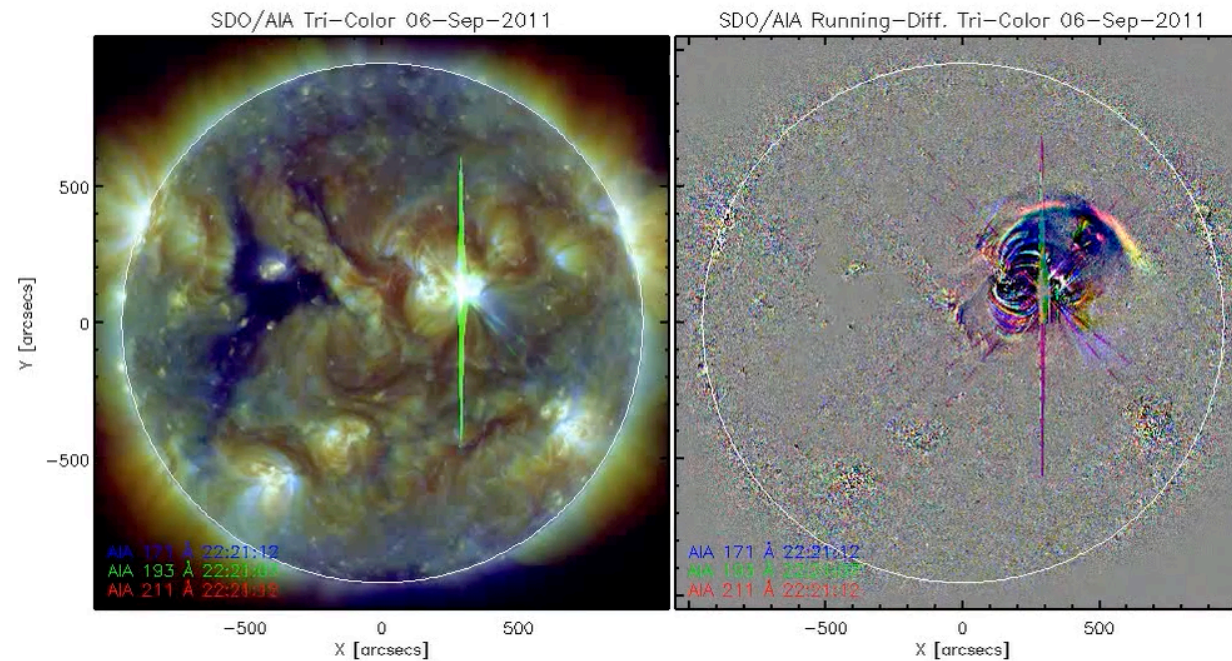
The twin STEREO spacecraft confirmed the existence of EIT waves in February 2009 when sunspot 11012 unexpectedly erupted. The blast hurled a billion-ton cloud of gas (a "CME") into space and sent a tsunami racing along the sun's surface. STEREO recorded the wave from two positions separated by 90 degrees,

More examples:

Moreton waves: <http://www.stce.be/news/222/welcome.html>

EIT waves: <http://www.stce.be/news/222/welcome.html> and <http://www.stce.be/news/241/welcome.html>

EUV WAVE



SDO/AIA



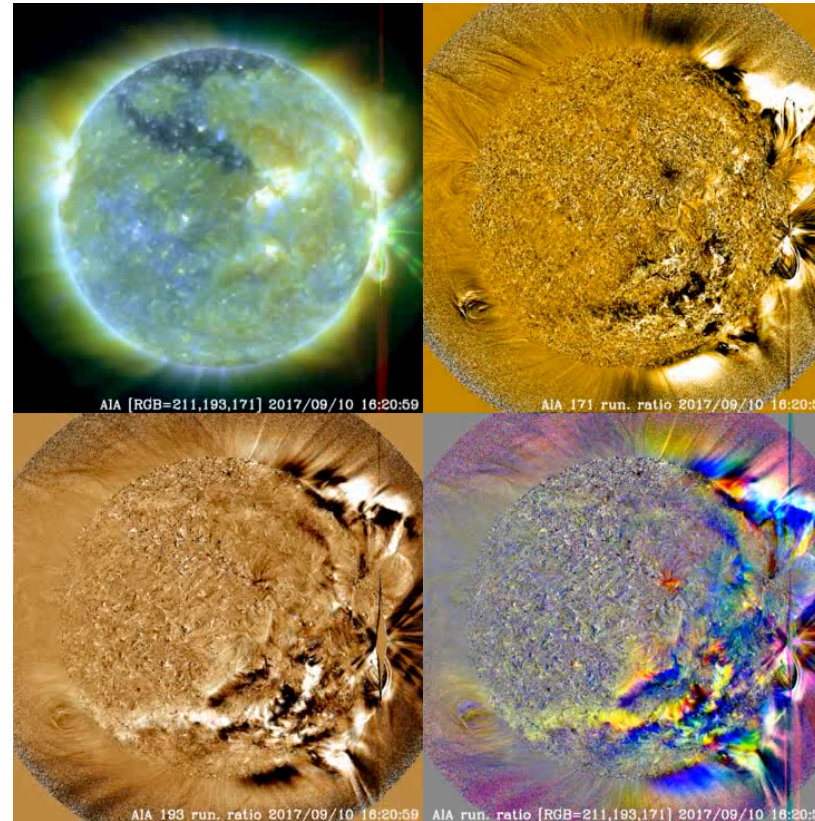
Shen et al. (2014): A Chain of Winking (Oscillating) Filaments Triggered by an Invisible Extreme-ultraviolet Wave

<http://adsabs.harvard.edu/abs/2014ApJ...786..151S>

In this paper, we present an interesting observational study of a chain of winking filaments that was in association with a GOES X2.1 flare in the NOAA active region AR11283 (N13W18) on 2011 September 6. The flare was produced with a remarkable EUV wave propagating mainly in the northwest direction, which not only triggered the oscillation of three filaments in the northwest of AR11283, but also launched the oscillation of a long filament and the occurrence of a small jet in the eastern hemisphere, where the wave signature is very weak or even invisible.

Based on our analysis results, we conclude that the EUV wave is a **good agent for triggering and connecting successive but separated solar activities** in the solar atmosphere.

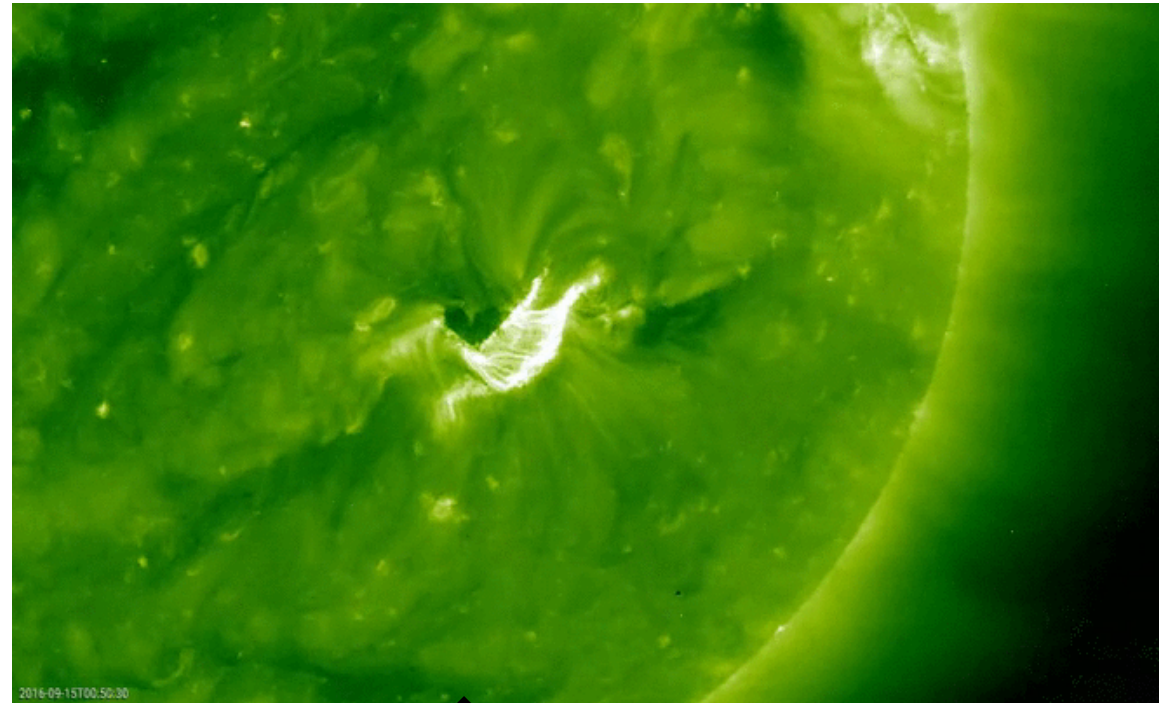
GLOBAL EUV WAVE



EUV-waves can be **stopped**, **reflected** and **refracted** at the boundary of coronal holes or near active regions.

Liu et al. (2018): A Truly Global Extreme Ultraviolet Wave from the SOL2017-09-10 X8.2+ Solar Flare-Coronal Mass Ejection
<http://adsabs.harvard.edu/abs/2018ApJ...864L..24L>
from X8 flare

CORONAL DIMMING (~~TRANSIENT CORONAL HOLE~~)



STEREO/EUVI



A coronal dimming occurs in the wake of coronal mass ejection and is usually interpreted as mass depletions due to the loss or rapid expansion of the overlying corona (the plasma is hurled away in the coronal mass ejection). This darkening is temporary as the plasma is supplemented again. That is why they are also called transient coronal holes. However, this is a confusing name as a coronal hole is a completely different type of structure.

Sources:

<http://www.stce.be/news/362/welcome.html>

Mason et al. (2016): Relationship of EUV Irradiance Coronal Dimming Slope and Depth to Coronal Mass Ejection Speed and Mass

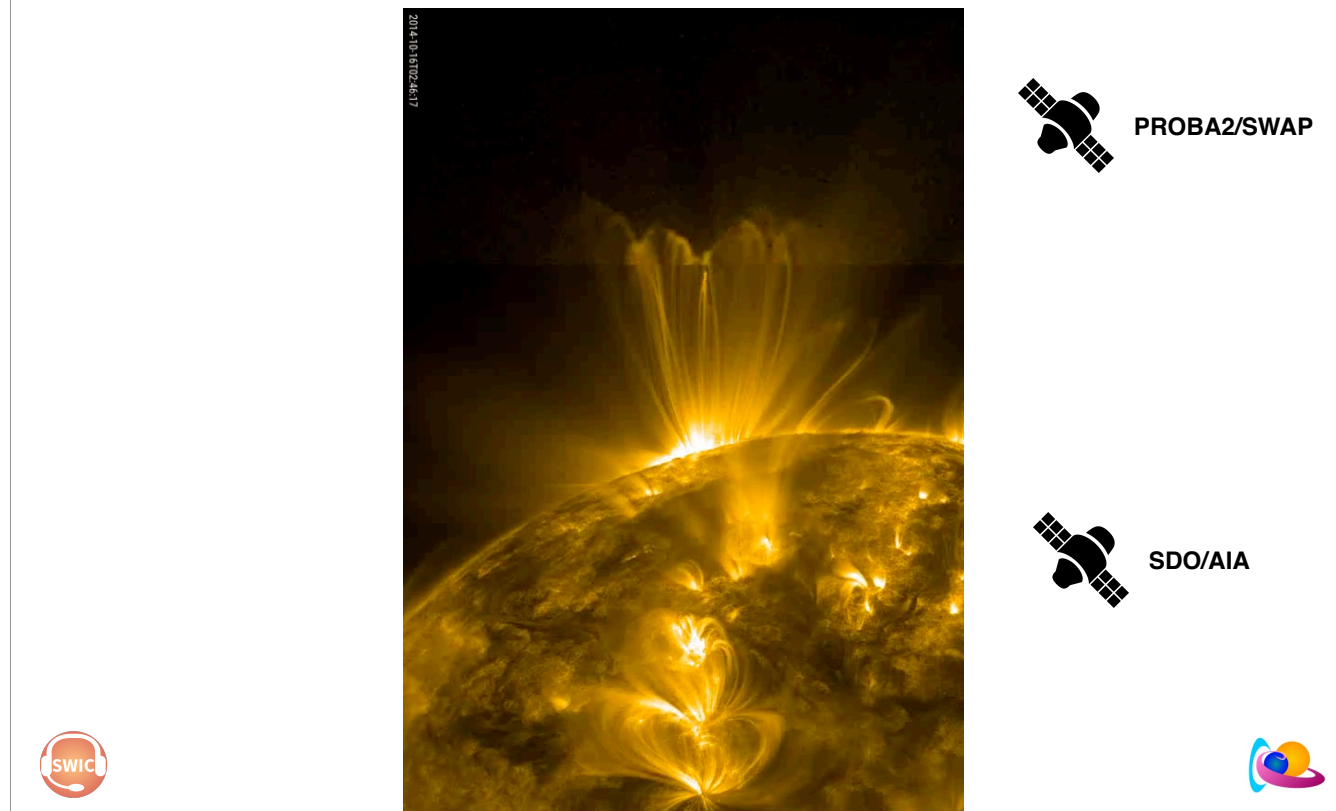
<http://adsabs.harvard.edu/abs/2016ApJ...830...20M>

<http://www.stce.be/news/362/welcome.html>

Cheng et al. (2016): The Nature of CME-flare-Associated Coronal Dimming

<http://adsabs.harvard.edu/abs/2016ApJ...825...37C>

POST-ERUPTION CORONAL LOOPS



Post-flare loops indicate a long duration event and thus the presence of a CME.

This M2-event finished with an **arcade**, which is the technical term for a series of post-flare coronal loops. Interestingly, these post-flare loops continued to grow, first reaching the limit of AIA's Field-Of-View (FOV), and then continuing to grow even beyond AIA's FOV. Fortunately, PROBA2's wider-field SWAP telescope was able to monitor this arcade in its full glory till its disappearance.

The loops of this long duration arcade were **visible for about 2.5 days** (60 hours!), and at their maximum height, they were towering at least 340.000 km above the solar surface. That's not far from the average Earth-Moon distance!

Sources:

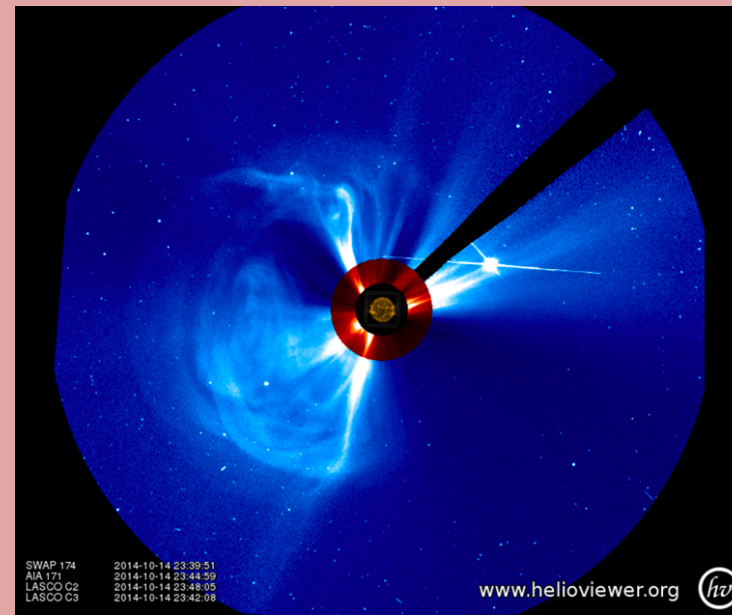
STCE: <http://www.stce.be/news/316/welcome.html>

STCE: <http://www.stce.be/news/331/welcome.html>

STCE: <http://www.stce.be/news/274/welcome.html>

CME - OVERVIEW

- Model
- On-disk signatures
 - Filaments
 - Waves
 - Dimming
 - Post-eruption arcade
- Characteristics



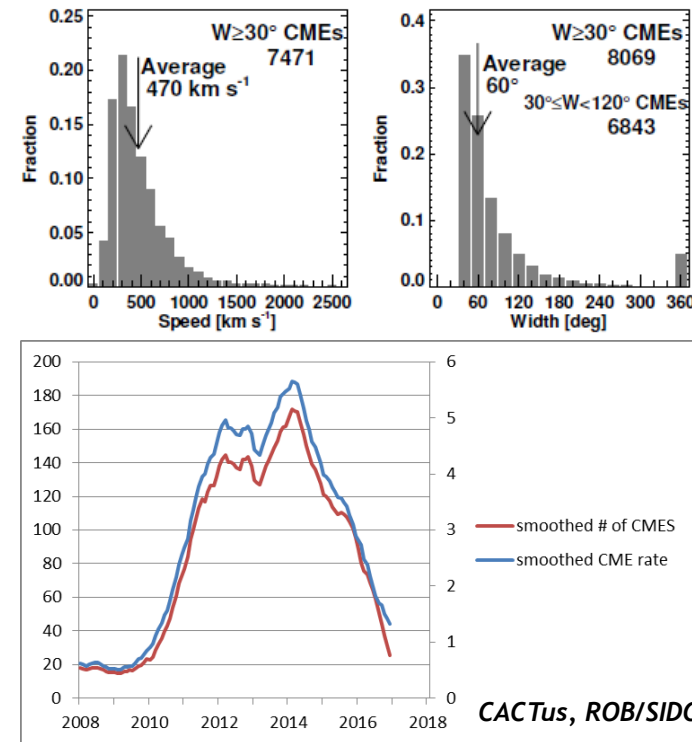
PROBA2/SWAP
SDO/AIA
SOHO/LASCO



CORONAL MASS EJECTION

Characteristics

- Average speed
 - 470 km/s
 - Fast transit event (<24h)
- Average width
 - 60 degrees
- Average Mass
 - 10^{12} kg
 - ~ medium sized mountain
- Number per day
 - 1-6 / day



Source file: Webb et al. (2012): Coronal Mass Ejections: Observations
<http://link.springer.com/article/10.12942/lrsp-2012-3>

A nice example a fairly recent (2012) Fast Transit Event can be found in the STCE News item: A CME with an Olympic Speed
<http://www.stce.be/news/152/welcome.html>

This CME had a transit time of about 19 hours, but was directed towards ST-A, not Earth.

It is believed that, if the CME had been earth-directed, the space weather consequences would have been similar to the Carrington event.

CORONAL MASS EJECTION

Terminology

- Width
 - Narrow: $< 20^\circ$
 - Partial halo: $> 120^\circ$
 - (Full) halo: 360°
- Shape halo
 - Symmetric
 - Asymmetric
- Origin
 - Frontside/Farside
- De- & acceleration

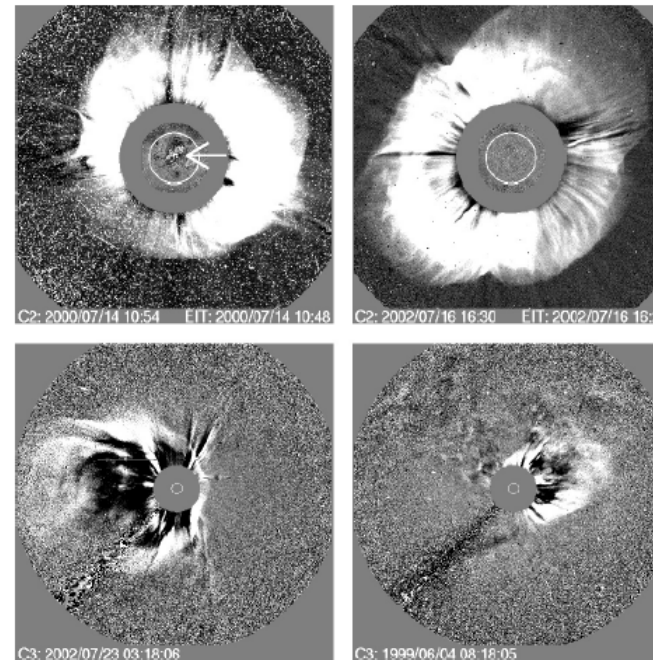


Figure 4: Examples of a variety of halo CME observations, clockwise: a frontside full halo (arrow shows likely source near Sun center); a backside full halo; a partial halo; and an asymmetric full halo. Image reproduced with permission from Gopalswamy *et al.* (2003a).



Source file: Webb et al. (2012): Coronal Mass Ejections: Observations
<http://link.springer.com/article/10.12942/lrsp-2012-3>

Some CMEs appear as narrow jets, some arise from pre-existing coronal streamers (the so-called streamer blowouts), while others appear as wide almost global eruptions. CMEs spanning very large angular ranges are probably not really global, but rather have a large component along the Sun-observer line and so appear large by perspective. These include the so-called halo CMEs. The CDAW CME catalog defines a “partial halo” as a CME with an apparent position angle range $> 120^\circ$. Hence, again, the definition of a CME is restricted by its viewing perspective.

Partial and full halo CMEs occur at a rate of about 10% that of all CMEs, but 360° halo CMEs are only detected at a rate of $\sim 4\%$ of all CMEs.

CMEs that are aligned near the relative disk center tend to be more geoeffective while those nearer the relative solar limb are less so. The vast majority of the most intense geomagnetic storms of Cycle 23, for example, were caused by halo CMEs (Gopalswamy, 2010a).

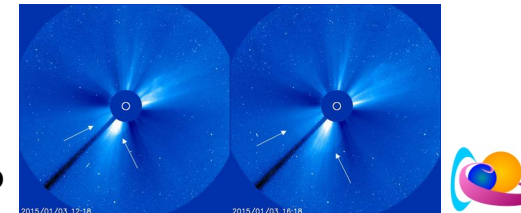
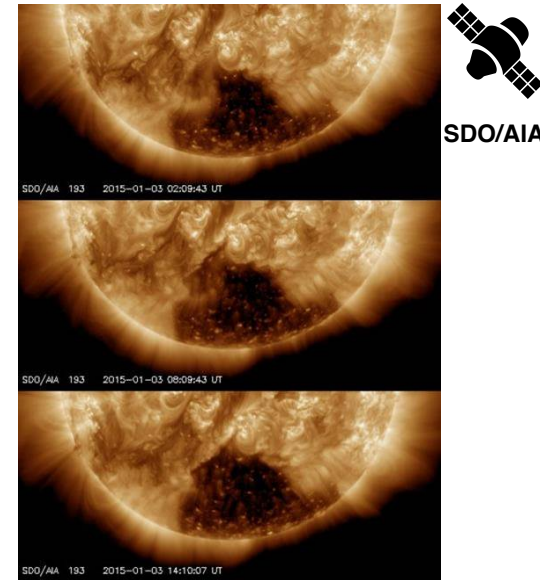
Because of their increased sensitivity, field of view and dynamic range, the SOHO/LASCO and STEREO/COR coronagraphs now frequently observe halo CMEs, which appear as expanding, circular brightenings that completely surround the coronagraphs’ occulting disks (Figure 4). Observations of associated activity on the solar disk are necessary to help distinguish whether a halo CME was launched from the front or backside of the Sun relative to the observer. This has had limited success, as front-sided CMEs that do not have a solar surface association can be mistaken for back-sided events. In recent years several CMEs have been observed by the “three eyes” of STEREO-B, LASCO and STEREO-A by a variety of viewing points, thus reducing this latter problem.

Yashiro et al. (2004) found that slow CMEs tend to accelerate and fast CMEs decelerated through the LASCO field of view, with those around the solar wind speed having constant speeds. Thus, CMEs attain fast acceleration low in the corona until gravity and other drag forces slow them further out. This process continues into the interplanetary medium.

CORONAL MASS EJECTION

Terminology

- Stealth CME
 - No obvious surface signature
 - Rather slow (< 300 km/s)
 - Rather faint
- CME cannibalism
 - 2nd CME overtakes 1st
 - Enhanced geomagnetic storms
- Deflection
 - By corona holes, CME,...



Stealth CME

Source file: Webb et al. (2012): Coronal Mass Ejections: Observations

<http://link.springer.com/article/10.12942/lrsp-2012-3>

The absence of solar surface activity with observed CME activity is not a new observation (Howard and Harrison, 2012). The launch of STEREO in 2006, however, afforded us the opportunity to study the origins of CMEs simultaneously from multiple lines of sight. Robbrecht et al. (2009a) presented a study of a streamer blowout CME without a clear source region. The STEREO spacecraft were sufficiently widely separated (53°) that the CME and its source region could be viewed edge-on in STEREO A and face-on in STEREO B. STEREO B saw the CME as a faint halo and it was detected in-situ as a magnetic cloud 5 days later. Robbrecht et al. suggested that the CME originated high enough up in the corona such that no surface signatures were evident. Subsequently, Ma et al. (2010) performed a statistical study of all CMEs observed during the first 8 months of 2009 when the STEREO lines of sight were nearly perpendicular to each other. They found that about a third of the CMEs were “stealth”, having no distinct surface association, and tending to be slow, i.e., < 300 km s⁻¹. Faint coronal changes could be detected in about half of the stealth CMEs, again suggesting a higher launch site. It is noted that this period was during the recent unusual extended solar minimum, so the fraction of such CMEs may be different at other times.

A good example is in this STCE Newsitem: The curious case of a strong storm

<http://www.stce.be/news/290/welcome.html>

More info at Howard et al. (2012): Stealth Coronal Mass Ejections: A Perspective

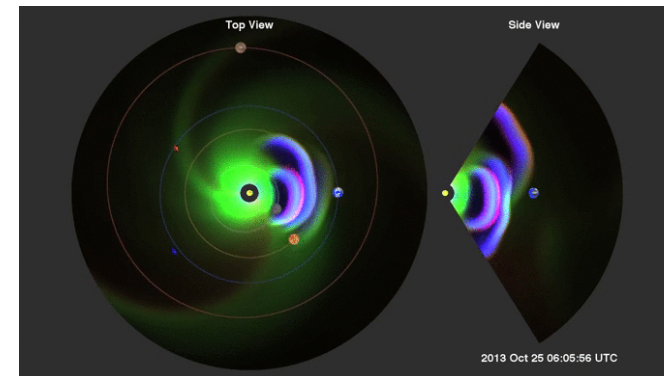
<http://adsabs.harvard.edu/abs/2013SoPh..285..269H>

CORONAL MASS EJECTION



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- Deflection
 - By coronal holes, CME,...



Animation from NASA/GSFC at <https://svs.gsfc.nasa.gov/4099>

In this research model run, the Sun has launched three coronal mass ejections (CMEs) which may merge into a single front as it expands into the solar system. These events are sometimes called 'cannibal' CMEs.

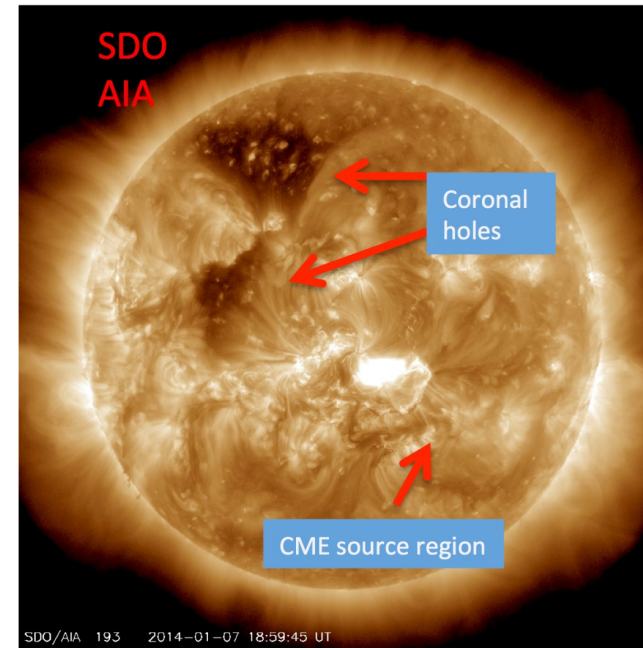
This model run is based on estimated parameters from solar events of October 23–24, 2013

Also at https://science.nasa.gov/science-news/science-at-nasa/2001/ast27mar_1

CORONAL MASS EJECTION

Terminology

- Stealth CME
 - No obvious surface signature
 - Rather slow (<300 km/s)
 - Rather faint
- CME cannibalism
 - 2nd CME overtakes 1st
 - Enhanced geomagnetic storms
- Deflection
 - By coronal holes, CME,...



A good example of deflections is at https://science.gsfc.nasa.gov/674/swl_research.html

Extreme Solar Wind Deflects CMEs

A very fast CME was observed on January 7, 2014. Preliminary data analysis and all 8 community forecasts reported in GSFC's Space Weather Scoreboard indicated rapid arrival at Earth and a major geo-magnetic storm. However, the CME arrived at Earth ~19 hr after the predicted time, and the geomagnetic storm was weak ($K_p < 3$). What happened? Detailed analysis by the CCMC/SWRC team identified possible causes for the gap between predicted vs. actual outcome. The solar wind coming from the nearby coronal holes was extremely fast – 950 km/s at Earth (very rare!) and deflected the CME away from the Earth. However, the solar wind speed assumed at the lower boundary of the CME transport model (WSA-ENLIL) was too low – 750 km/s (maximum allowed value). Therefore the model CME propagated to Earth much too slowly. Previously the same coronal hole did not produce such high speed wind, so the strong deflection was a surprise. We know that CMEs can be deflected by a coronal hole, so a CME that seems to be Earth-directed can be deflected from the Earth-Sun line. The simulations did not predict that the deflection would be so large that the CME only hit a glancing blow to the Earth.

Deflections

Kay et al. (2015): Global Trends of CME Deflections Based on CME and Solar Parameters

<http://adsabs.harvard.edu/abs/2015ApJ...805..168K>

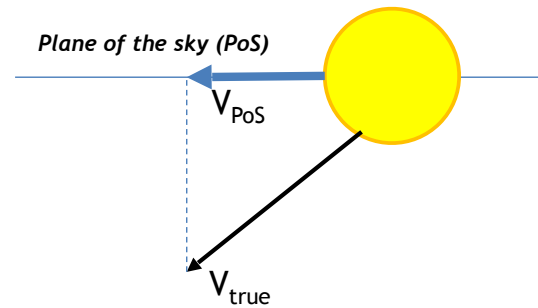
Forecasting space weather effects relies on knowledge of the path of a CME. Observations commonly show significant non-radial deviations in the CME trajectories. Understanding these deflections will allow for more accurate space weather predictions. Coronal observations show that CMEs can undergo significant deflection close to the Sun, but it is often hard to disentangle the effects of deflection, rotation, and non-uniform expansion in the lower corona (Nieves-Chinchilla et al. 2012). Byrne et al. (2010) measure a latitudinal deflection of 30° below $7 \odot R$ for the 2008 December 12 CME. Kilpua et al. (2009) suggest that CMEs may not be able to penetrate the open magnetic field emanating from coronal holes (CHs). The CH magnetic field then guides CMEs toward the Heliospheric Current Sheet (HCS). Shen et al. (2011) and Gui et al. (2011) attribute the deflection to gradients in the background magnetic energy density, which would also cause CMEs to tend to deflect toward the HCS. As with the observed CMEs, the MHD CMEs tend toward regions of lower magnetic energy. In some cases, magnetic reconnection creates an imbalance in the magnetic energy, which causes a CME to deflect early in the eruption (Zuccarello et al. 2012; Lynch & Edmondson 2013). MHD simulations also show that CMEs can deflect due to interactions with other CMEs (Lugaz et al. 2012). Finally, there are also effects of CME rotation due to a torque created by differential forces along the CME's toroidal axis.

CORONAL MASS EJECTION

Speed

- We see the projected speed
 - Plane of the sky (PoS)
- We use the true speed
 - = Corrected PoS speed

$$V_{\text{true}} \geq V_{\text{PoS}}$$
 - From Type II radio bursts
 = shock speed



Created: 2012 Jul 15 0332 UT
 Date: 2012 07 12
 Prepared by the U.S. Dept. of Commerce, NOAA, Space Weather Prediction Center
 Please send comments and suggestions to SWPC.Webmaster@noaa.gov
 Missing data: ////
 Updated every 30 minutes.

Edited Events for 2012 Jul 12

#Event	Begin	Max	End	Obs	Q	Type	Loc/Frq	Particulars	Reg#
9900 +	1537	1649	1730	G15	5	XRA	1-8A	X1.4 4.6E-01	1520
9900	1610	////	1638	PAL	C	RSP	108-180	CIW/1	1520
9900 +	1614	1649	1732	SAG	G	RBR	410	6600	1520
9900 +	1614	1652	1706	SAG	G	RBR	1415	1100	1520
9900 +	1614	1653	1705	SAG	G	RBR	2695	800	1520
9900 +	1614	1653	1714	SAG	G	RBR	4995	480	1520
9900 +	1615	1653	1824	SAG	G	RBR	8800	430	1520
9900	1615	1652	2010	SAG	G	RBR	245	3900	1520
9900	1620	1655	1814	SAG	G	RBR	610	2400	1520
9900 +	1621	1654	1815	SAG	G	RBR	15400	220	1520
9900 +	1625	////	1653	SAG	C	RSP	025-082	II/2 1268	1520
9900 +	1638	////	2359	PAL	C	RSP	025-180	IV/2	1520



Source file: Webb et al. (2012): Coronal Mass Ejections: Observations
<http://link.springer.com/article/10.12942/lrsp-2012-3>

Some additional information on the relation between CMEs, CME shocks and Type II radio bursts can be found at http://www.ovsa.njit.edu/fasr/Chapter_15.pdf (Gopalswamy: Interplanetary Radio bursts, in Solar and Space Weather Radiophysics – Chapter 15).
 As well as in Gopalswamy et al. (2008): Coronal mass ejections, type II radio bursts, and solar energetic particle events in the SOHO era
<http://adsabs.harvard.edu/abs/2008AnGeo..26.3033G>
<http://www.ann-geophys.net/26/3033/2008/angeo-26-3033-2008.pdf>

Some good examples on how to calculate/deduce the CME speed:

- NASA: <http://rodshome.com/TLA/sunspots/CMEvelocity%20calc.pdf>
- Pohjolainen et al. (2007): CME Propagation Characteristics from Radio Observations
<http://adsabs.harvard.edu/abs/2007SoPh..244..167P>

Data of Type II bursts with derived shock speeds:

NOAA: <https://www.swpc.noaa.gov/products/solar-and-geophysical-event-reports>
 SWS: <http://www.sws.bom.gov.au/Solar/2/3>

Attention: The shock speed is usually (a bit / a lot) higher than the (corrected) CME speed.

INTERPLANETARY CORONAL MASS EJECTION

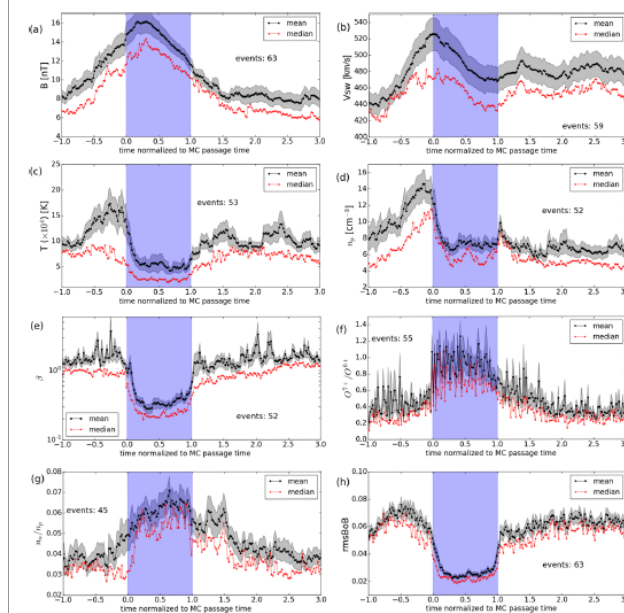
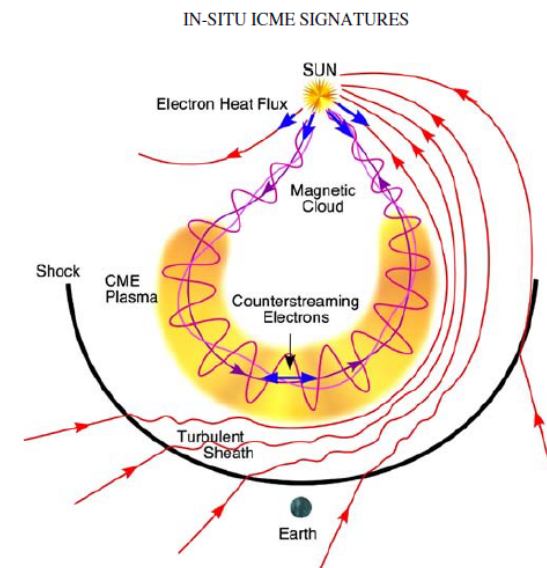


Figure 2. Schematic of the three-dimensional structure of an ICME and upstream shock, relating magnetic field, plasma, and BDE signatures.



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Figure right: Zurbuchen et al. (2006): In-Situ Solar Wind and Magnetic Field Signatures of Interplanetary Coronal Mass Ejections <http://adsabs.harvard.edu/abs/2006SSRv..123...31Z>

An interplanetary CME (ICME) is a CME of which the solar wind features are measured in situ by spacecraft at Earth or in the solar system.

Pending the mutual positions of the Earth and the CME, Earth may experience the following impacts from this (I)CMEs:

1. Nothing
2. Shock + Sheath
3. Shock + Sheath + Magnetic Cloud leg (long)
4. Shock + Sheath + Magnetic Cloud (head-on) + rarefied region
5. No shock, still magnetic cloud

These all give different signatures in the various solar wind parameters.

Rodriguez et al. (2016): Typical Profiles and Distributions of Plasma and Magnetic Field Parameters in Magnetic Clouds at 1 AU

<http://adsabs.harvard.edu/abs/2016SoPh..291.2145R>

Figure to the left: Rodriguez et al. (2016): Typical Profiles and Distributions of Plasma and Magnetic Field Parameters in Magnetic Clouds at 1 AU

<http://adsabs.harvard.edu/abs/2016SoPh..291.2145R>

Coronal mass ejections (CMEs) are large-scale solar eruptive events in which large amounts of plasma carrying magnetic flux and helicity (see e.g. Démoulin, Janvier, and Dasso, 2016, and references therein) are expelled into the interplanetary space. When sampled in situ by a spacecraft in the interplanetary medium, they are called interplanetary CMEs (ICMEs). Magnetic clouds (MCs) are an important subset of ICMEs that exhibit a particular internal magnetic field configuration resembling that of a flux rope. This is characterized by an enhanced magnetic field intensity, smooth rotation of its magnetic field vector, and low temperature (e.g. Burlaga, 1991).

The classical three-part structure of a CME (bright front, dark cavity, and dense core) is commonly also interpreted in terms of a magnetic flux rope propagating in the corona (see e.g. Illing and Hundhausen, 1986; Vourlidas et al., 2013). The bright front corresponds to the plasma pile-up in front of the flux rope, the cavity represents the bulk of the flux rope, and the dense core is the erupting prominence that is located in the bottom (concave-out) parts of the flux rope field lines. However, it is very difficult to identify the corresponding three-part morphology in ICMEs detected in situ (e.g. Kilpua et al., 2013a).

The plasma β (the ratio of the plasma pressure to the magnetic pressure) and the level of fluctuations in the magnetic field vector are the best parameters to define the boundaries of MCs. We find that one third of the events shows a peak in plasma density close to the trailing edge of the flux ropes.





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<https://events.oma.be/indico/event/21/>

The solar wind-magnetosphere coupling is governed by the duskward electric field $E_y \sim vB_z$. However, v varies only by a factor of 2 (maybe 5 in extreme events). B_z varies by a factor of 10 and is thus a parameter more important for predictions. To be geoeffective, the CME-associated disturbance should have a suitable magnetic field configuration: the interplanetary magnetic field (IMF) B_z component should be negative (southward), strong enough and long-lasting.

:Issued: 2023 Mar 18 1231 UTC
:Product: documentation at <http://www.sidc.be/products/meu>
#-----#
DAILY BULLETIN ON SOLAR AND GEOMAGNETIC ACTIVITY from the SIDC #
(RWC Belgium) #
#-----#

SIDC URSIGRAM 30318
SIDC SOLAR BULLETIN 18 Mar 2023, 1230UT
SIDC FORECAST (valid from 1230UT, 18 Mar 2023 until 20 Mar 2023)
SOLAR FLARES : C-class flares expected, (probability >=50%)
GEOMAGNETISM : Quiet (A<20 and K<4)
SOLAR PROTONS : Quiet
PREDICTIONS FOR 18 Mar 2023 10CM FLUX: 138 / AP: 007
PREDICTIONS FOR 19 Mar 2023 10CM FLUX: 140 / AP: 019
PREDICTIONS FOR 20 Mar 2023 10CM FLUX: 138 / AP: 029

COMMENT: The solar flaring activity was at moderate levels during the last 24 hours with one M-class flare and several C-class flares detected, with the most frequent sources being NOAA active regions 3254 and 3256. The largest flare was a M1.1 flare, peaking at 15:07 UTC on March 17, associated with active region NOAA 3254 (beta class). NOAA active region 3256 produced an impulsive C9.4 flare at 07:10 UTC on March 18. This event was also associated with Type IV radio emission. Other regions on the disc did not show any significant flaring activity. Further M-class flare activity is possible but not probable, while frequently C-class activity is expected in the next 24 hours.

A filament eruption in the southwestern quadrant was observed on March 17 from around 09:20UTC. The associated CME appears in SoHO/LASCO C2 coronagraph data from 10:23UTC onwards. The CME is directed to the south-west and the bulk of the CME is not expected to be Earth directed. However, a glancing blow of the shock may impact Earth at around 19:00 UTC on March 19. Another small filament eruptions occurred in the northwestern quadrant from around 17:09UTC and 20:09UTC on March 17. We are awaiting corresponding coronagraph data for further analysis. During the last 24 hours there were no other potentially Earth-directed CMEs detected in the available coronagraph observations.

The greater than 10 MeV proton flux was at almost nominal levels over the past 24 hours and is expected to remain so for the next 24 hours. The greater than 2 MeV electron flux remained below the 1000 pfu alert threshold and is expected to remain below this threshold during the next 24 hours. The 24h electron fluence was at normal levels and is expected to remain so.

Over the past 24 hours the solar wind parameters (ACE and DSCOVR) have been indicative of slow solar wind conditions. The solar wind speed ranged between 400 km/s and 450 km/s. The interplanetary magnetic field magnitude was about 6 nT. The magnetic field orientation was predominantly in the positive sector (field directed away from the Sun). Similar slow solar wind regime is expected on March 18 with a slight wind speed enhancement possible for late on March 19, due to expected influence of the small equatorial coronal hole of positive polarity with a chance of being mixed with glancing blow from a CME which left the solar surface around 10 UTC on March 17th.

The geomagnetic conditions over the past 24 hours were globally and locally quiet to unsettled (NOAA Kp and K Bel 1-3). Quiet conditions are expected for March 18 with active to minor storm conditions possible for late on March 19 and March 20, due to expected arrival of the high speed stream and a possible glancing blow from a CME.

TODAY'S ESTIMATED ISN : 044, BASED ON 14 STATIONS.

SOLAR INDICES FOR 17 Mar 2023
WOLF NUMBER CATANIA : 110
10CM SOLAR FLUX : 134
AK CHAMBON LA FORET : 014
AK WINGST : 007
ESTIMATED AP : 007
ESTIMATED ISN : 073, BASED ON 18 STATIONS.

NOTICEABLE EVENTS SUMMARY
DAY BEGIN MAX END LOC XRAY OP 10CM Catania/NOAA RADIO_BURST_TYPES
17 1504 1507 1511 S22W65 M1.0 SN 12/3247
END



CME/ICME

